Last lecture (6)

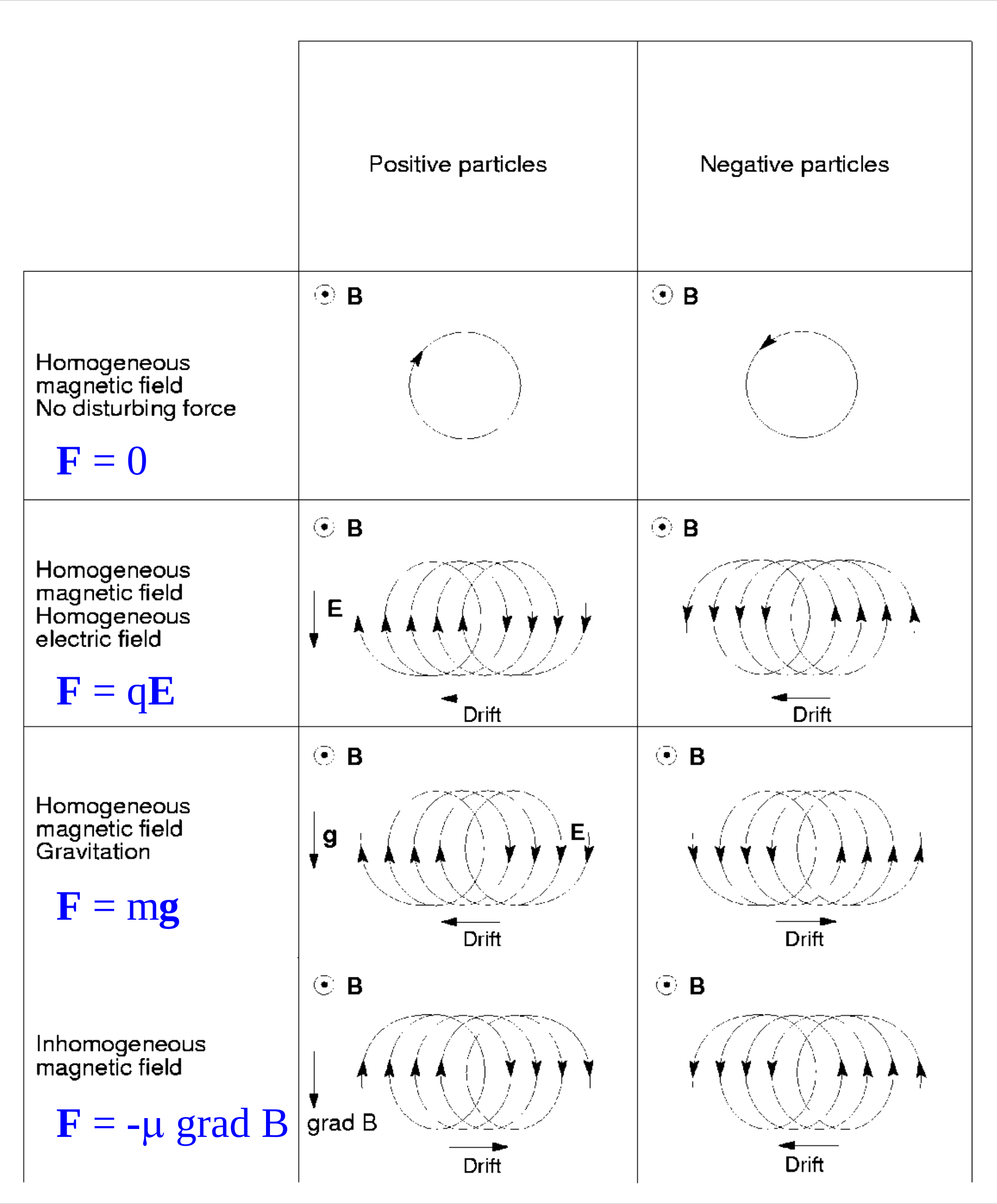
- Drift motion
- Ionospheric conductivities
- Particle motion in the magnetosphere

Today's lecture (7)

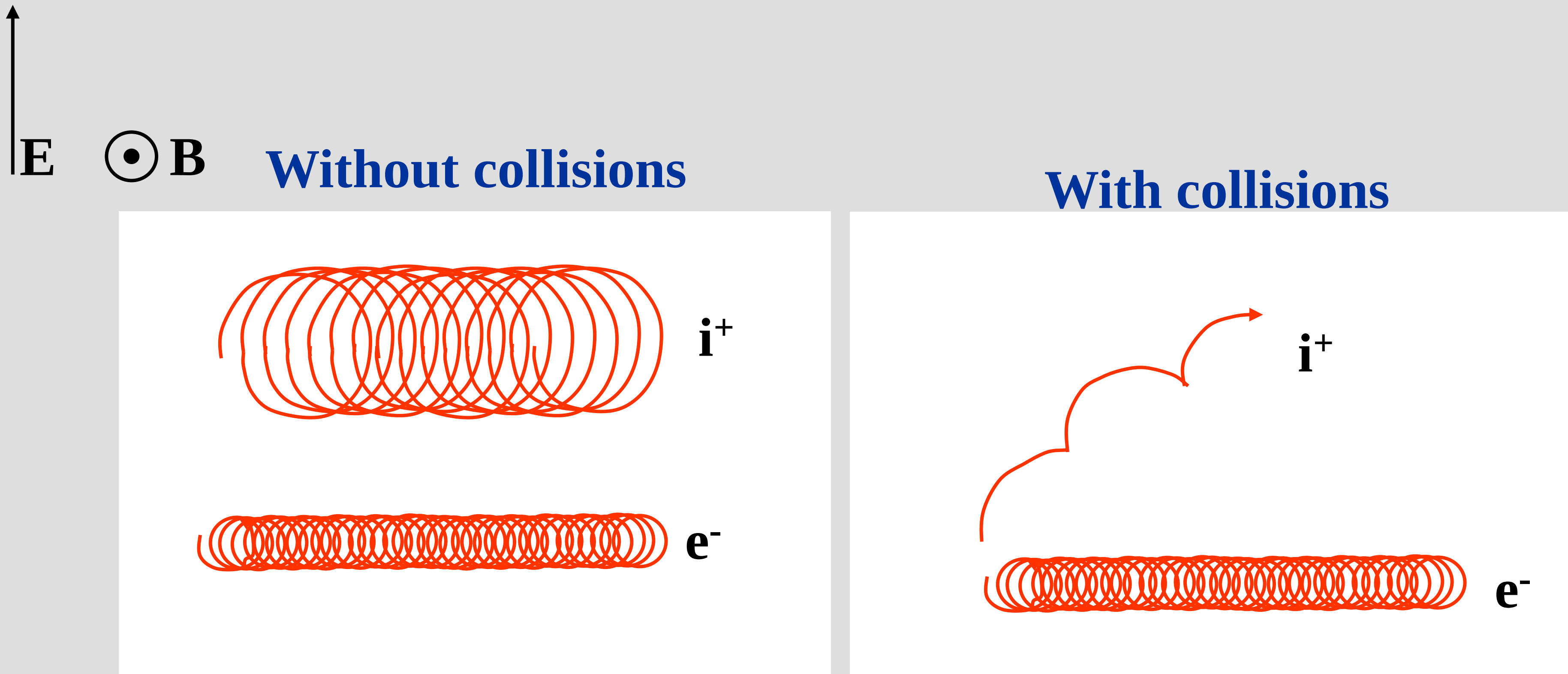
- Particle motion in the magnetosphere
- Other magnetospheres
- Aurora

Drift motion

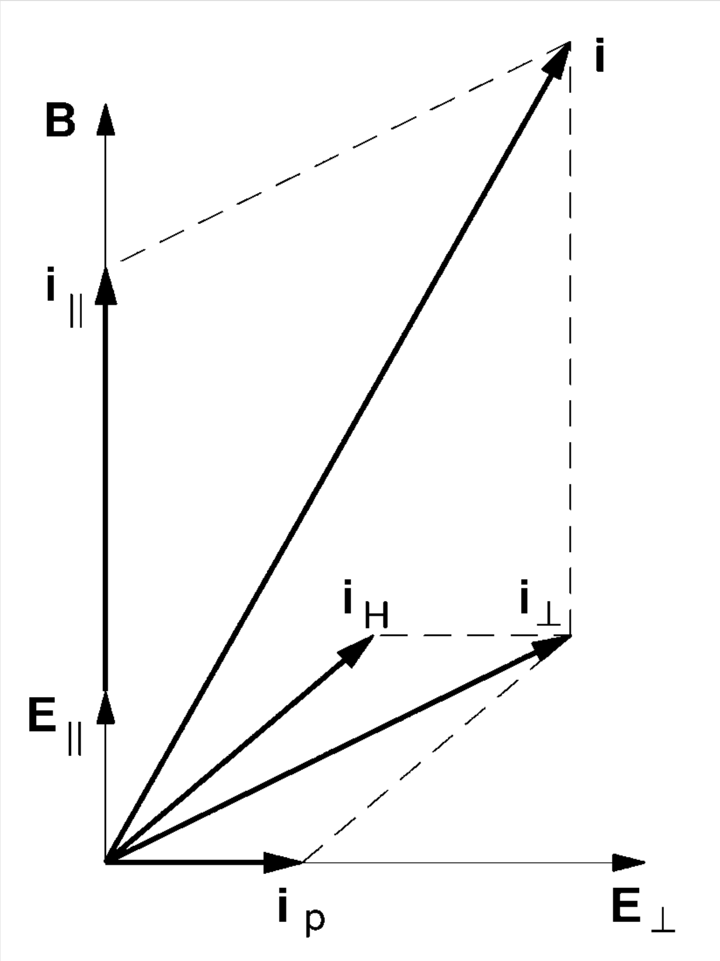
$$\mathbf{u}_{drift} = \frac{\mathbf{F} \times \mathbf{B}}{qB^2}$$



ExB-drift

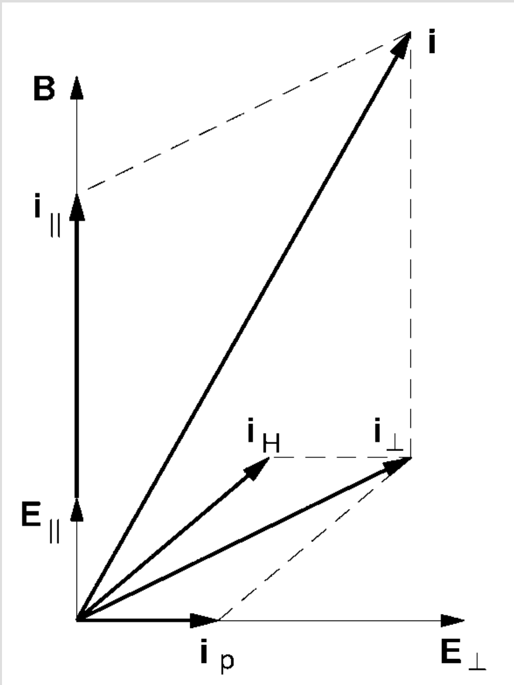


Electric conductivity in a magnetized plasma



- $i_{//}$ = parallel current
- i_P = Pedersen current
- i_H = Hall current

Electric conductivity in a magnetized plasma II



$$\sigma_P = \sigma_e \frac{1}{1+\omega_{ge}^2\tau_e^2} + \sigma_i \frac{1}{1+\omega_{gi}^2\tau_i^2}$$

$$\sigma_H = \sigma_e \frac{\omega_{ge}\tau_e}{1+\omega_{ge}^2\tau_e^2} - \sigma_i \frac{\omega_{gi}\tau_i}{1+\omega_{gi}^2\tau_i^2}$$

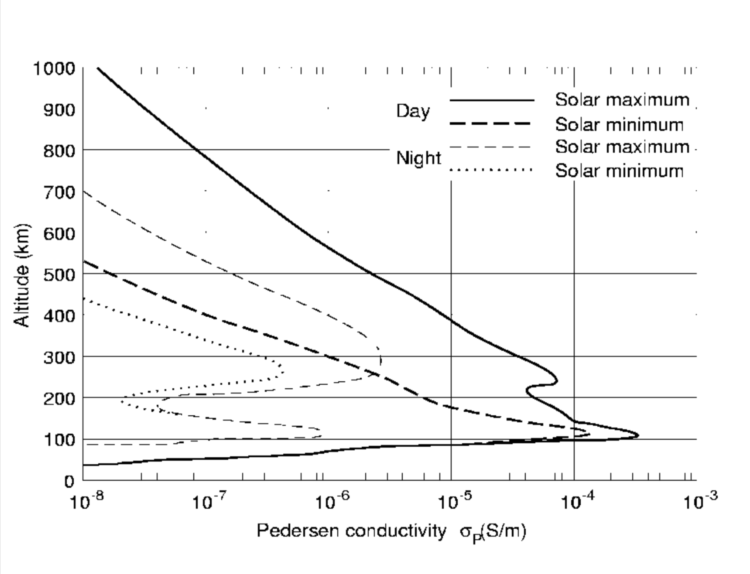
$$\sigma_{//} = \sigma_e + \sigma_i$$

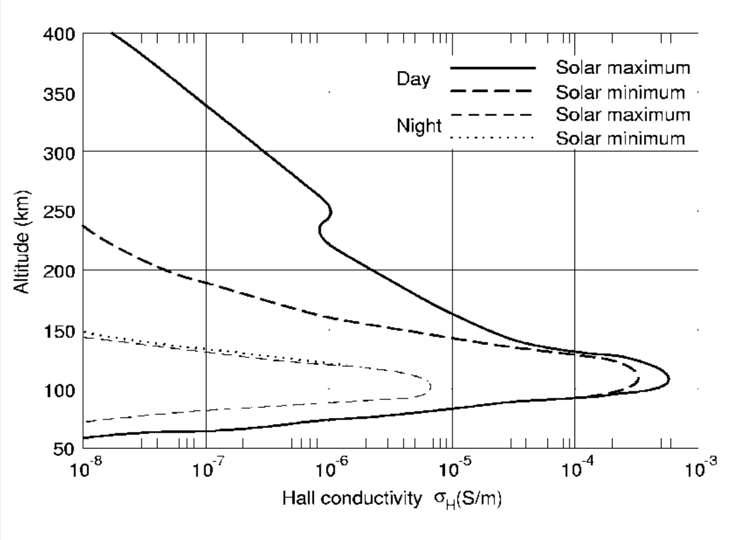
$$\sigma_e = e^2 n \tau_e / m_e \qquad \sigma_i = e^2 n \tau_i / m_i$$

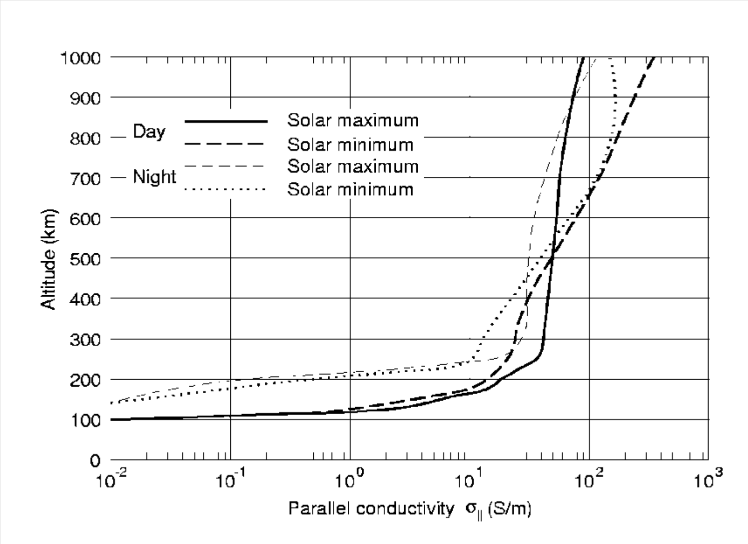
$$i_{//} = \sigma_{//} E_{//}$$

$$\left. \begin{array}{l} i_P = \sigma_P E_\perp \\ i_H = \sigma_H E_\perp \end{array} \right\} \quad \text{or} \quad \mathbf{i}_\perp = \sigma_P \mathbf{E}_\perp + \sigma_H \frac{\mathbf{B}\times\mathbf{E}_\perp}{B}$$

Ionospheric conductivities







Geomagnetic field

Approximated by a dipole close to Earth.

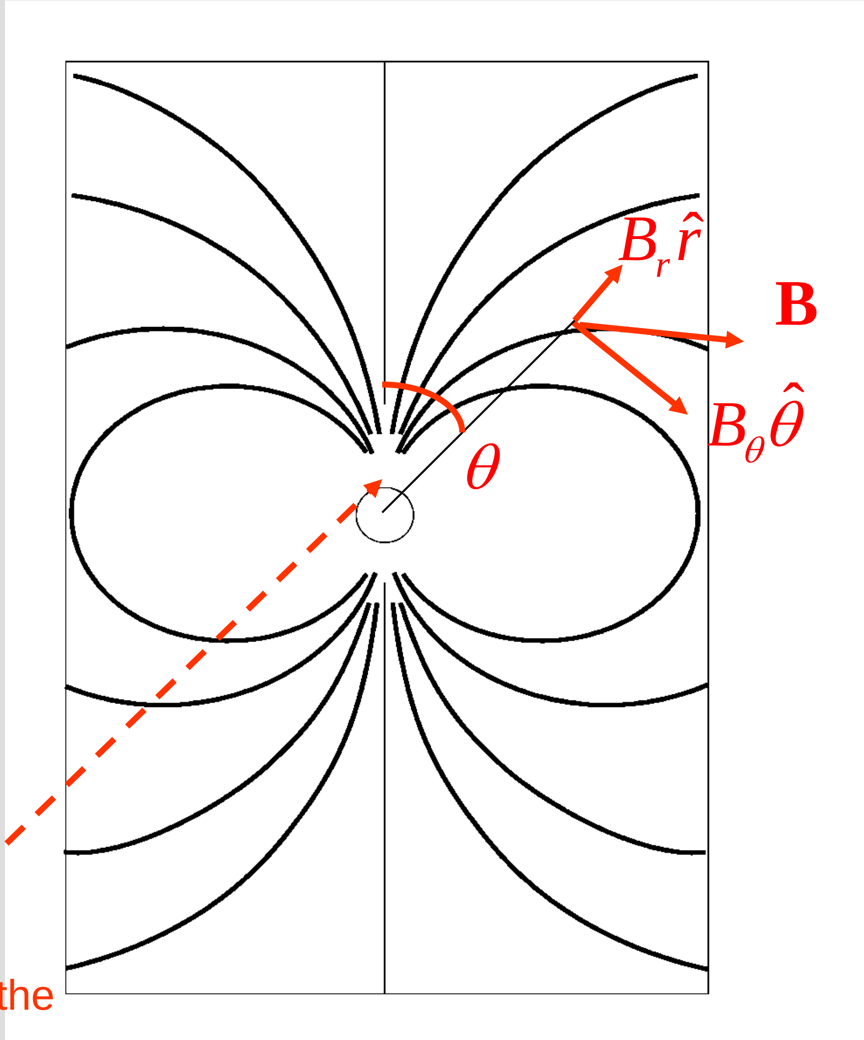
$$B_r = B_p \left(\frac{R_E}{r}\right)^3 \cos\theta$$

$$B_\theta = \frac{B_p}{2} \left(\frac{R_E}{r}\right)^3 \sin\theta$$

$$a = \frac{2\pi R_E^3 B_p}{\mu_0}$$

magnetic dipole moment

Magnetic field at the "north pole"



Geomagnetic field

Alternative formulation of dipole field

$$B_r = B_p\left(\frac{R_E}{r}\right)^3 \cos\theta$$

$$B_\theta = \frac{B_p}{2}\left(\frac{R_E}{r}\right)^3 \sin\theta$$

$$B_r = \frac{\mu_0 a}{2\pi}\frac{1}{r^3}\cos\theta$$

$$B_\theta = \frac{\mu_0 a}{2\pi}\cdot\frac{1}{2}\cdot\frac{1}{r^3}\sin\theta$$

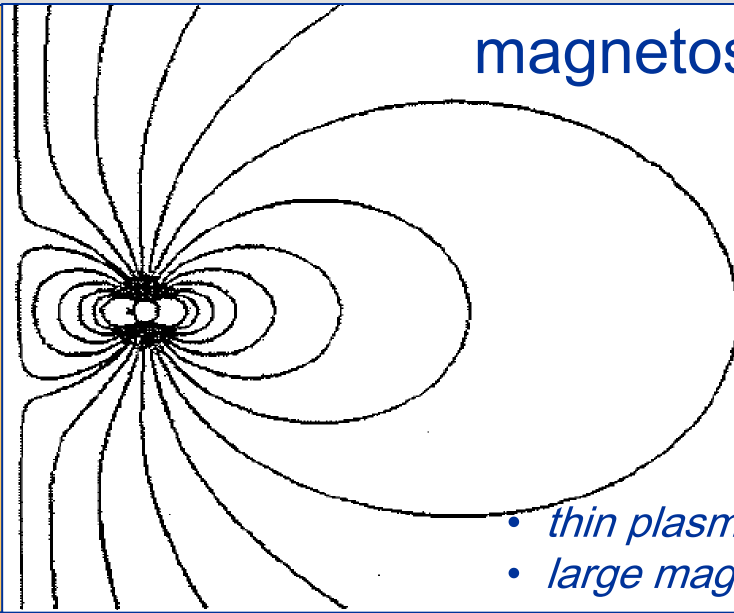
$$a = \frac{2\pi R_E^3 B_p}{\mu_0}$$

magnetic dipole moment

Stand-off distance from pressure balance

solar wind

- dense plasma with ***high conductivity***
- *weak magnetic field*



magnetosphere

- *thin plasma*
- *large magnetic field*

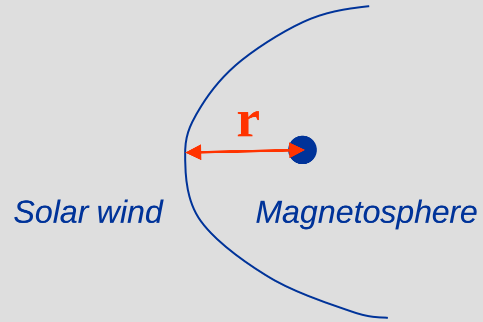
Dynamic pressure:

$$p_d = \rho_{SW} v_{SW}^2$$

Magnetic pressure:

$$p_B = \frac{B^2}{2\mu_0}$$

Magnetopause ”stand-off distance”



Dynamic pressure: $p_d = \rho_{SW} v_{SW}^2$

Magnetic pressure: $p_B = \dfrac{1}{2\mu_0} B^2$

Dipole field strength (in equatorial plane): $B = \dfrac{\mu_0 a}{4\pi} \dfrac{1}{r^3}$

$p_d = p_B \Rightarrow$

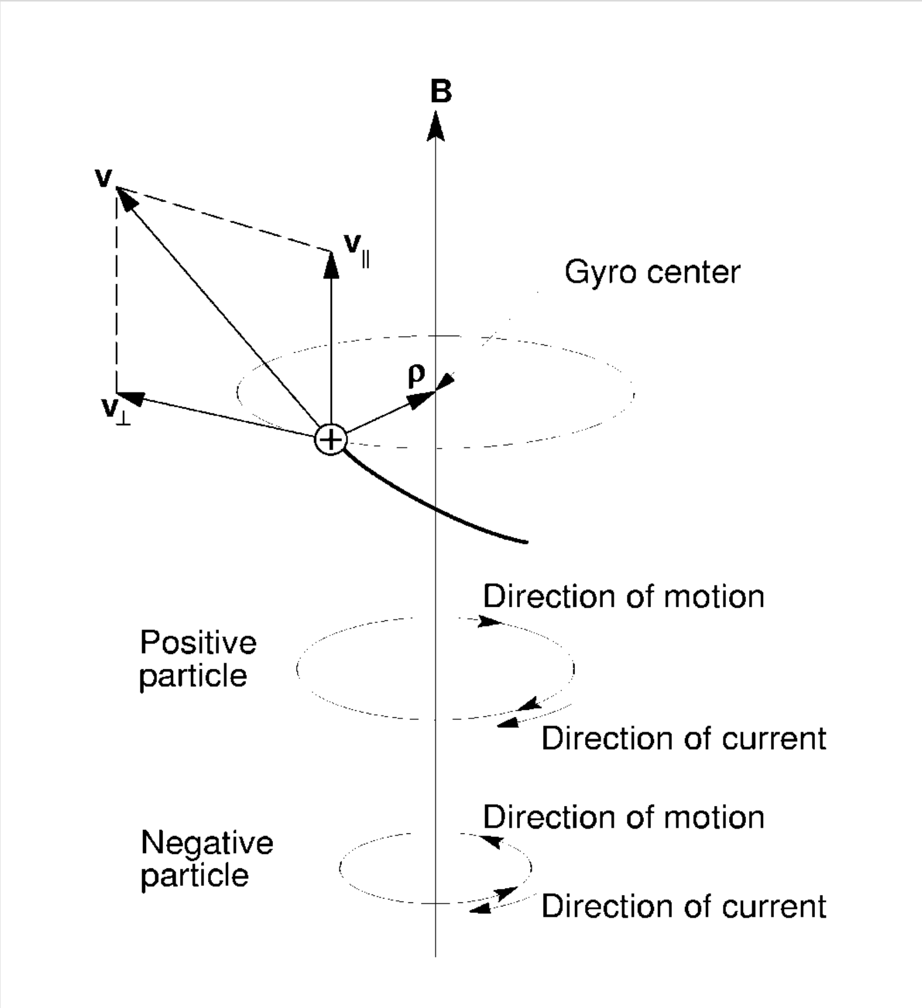
$$\rho_{SW} v_{SW}^2 = \left[\frac{\mu_0 a}{4\pi} \frac{1}{r^3} \right]^2 / 2\mu_0 \Rightarrow$$

$$r = \left(\frac{\mu_0 a}{4\pi} \right)^{1/3} \left(2\mu_0 \rho_{SW} v_{SW}^2 \right)^{-1/6}$$

$a = 8 \times 10^{22}$ Am^2, v=500 km/s, $\rho_{SW} = 10^7 \times 1.7 \times 10^{-27}$ kg/m^3:

r = 7 R$_e$ (1 R$_e$ = 6378 km)

Particle motion in magnetic field



gyro radius

$$\rho = \frac{mv_{\perp}}{qB}$$

gyro frequency

$$\omega_g = \frac{qB}{m}$$

magnetic moment

$$\mu = IA = q\, f_g \pi\, \rho^2 = mv_{\perp}{}^2/2B$$

Adiabatic invariant

DEFINITION:

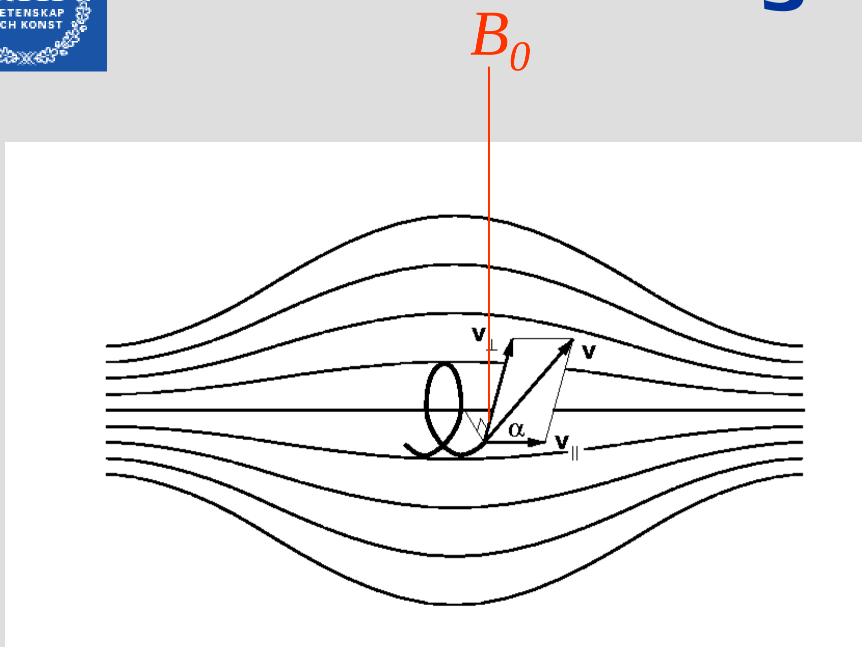
An **adiabatic invariant** is a property of a physical system which stays constant when changes are made slowly.

By 'slowly' in the context of charged particle motion in magnetic fields, we mean much slower than the gyroperiod.

'First adiabatic invariant' of particle drift:

$$\mu = \frac{mv_\perp^2}{2B}$$

Magnetic mirror



$mv^2/2$ constant (energy conservation) ➡

$$\frac{\sin^2\alpha}{B} = konst$$

particle turns when α = 90° ➡

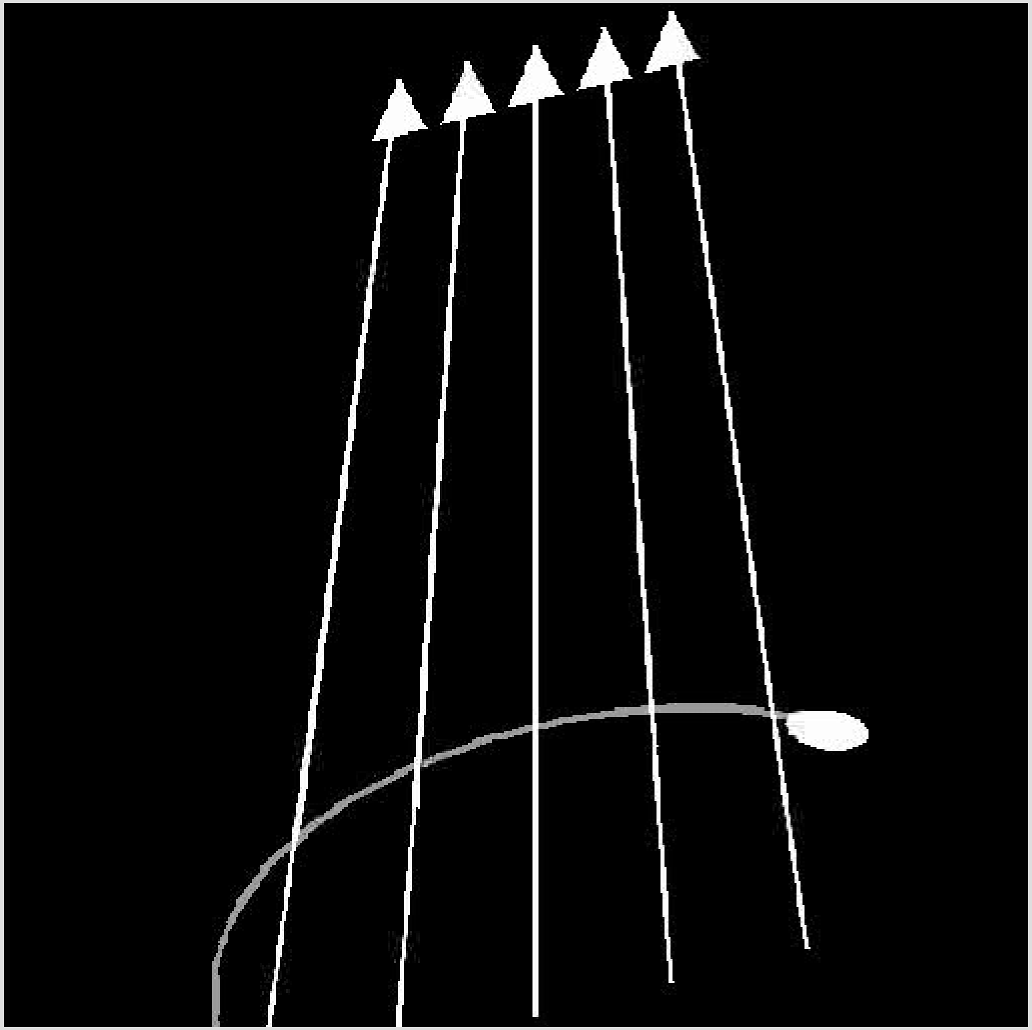
$$\frac{B_{turn}}{\sin^2 90^\circ} = \frac{B_0}{\sin^2\alpha}$$ ➡

$$B_{turn} = \frac{B_0}{\sin^2\alpha}$$

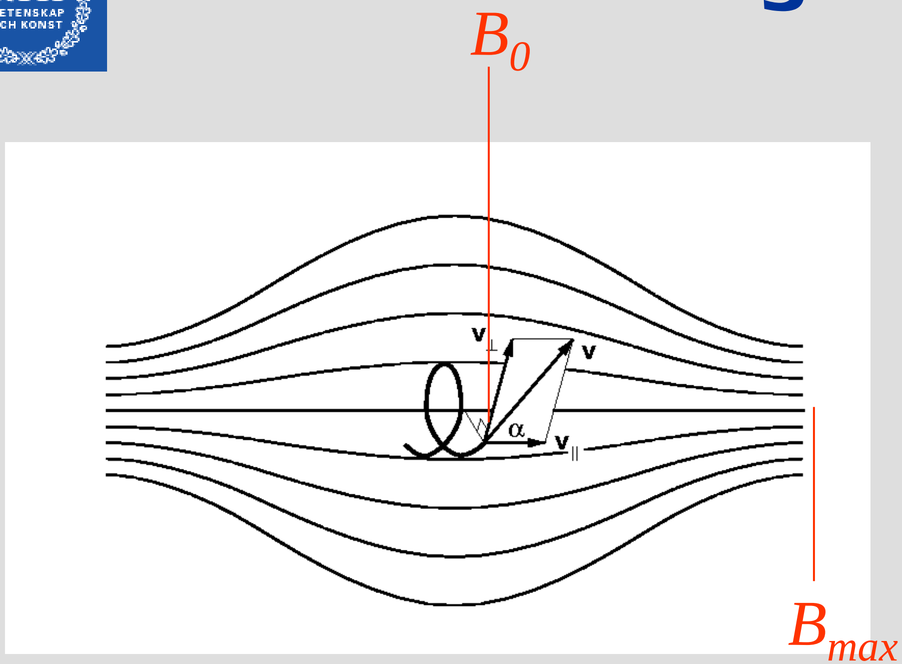
The magnetic moment μ is an ***adiabatic invariant***.

$$\mu = \frac{mv_\perp^2}{2B} = \frac{mv^2\sin^2\alpha}{2B}$$

Magnetic mirror



Magnetic mirror



mv²/2 constant (energy conservation) →

$$\frac{\sin^2\alpha}{B} = konst$$

particle turns when α = 90° →

$$B_{turn} = B_0 / \sin^2\alpha$$

If maximal *B*-field is B_{max} a particle with pitch angle α can only be turned around if

$$B_{turn} = B_0 / \sin^2\alpha \le B_{\max}$$ →

$$\alpha > \alpha_{lc} = \arcsin\sqrt{B_0 / B_{\max}}$$

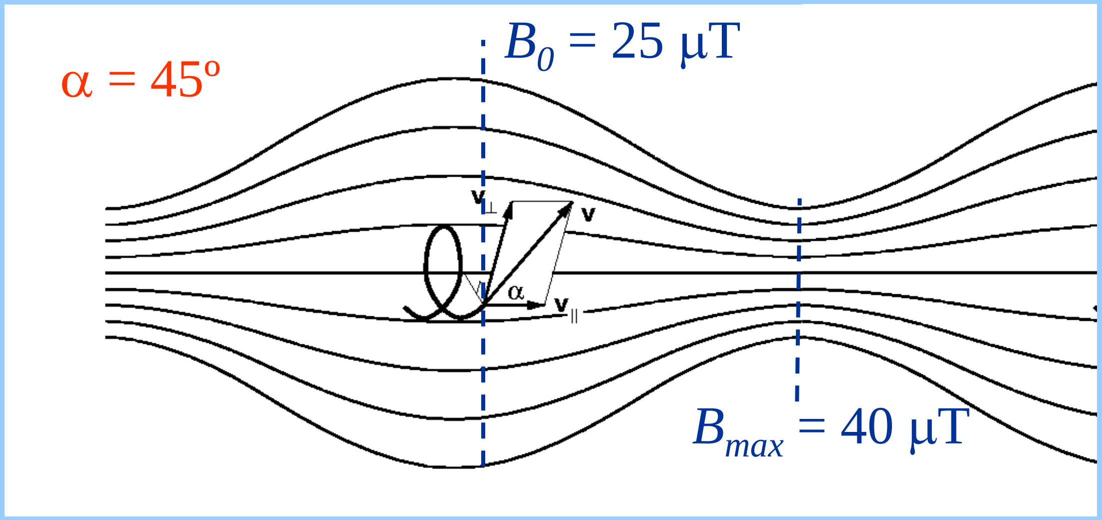
The magnetic moment μ is an ***adiabatic invariant***.

$$\mu = \frac{mv_\perp^2}{2B} = \frac{mv^2\sin^2\alpha}{2B}$$

Particles in ***loss cone*** :

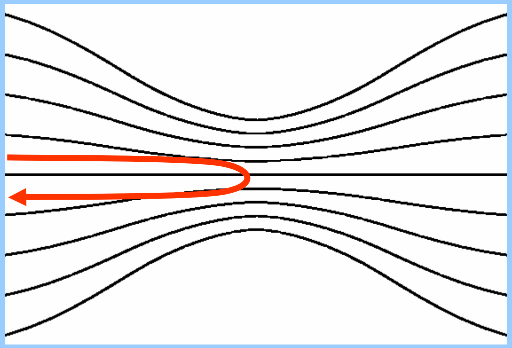
$$\alpha < \alpha_{lc}$$

What will happen to the particle?

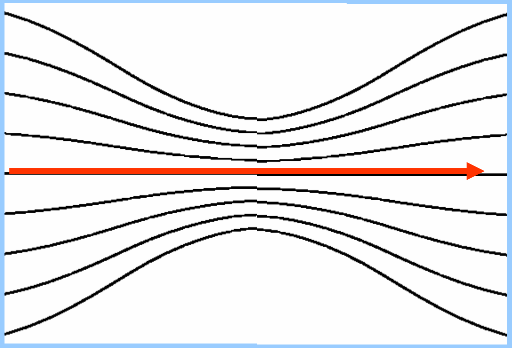


$$\alpha_{lc} = \arcsin\sqrt{B_0 / B_{max}}$$

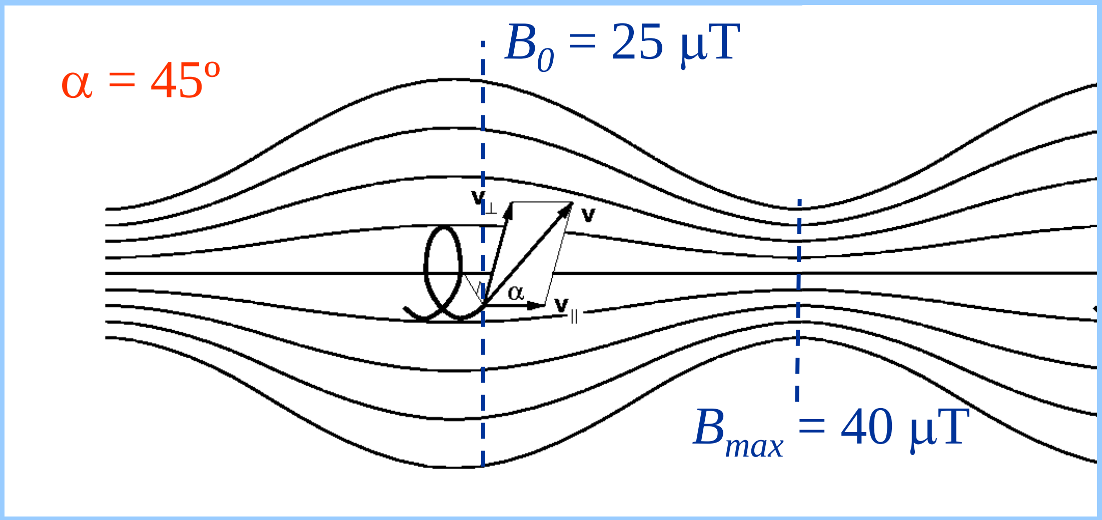
Blue — It will mirror



Yellow — It will escape

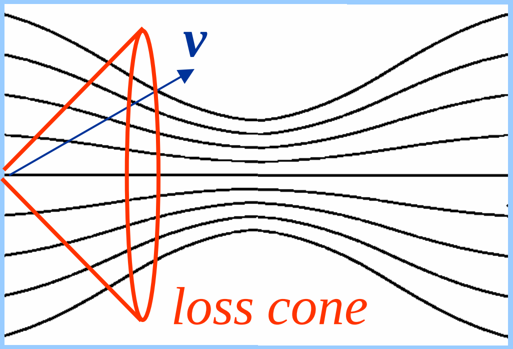


What will happen to the particle?



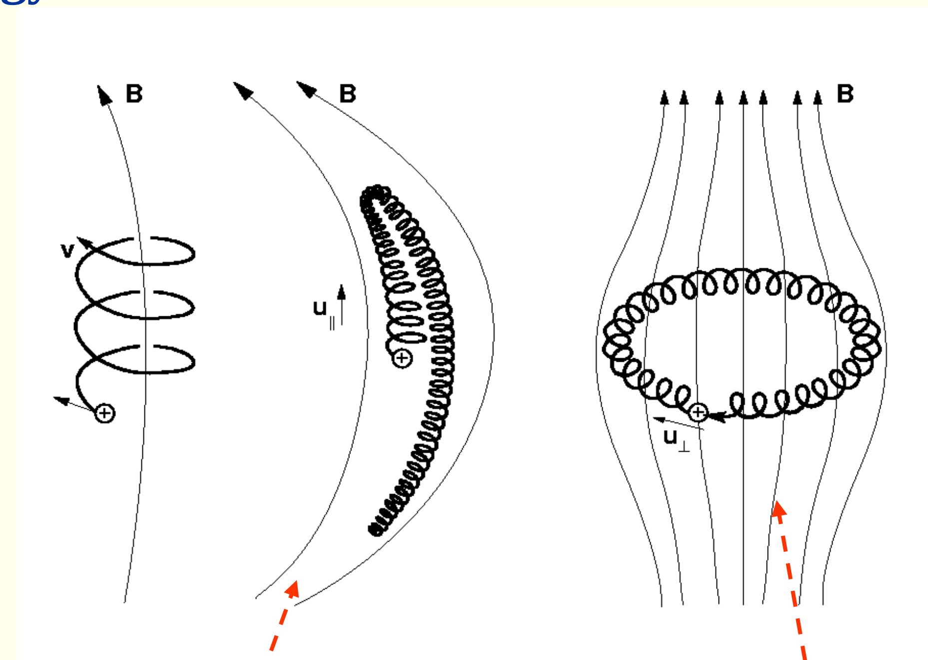
$$\alpha_{lc} = \arcsin\sqrt{B_0 / B_{\max}} = \arcsin\sqrt{25/40} = 52°$$

Yellow It will escape



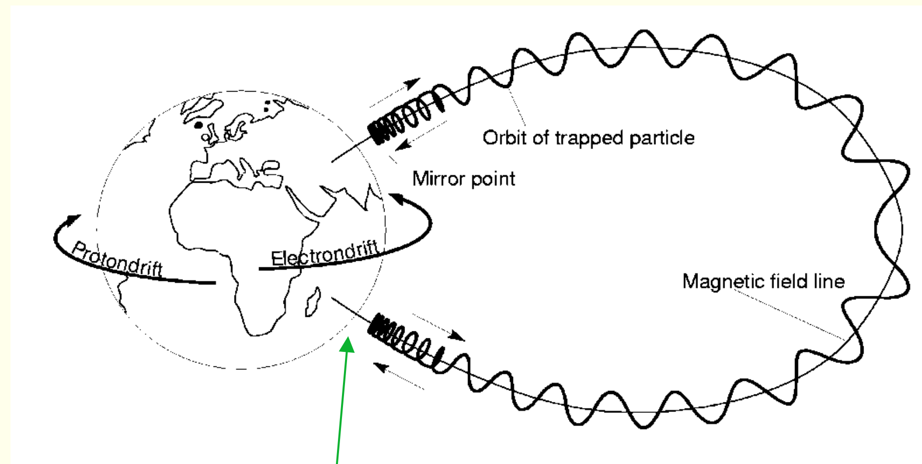
Particle motion in geomagnetic field

gyration **longitudinal oscillation** **azimuthal drift**



Magnetic mirror

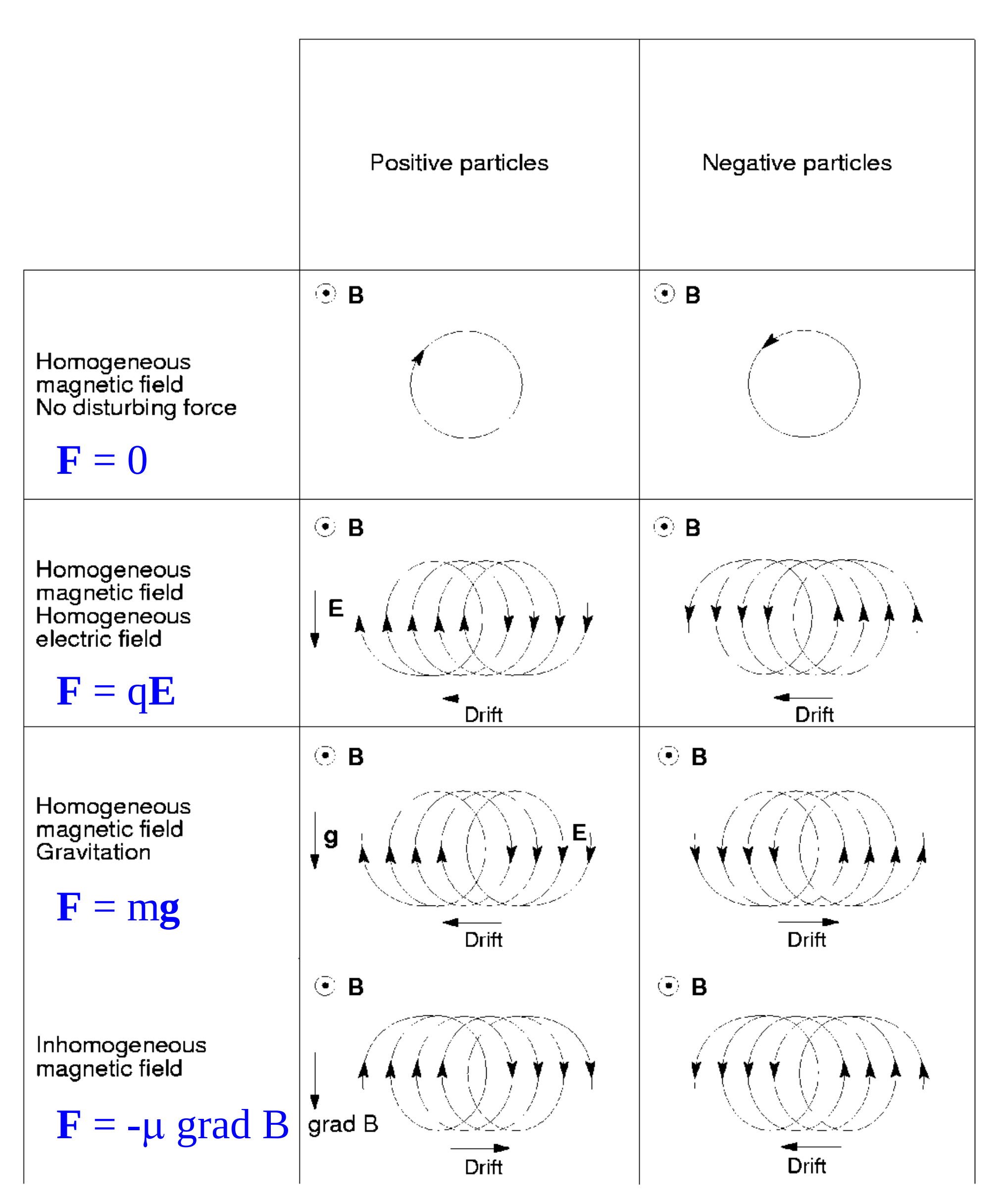
grad B drift



Particles in the loss cone create the aurora!

Drift motion

$$\mathbf{u} = \frac{\mathbf{F} \times \mathbf{B}}{qB^2}$$

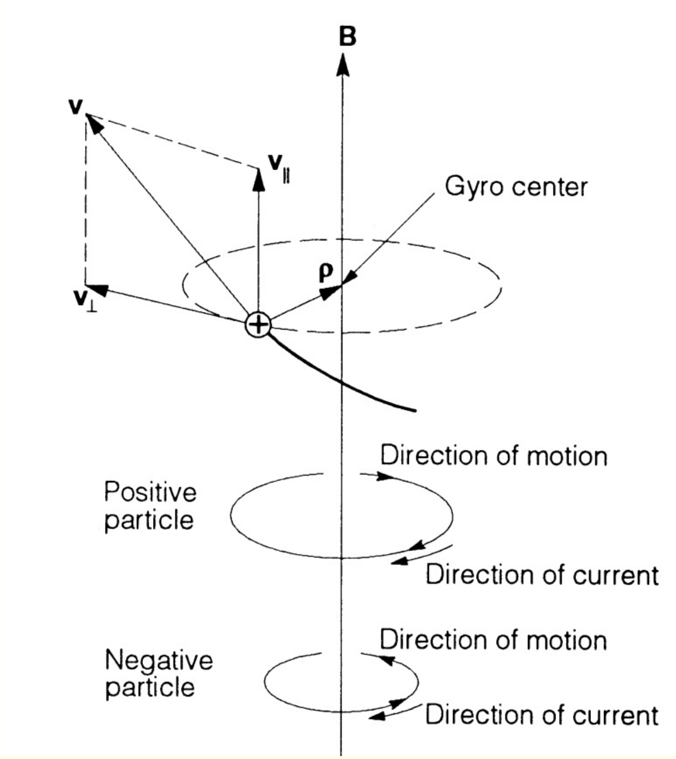


Force on magnetic dipole

$$\boldsymbol{\mu} \sim -\mathbf{B} \quad \Rightarrow \quad \boldsymbol{\mu} = -\mu \frac{\mathbf{B}}{B}$$

$$\mathbf{F} = \nabla(\boldsymbol{\mu} \cdot \mathbf{B}) = -\mu \nabla \left(\frac{\mathbf{B}}{B} \cdot \mathbf{B} \right) =$$

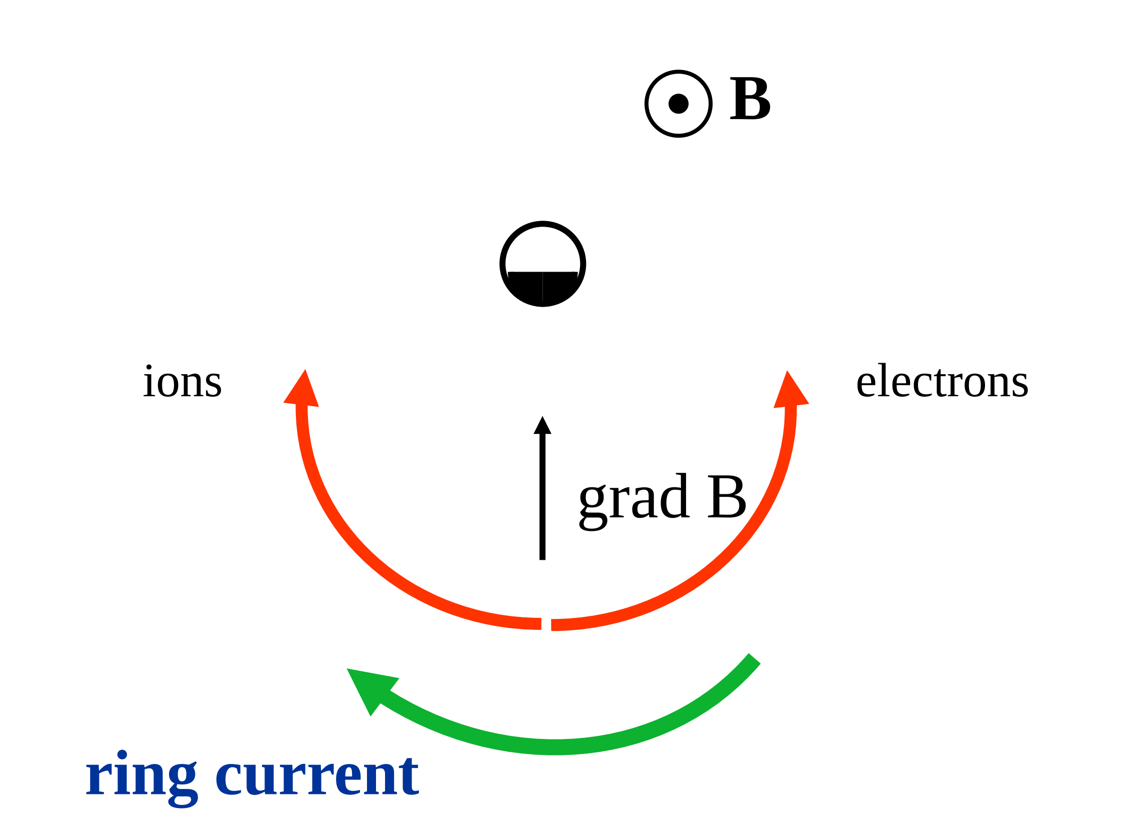
$$= -\mu \nabla \left(\frac{B^2}{B} \right) = -\mu \nabla B$$



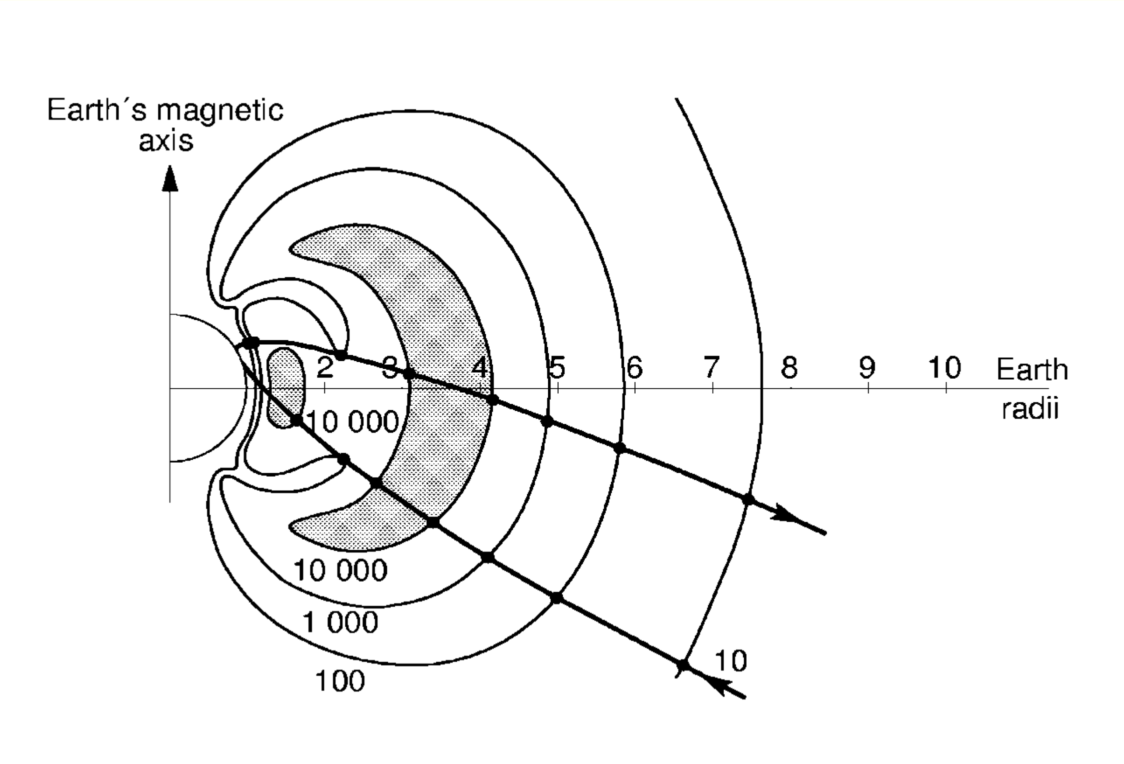
Ring current and particle motion

$$\mathbf{u} = -\frac{\mu \nabla B \times \mathbf{B}}{qB^2}$$

$$\mu = \frac{m v_\perp^2}{2B}$$



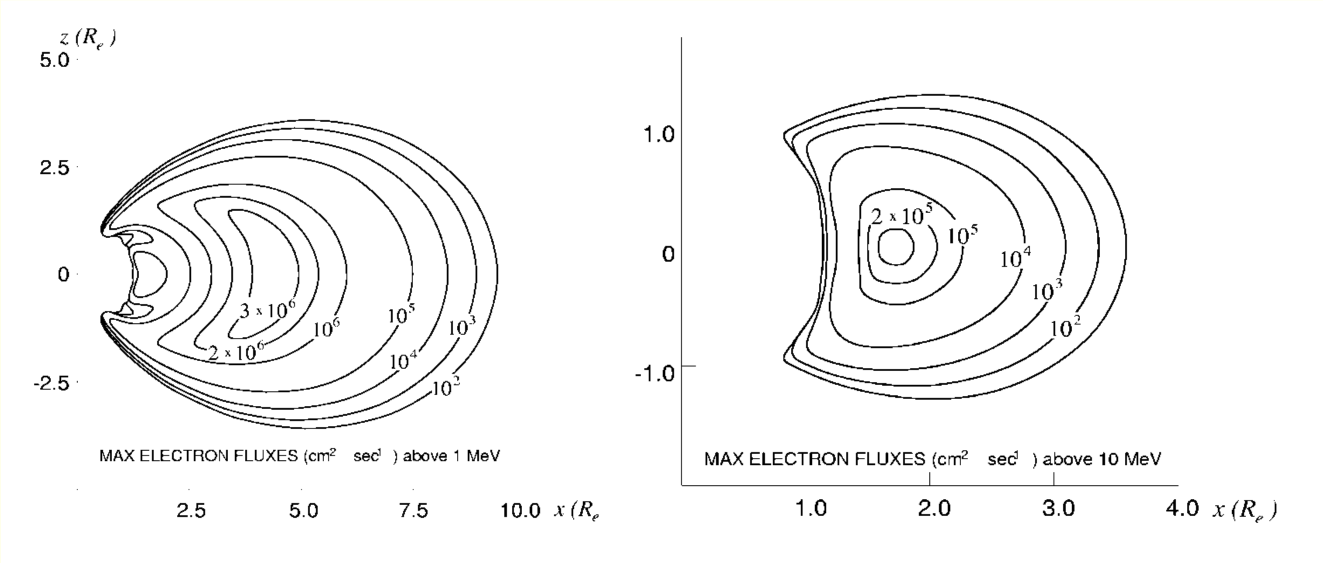
Radiation belts



I. Van Allen belts

- Discovered in the 50s , Explorer 1
- Inner belt contains protons with energies of ~30 MeV
- Outer belt (Explorer IV, Pioneer III): electrons, *W* >1.5 MeV

Radiation belts



- At lower energies there is a more or less continous population of energetic particles in the inner magnetosphere. (Inner part of ***plasma sheet***)

- source: **CRAND** (Cosmic Ray Albedo Neutron Decay).
- a danger for satellites and astronauts.
- associated with a current (***ring current***) which distorts the inner part of the geomagnetic field.

CRAND (Cosmic Ray Albedo Neutron Decay

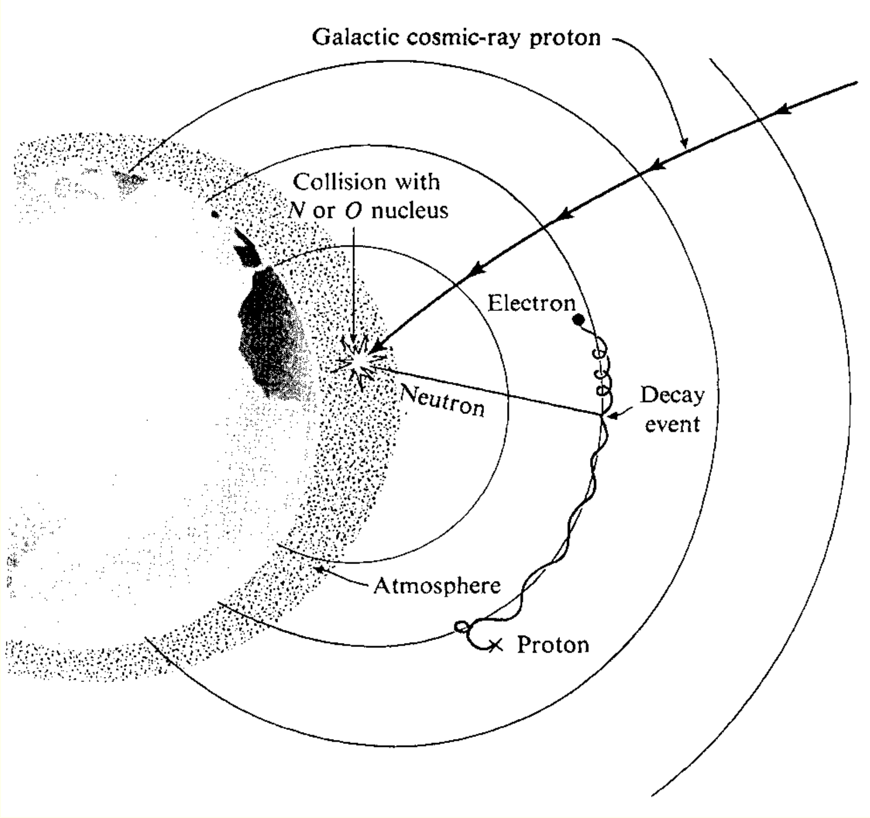
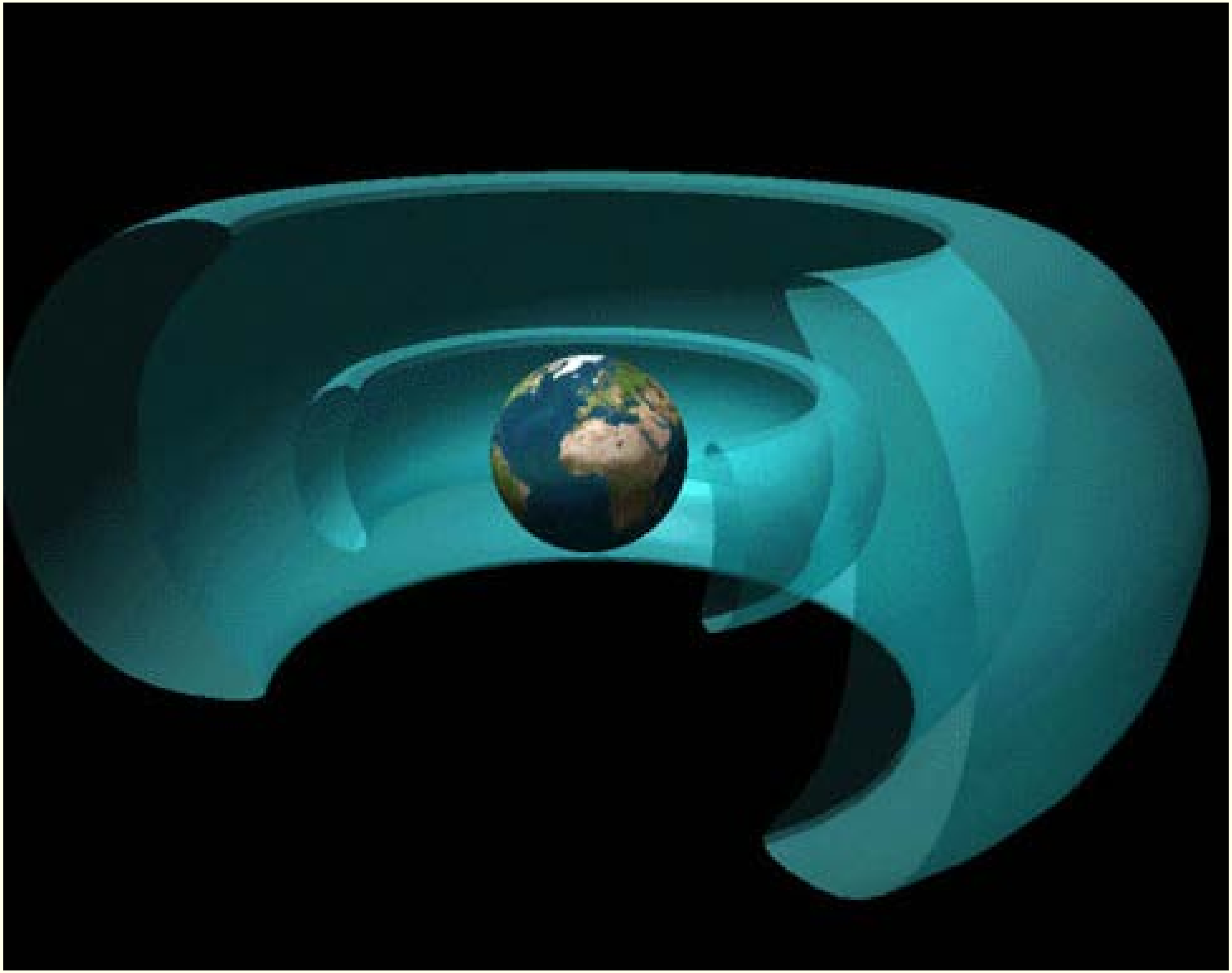


Figure 8. An illustration of the CRAND process for populating the inner radiation belts [Hess, 1968].

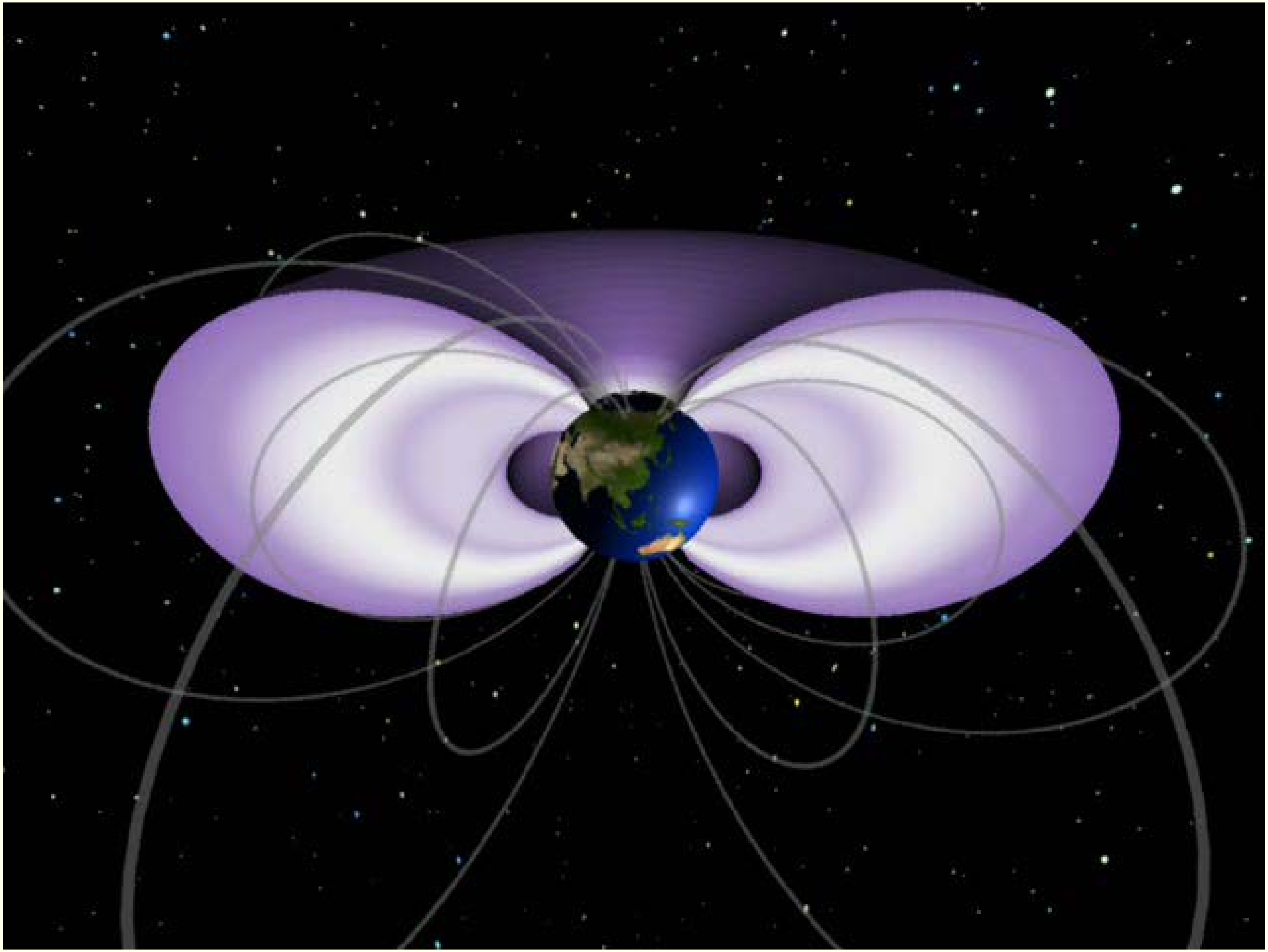
Collisions between cosmic ray particles and the Earth create new particles. Among these are neutrons, that are not affected by the magnetic field. They decay, soom eof them when they happen to be in the radiation belts. The resulting protons and electrons are trapped in the radiation belts.

This contribution to the radiation belts are called the ***neutron albedo***.

Radiation belts



Radiation belts

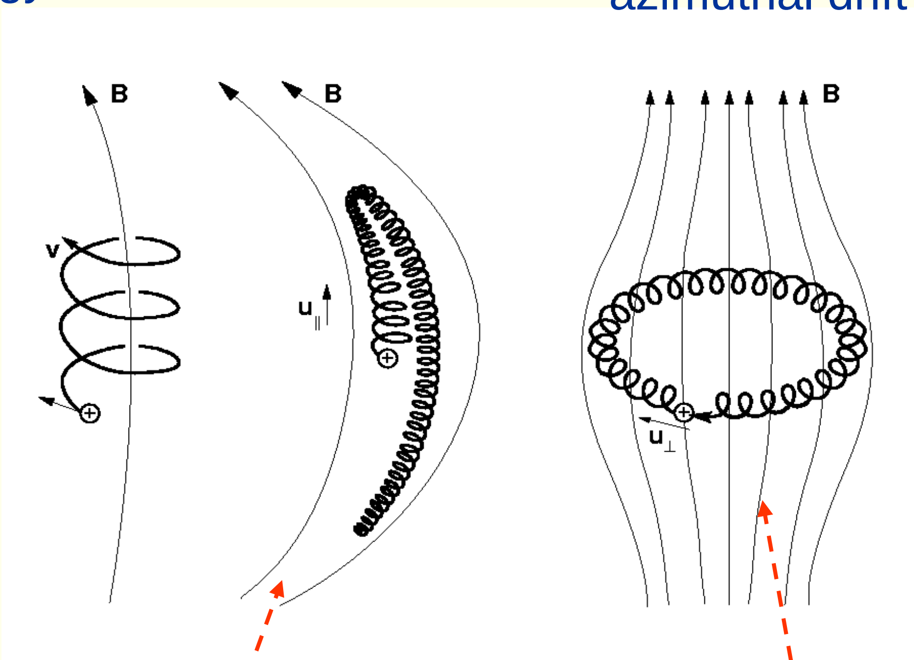


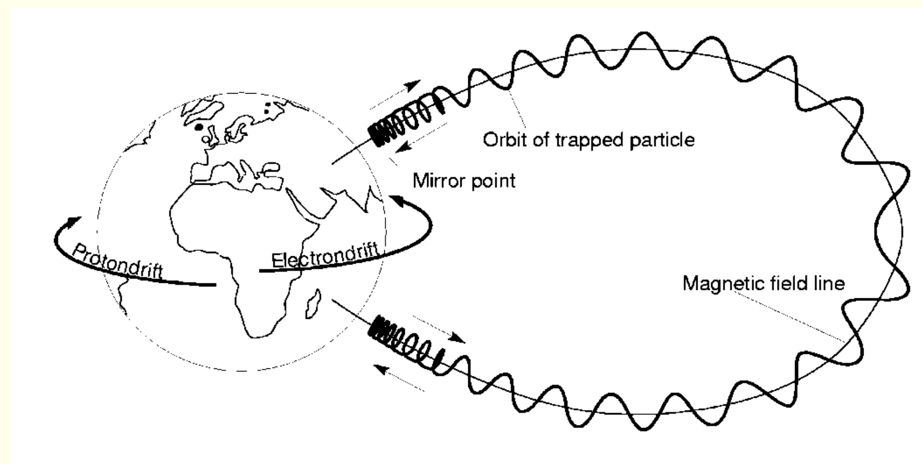
Particle motion in geomagnetic field

longitudinal oscillation

gyration

azimuthal drift

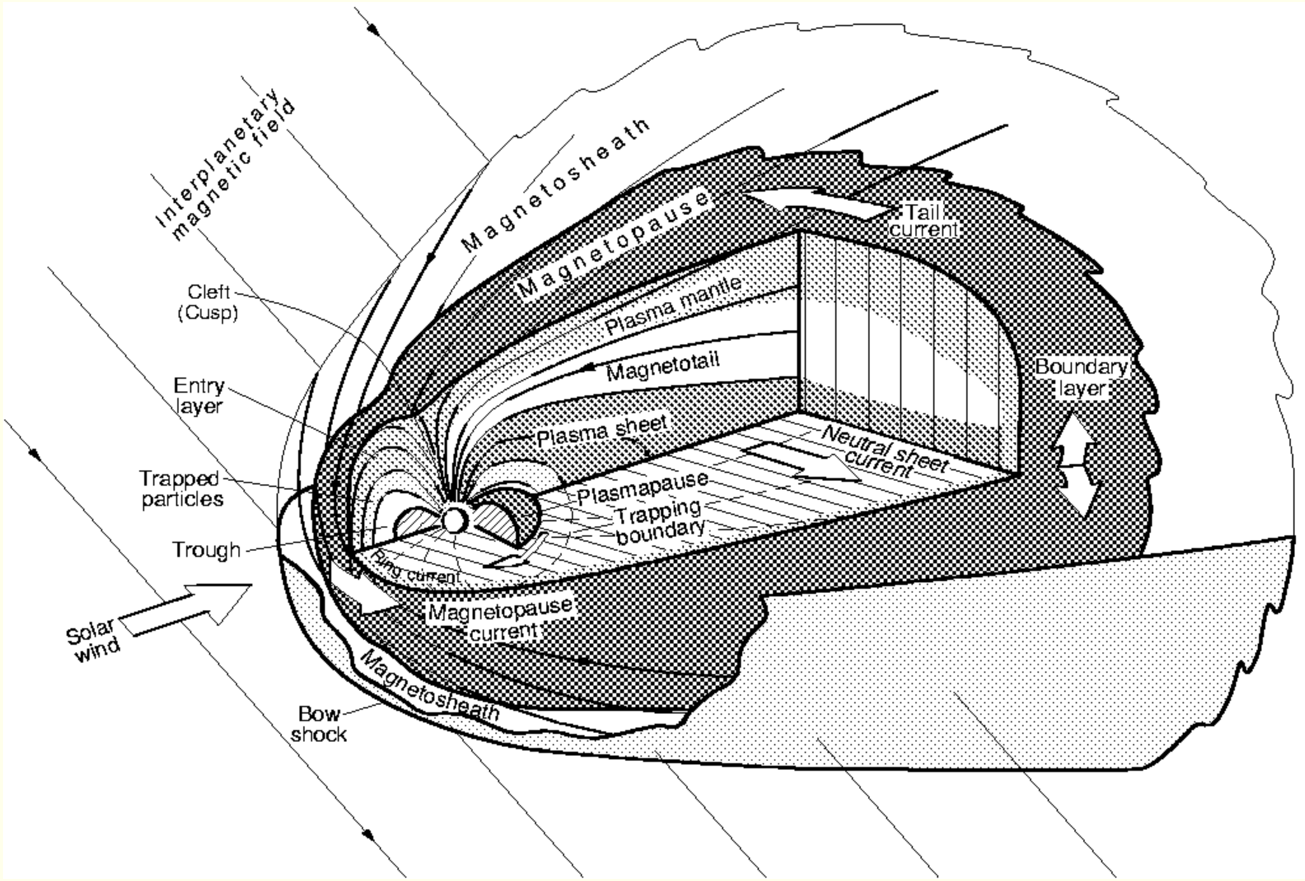




Magnetic mirror

grad B drift

Structure of magnetosphere

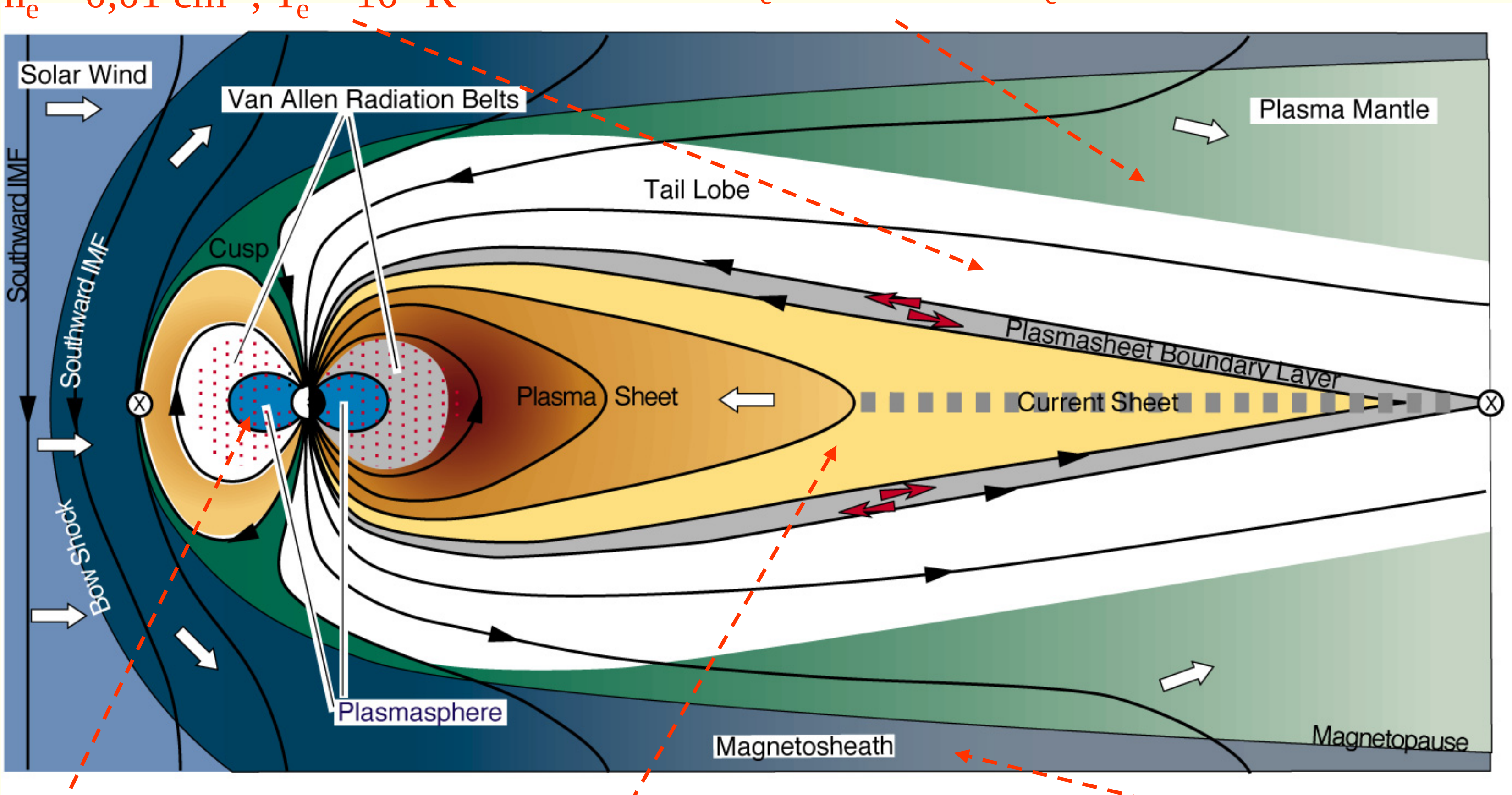


- The plasma in the is made up of approximately equal parts of H^+ and O^+.
- Plasma populations organized by geomagnetic field.
- Particles will mirror between northern and southern hemispheres on closed field lines

Magnetospheric structure

polar plumes = tail lobe
$n_e \sim 0{,}01\ cm^{-3}$, $T_e \sim 10^6$ K

plasma mantle
$n_e \sim 0{,}1\text{-}1\ cm^{-3}$, $T_e \sim 10^6$ K

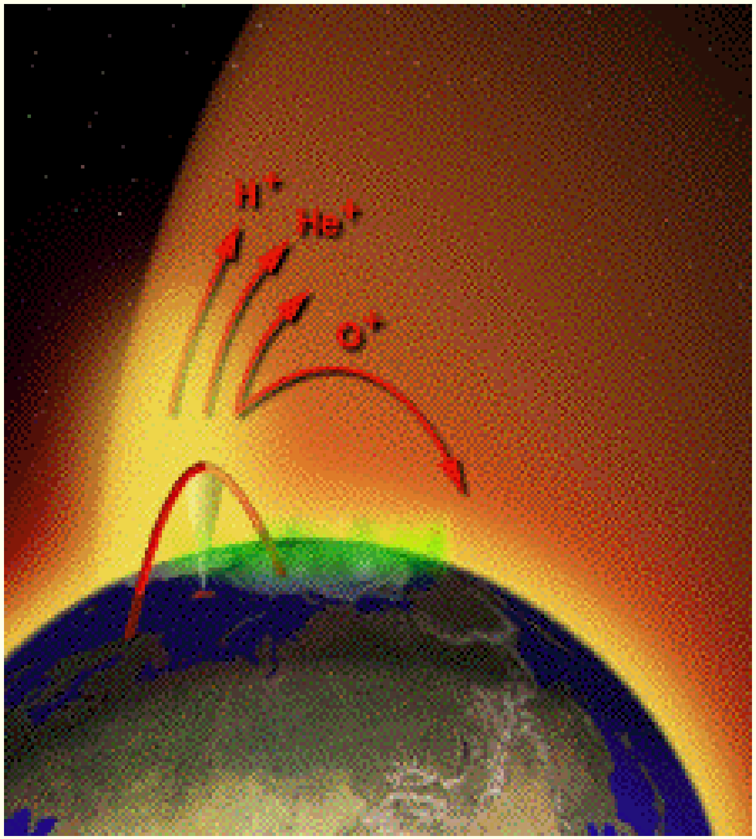


plasmasphere:
$n_e \sim 10\text{-}100\ cm^{-3}$, $T_e \sim 1000$ K

plasma sheet:
$n_e \sim 1\ cm^{-3}$, $T_e \sim 10^7$ K

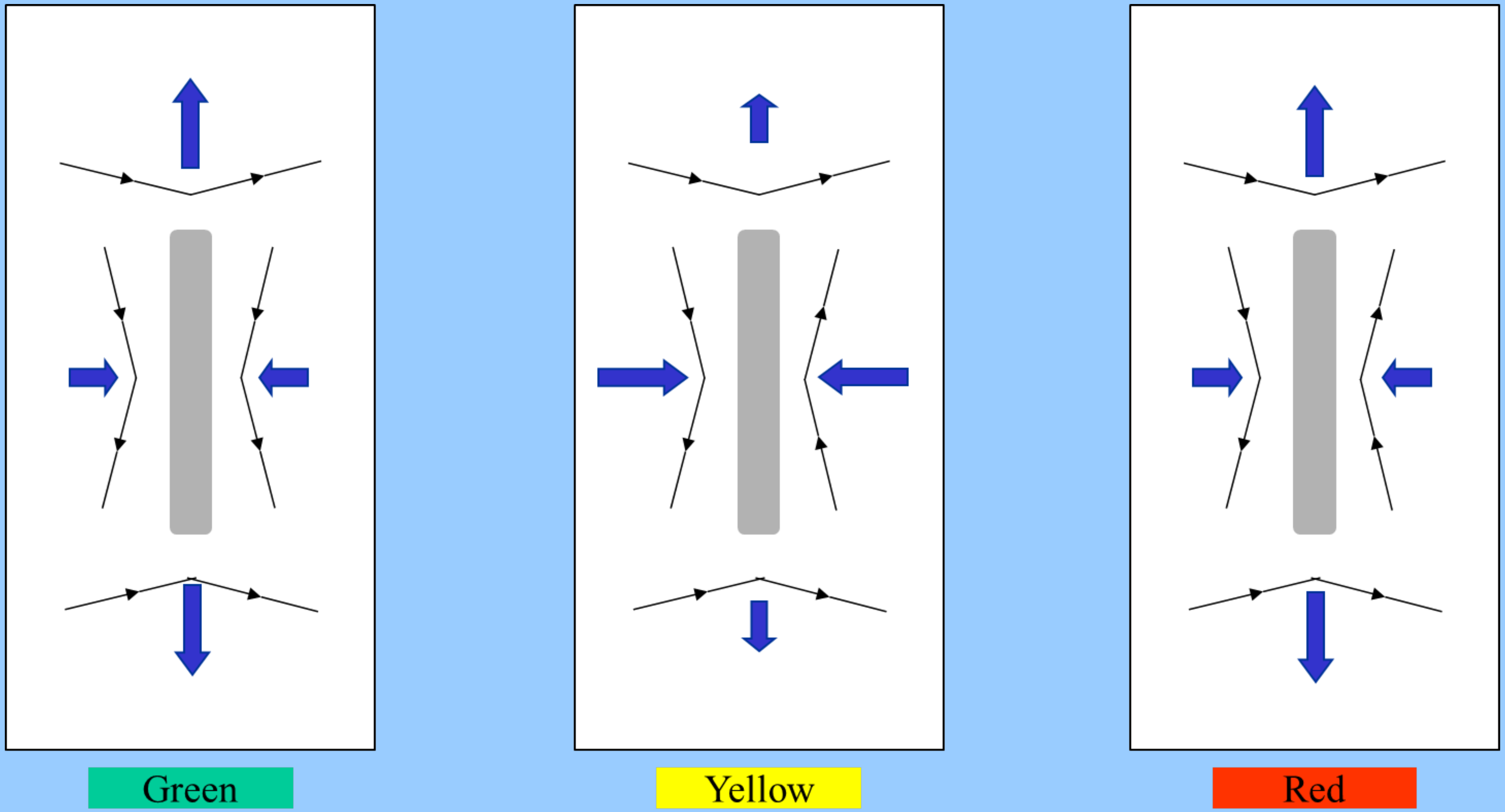
magnetosheath:
$n_e \sim 5\ cm^{-3}$, $T_e \sim 10^6$ K

Outflow from the ionosphere

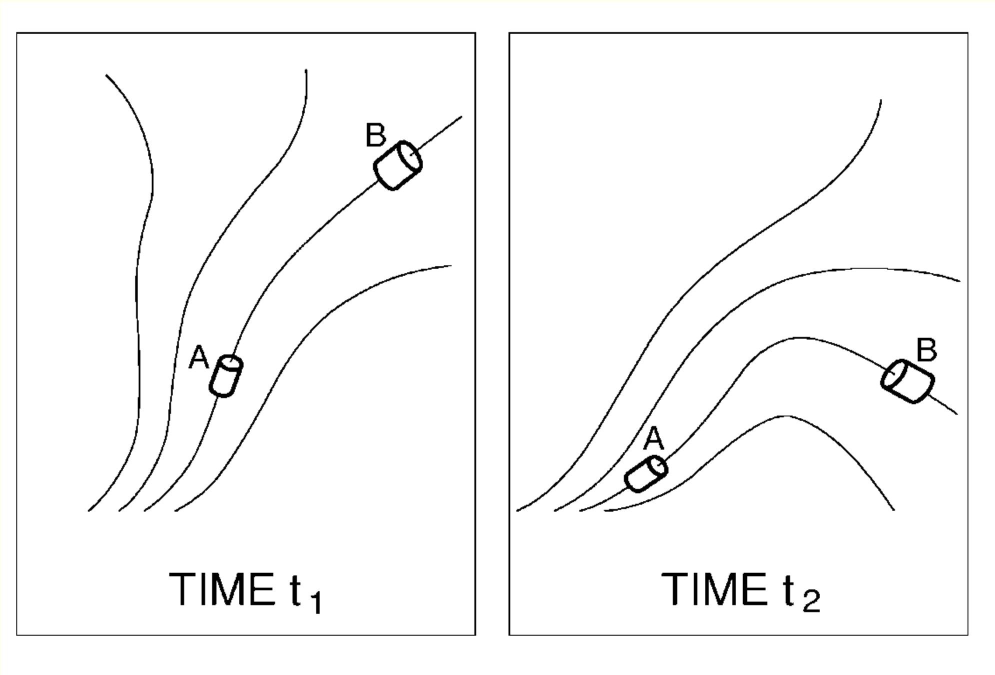


An important source for the magnetospheric plasma. Research is ongoing.

Magnetic reconnection



Frozen in magnetic field lines

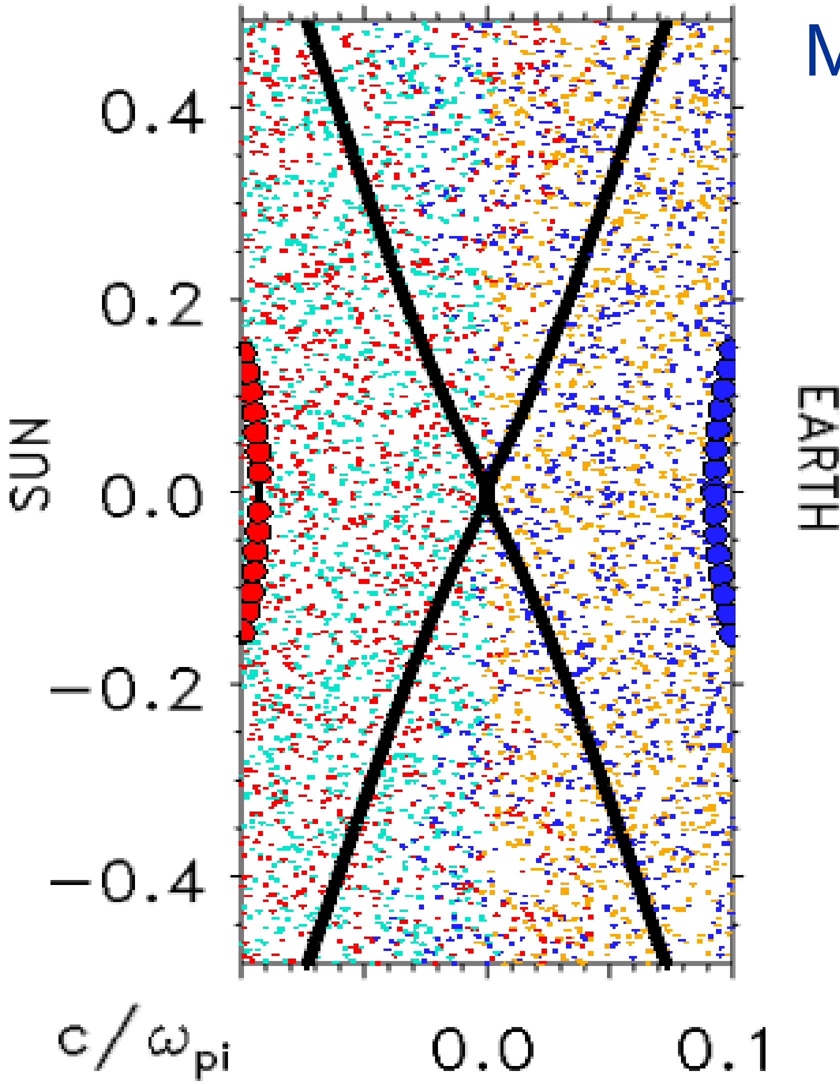


In fluid description of plasma two plasma elements that are connected by a common magnetic field line at time t_1 will be so at any other time t_2.

This applies if the magnetic Reynolds number is large:

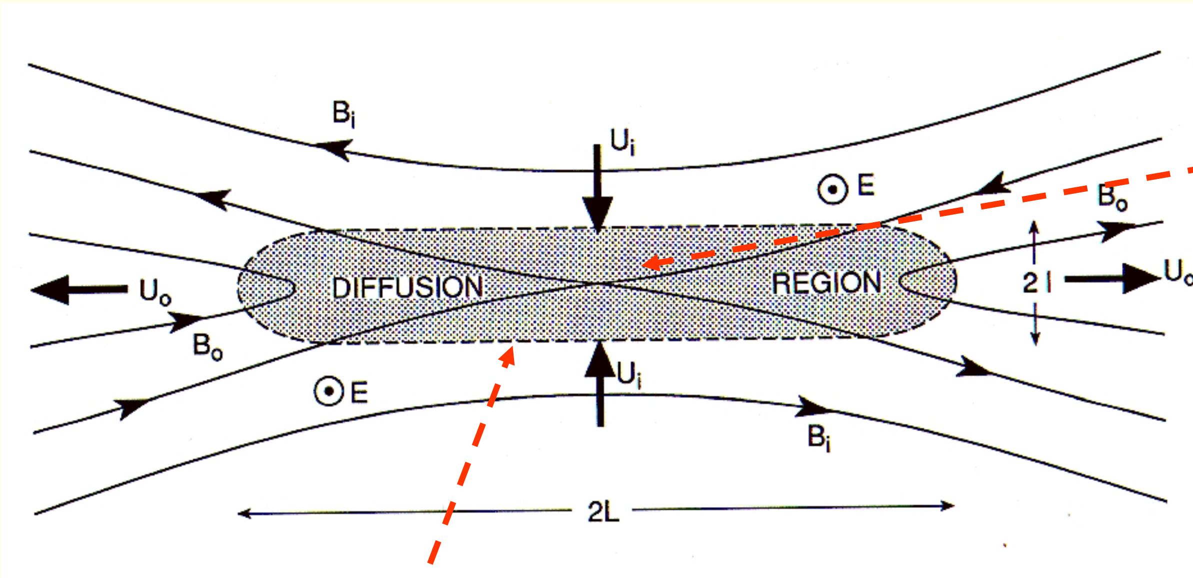
$$R_m = \mu_0 \sigma l_c v_c >> 1$$

An example of the collective behaviour of plasmas.



Magnetic reconnection

Reconnection



In ‘diffusion region’:

$R_m = \mu_0 \sigma l v \sim 1$

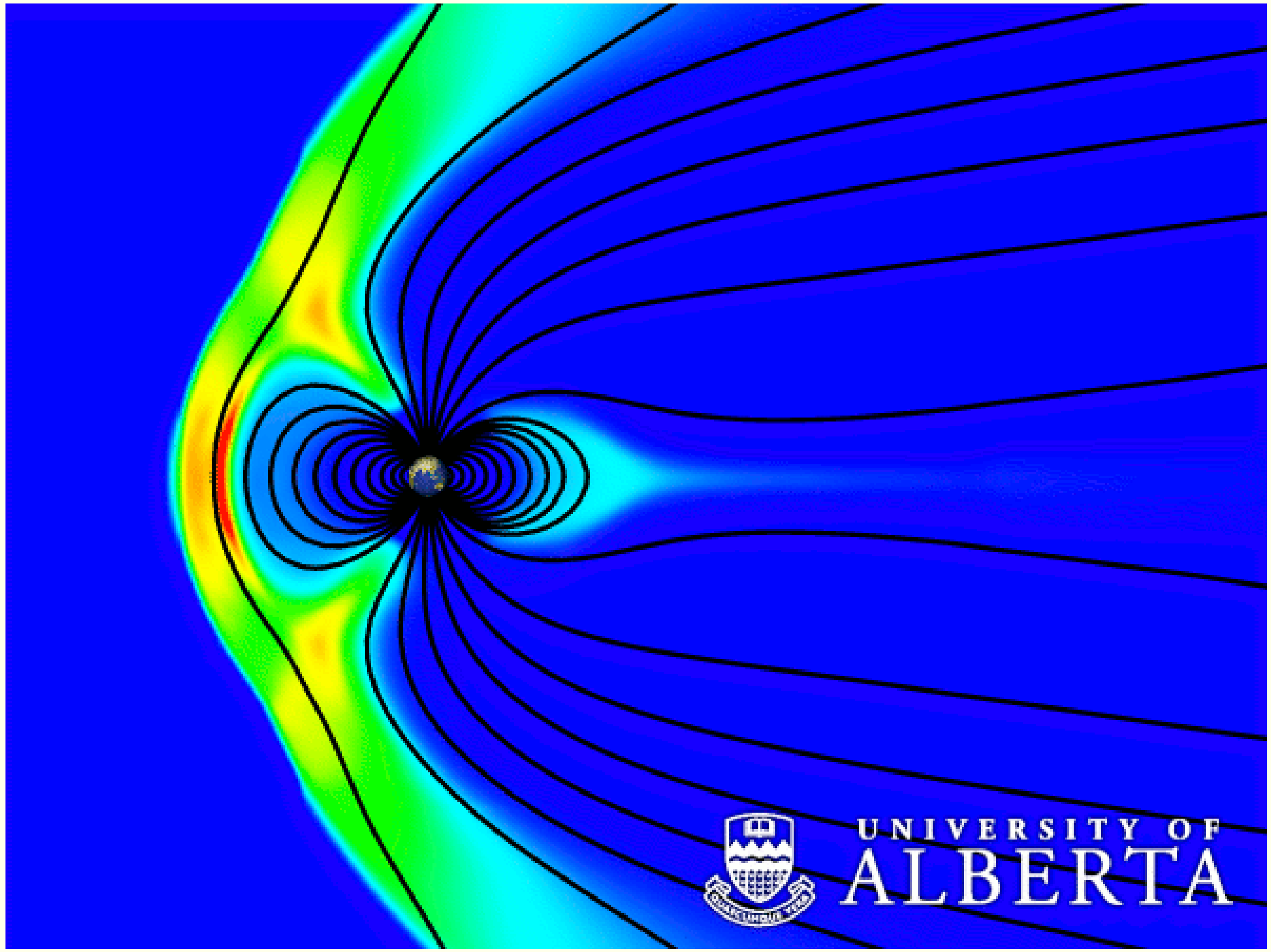
Thus: condition for frozen-in magnetic field breaks down.

A second condition is that there are two regions of magnetic field pointing in *opposite* direction:

- Field lines are “cut” and can be re-connected to other field lines
- **Magnetic energy is transformed into kinetic energy ($U_o >> U_i$)**
- **Plasma from different field lines can mix**

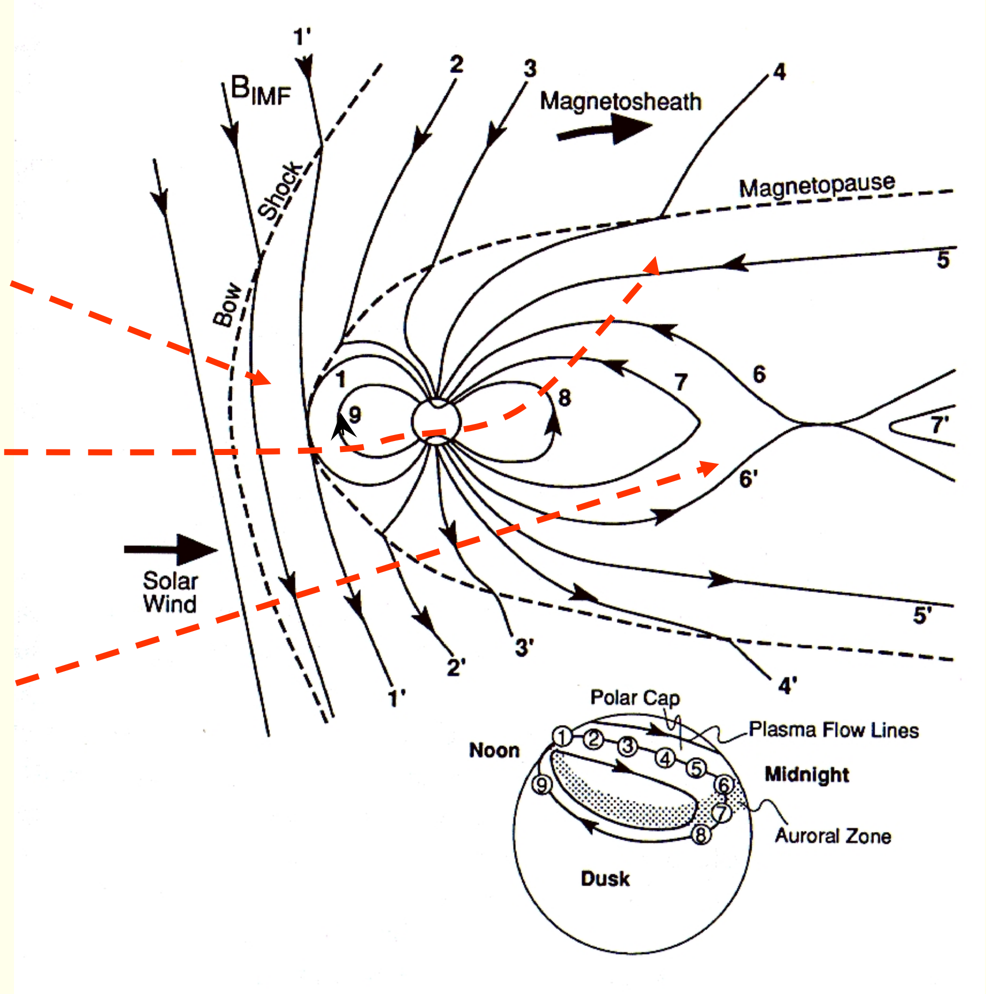
Reconnection and plasma convection

Solar wind

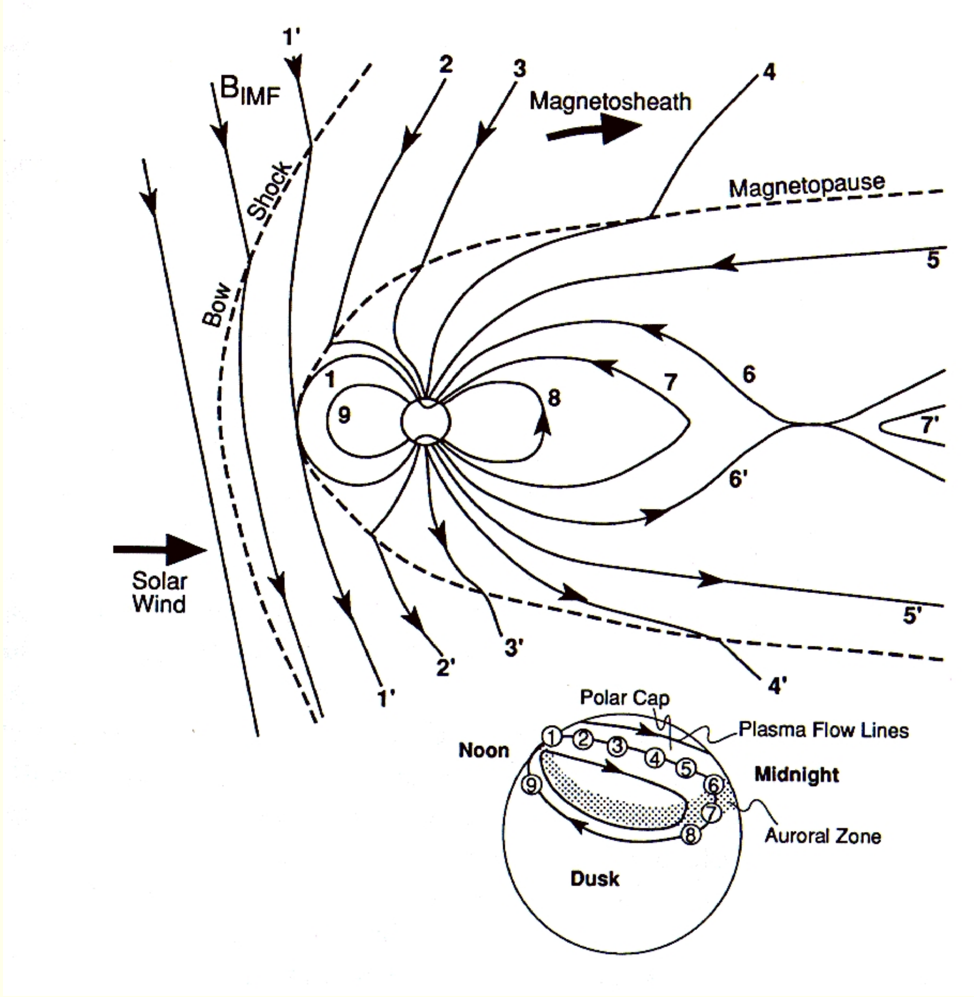


Reconnection och plasma convection

- Reconnection on the dayside "re-connects" the solar wind magnetic field and the geomagnetic field
- In this way the plasma convection in the outer magnetosphere is driven
- In the night side a second reconnection region drives the convection in the inner magnetosphere.
  The reconnection also heats the plasmasheet plasma.

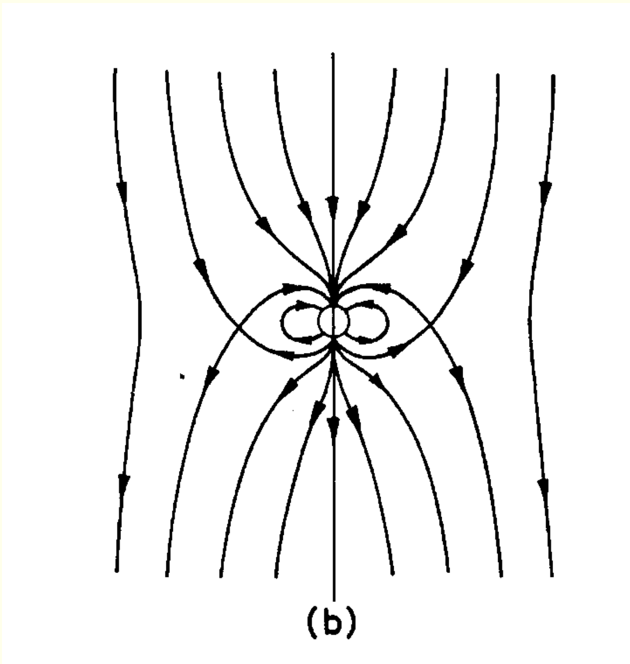


What happens if IMF is northward instead?

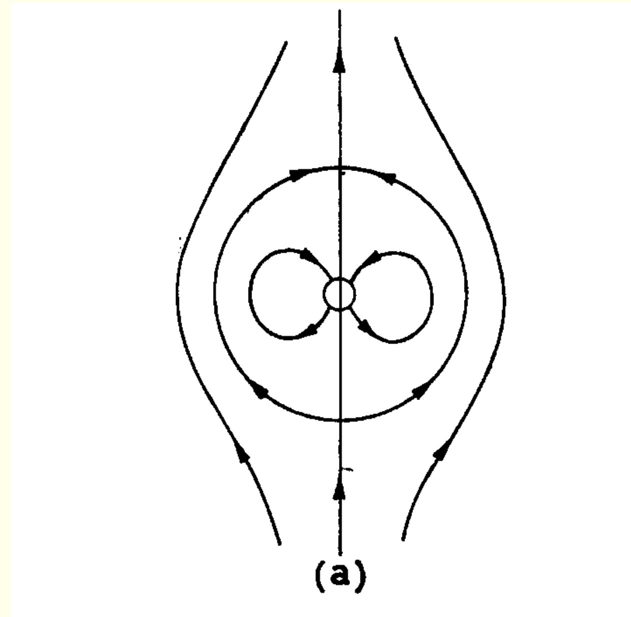


Magnetospheric dynamics

open magnetosphere



closed magnetosphere



southward

Interplanetary magnetic field (IMF)

northward

What do the magnetospheres of the other planets look like?

Planetary magnetospheres

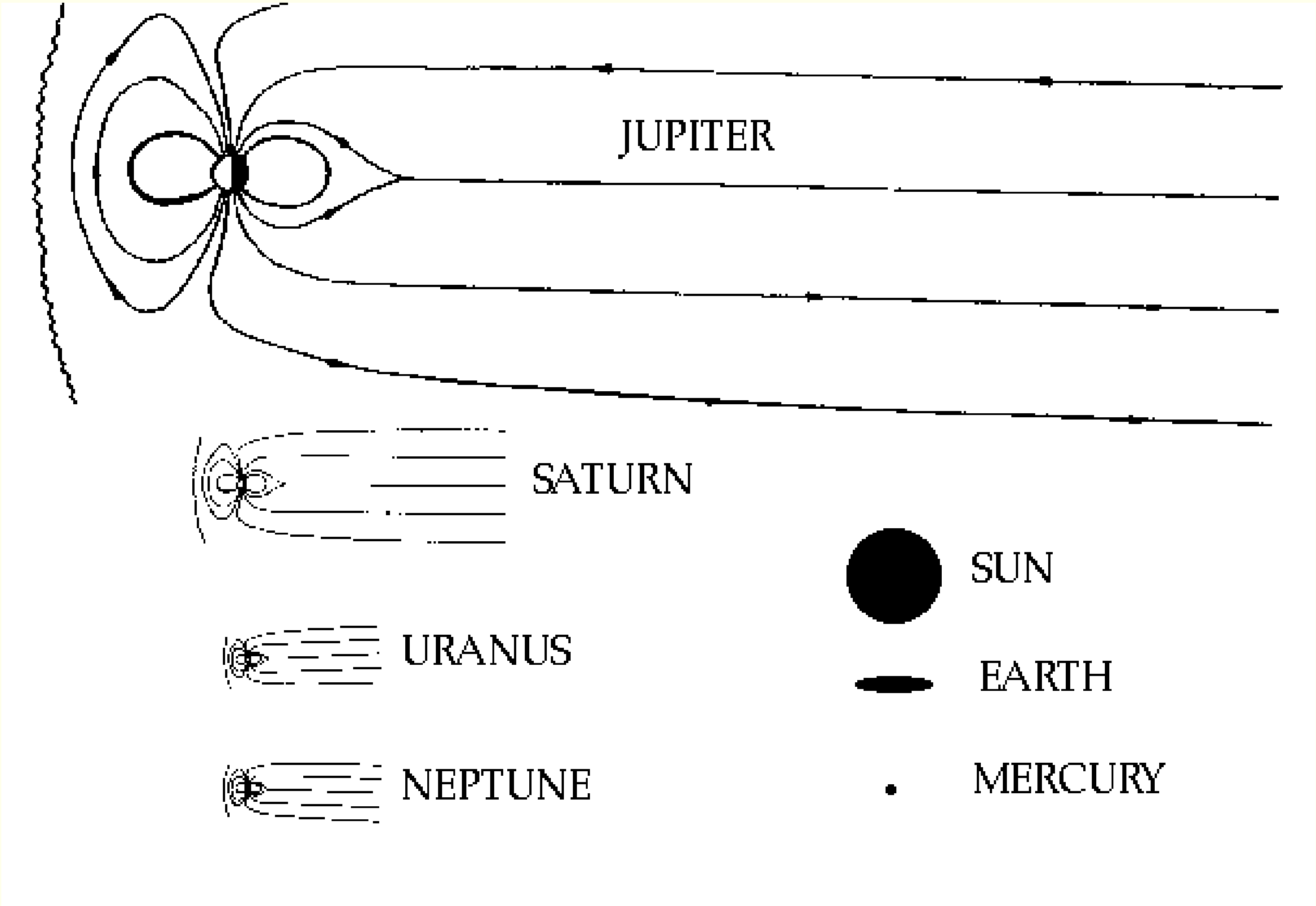
	Radius Earth radii	Spin period (days)	Equatorial field strength (μT)	Magnetic axis direction relative to spin axis	Polarity relative to Earth´s	Typical magneto-pause distance (planetary radii)
Mercury	0.38	58.6	0.35	10^o	Same	1.1
Venus	0.95	243	< 0.03	-	-	1.1
Earth	1.0	1	31	11.5^o	Same	10
Mars	0.53	1.02	0.065	-	Opposite	?
Jupiter	11.18	0.41	410	10^o	Opposite	60-100
Saturn	9.42	0.44	40	$<1^o$	Opposite	20-25
Uranus	3.84	0.72	23	60^o	Opposite	18-25
Neptune	3.93	0.74	20-150*)	47^o	Opposite	26**)

*) The magnetic field differs greatly from a dipole field. The numbers represent maximum and minimum strength at the planetary surface

**) Based on single passage

Very weak magnetic fields

Relative size of the magnetospheres



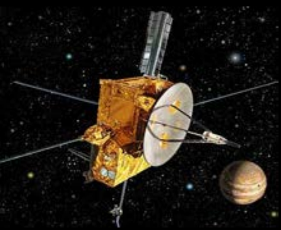
Comparative magnetospheres

In situ observations

Space probe	Planets	Observations
Mariner 10	Mercury	*1974 – 1975*
*Messenger **	Mercury	*2008 –*
Pioneer 10,11	Jupiter, Saturn	*1973 – 1979*
Voyager 1,2	Jupiter, Saturn, Uranus, Neptune	*1977 – 1989*
Ulysses	Jupiter	*1992*
*Galileo**	Jupiter	*1995 – 2003*
*Cassini**	Jupiter, Saturn	*2004 –*
New Horizons	Jupiter	*2007*

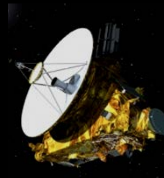
* Orbiters

Mariner 10

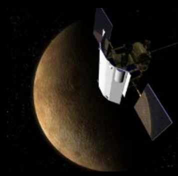


Ulysses

Galileo



New Horizons



Messenger

Pioneer 10

Pioneer 11

Voyager 1 and 2

Cassini

Mercury

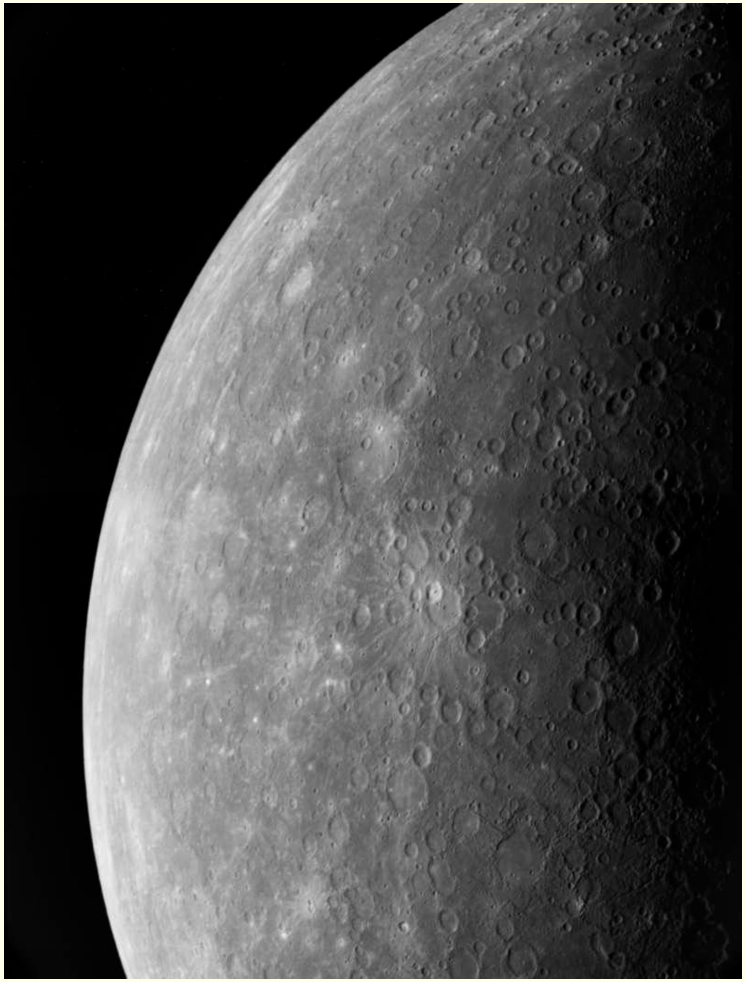
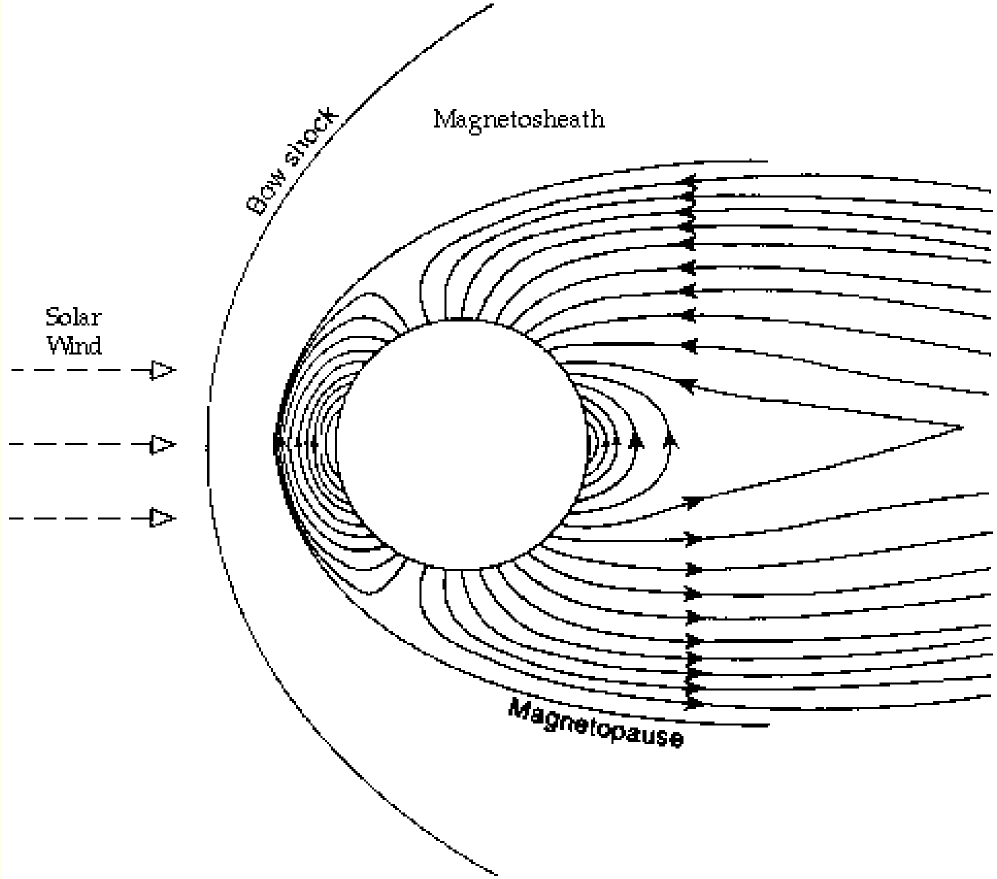


Photo from Mariner 10

- $r_M = 0{,}38\ r_E$
- $m_M = 0{,}06\ m_E$
- distance from sun: 0,39 AU
- no or extremely thin atmosphere

Mercury's magnetosphere



- rather large magnetic field
- the high solar wind density close to the sun makes the magneto-sphere very small
- no ionosphere

Venus

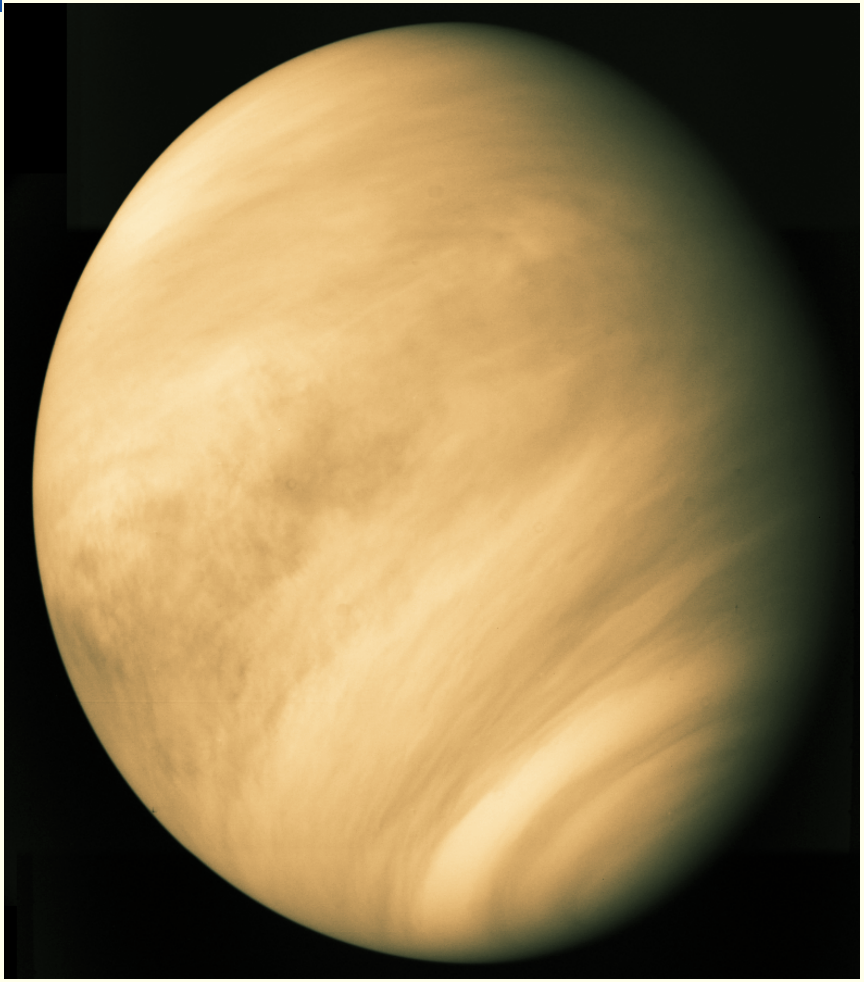
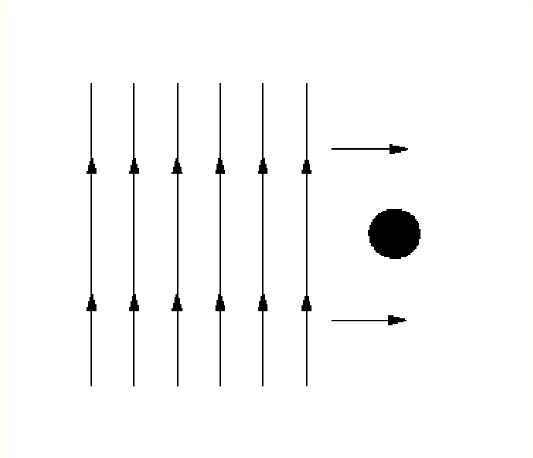


Photo from Galileo

- $r_V = 0{,}95\ r_E$
- $m_V = 0{,}82\ m_E$
- distance from sun : 0,72 AU
- very dense atmosphere
    - ~ 90 atm
    - 96% CO_2
- very weak magnetic field

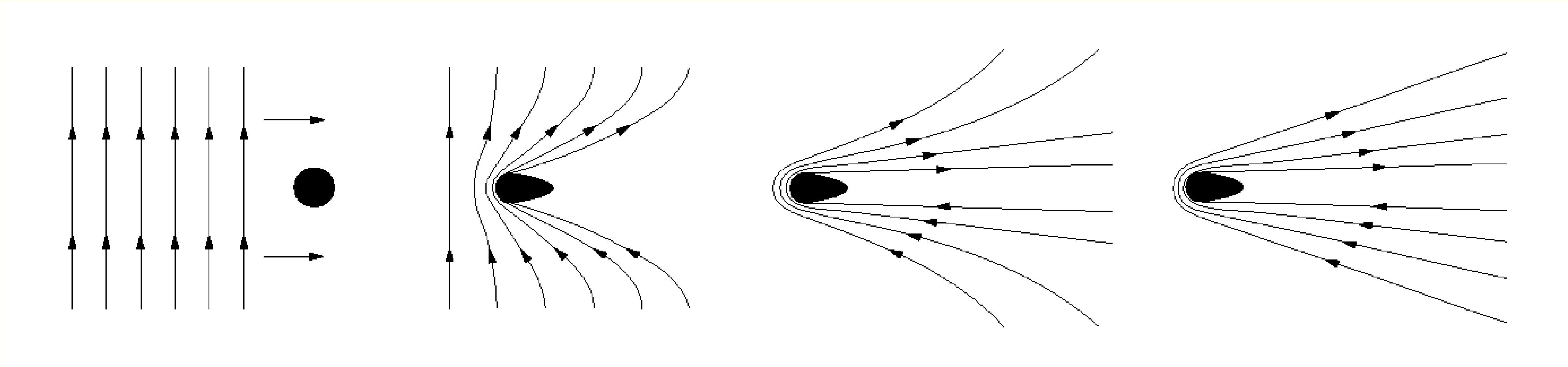
Comets, induced magnetotail



- The ***coma*** of the comet is ionized when the comet gets close to sun

What will happen when the interplanetary magnetic field (which is frozen into the solar wind) hits the coma?

Comets, induced magnetotail



- The ***coma*** of the comet is ionized when the comet gets close to sun
- This plasma stops the field lines from passing the comet, due to the frozen-in magnetic field
- Venus and Mars have similar induced, weak ”magnetospheres”

Mars

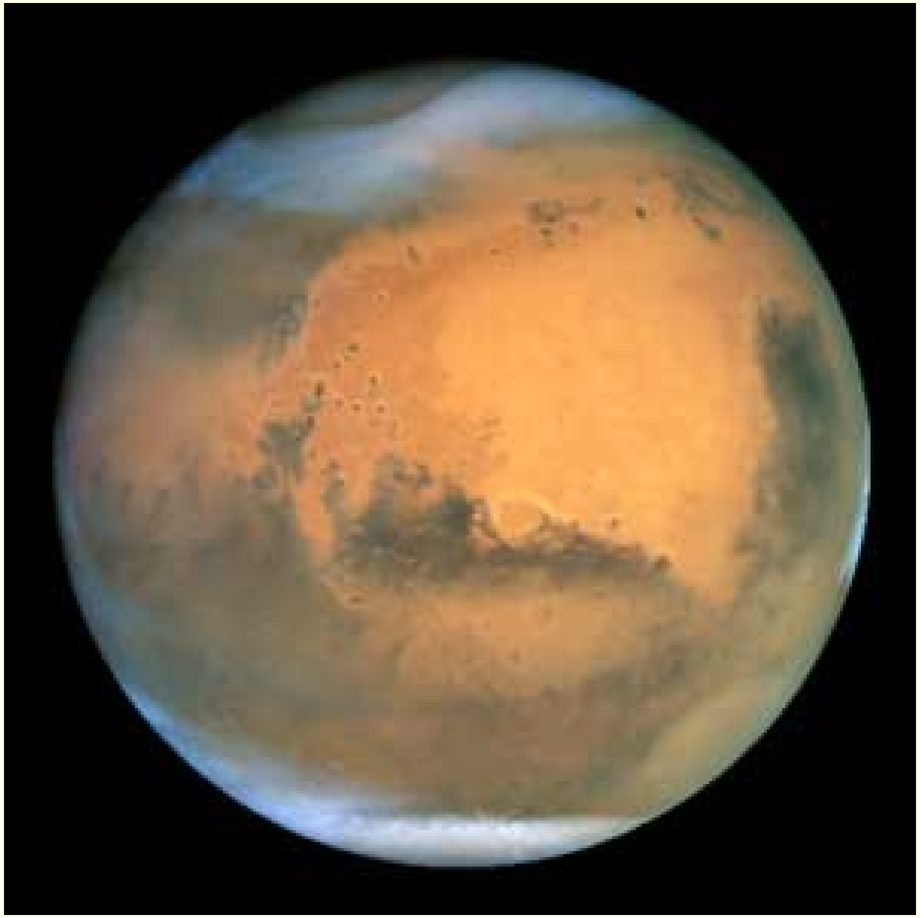


Photo from Hubble Space Telescope

- $r_M = 0,53\ r_E$
- $m_M = 0,11\ m_E$
- distance from sun : 1,52 AU
- very thin atmosphere
  - ~ 0.01 atm.
  - 95% CO_2
- very weak magnetic field

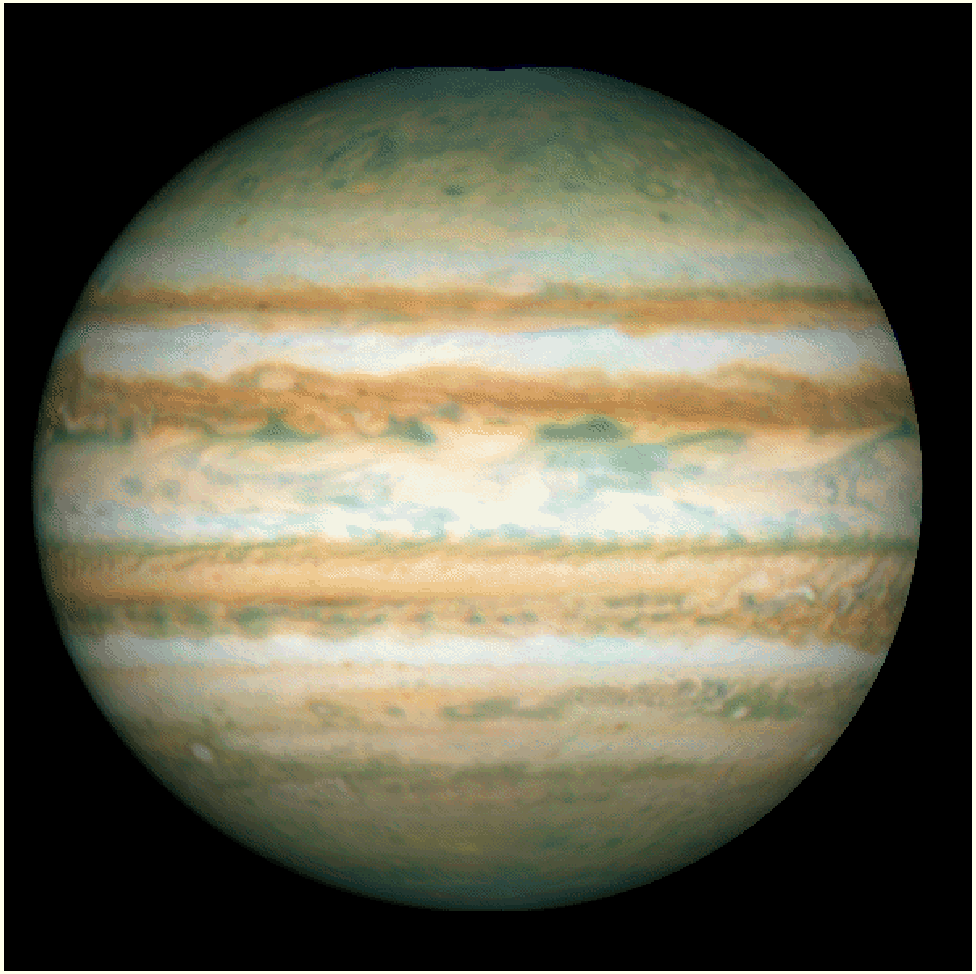
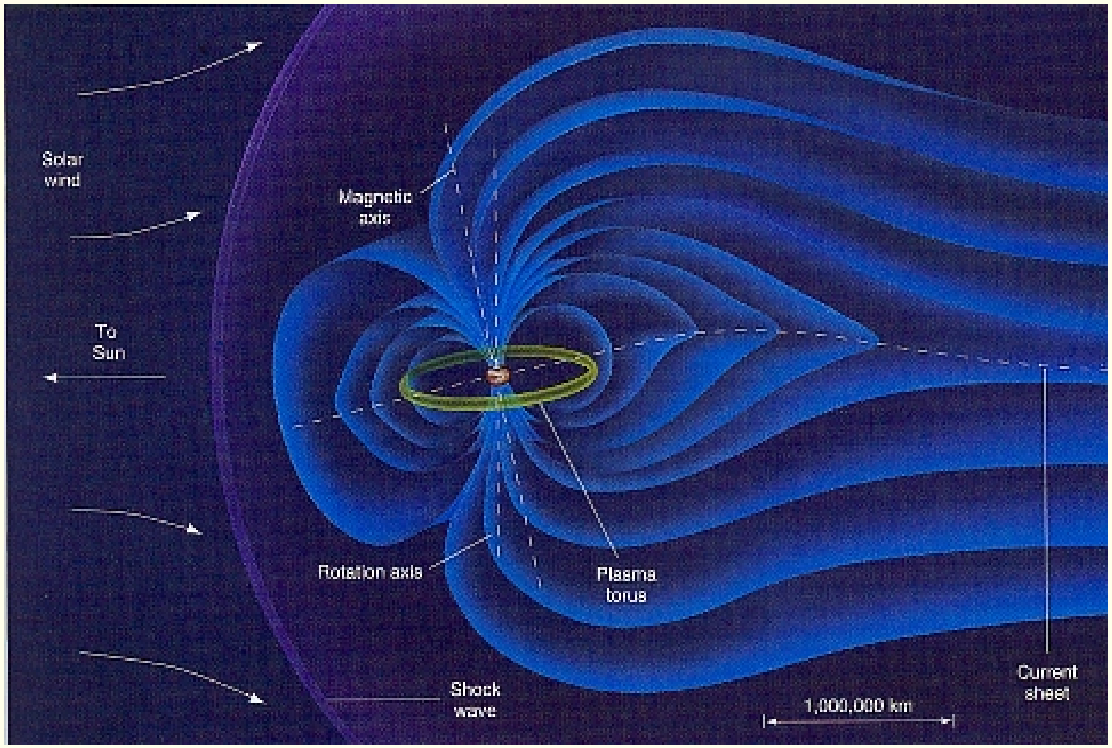


Photo from Hubble Space Telescope

Jupiter

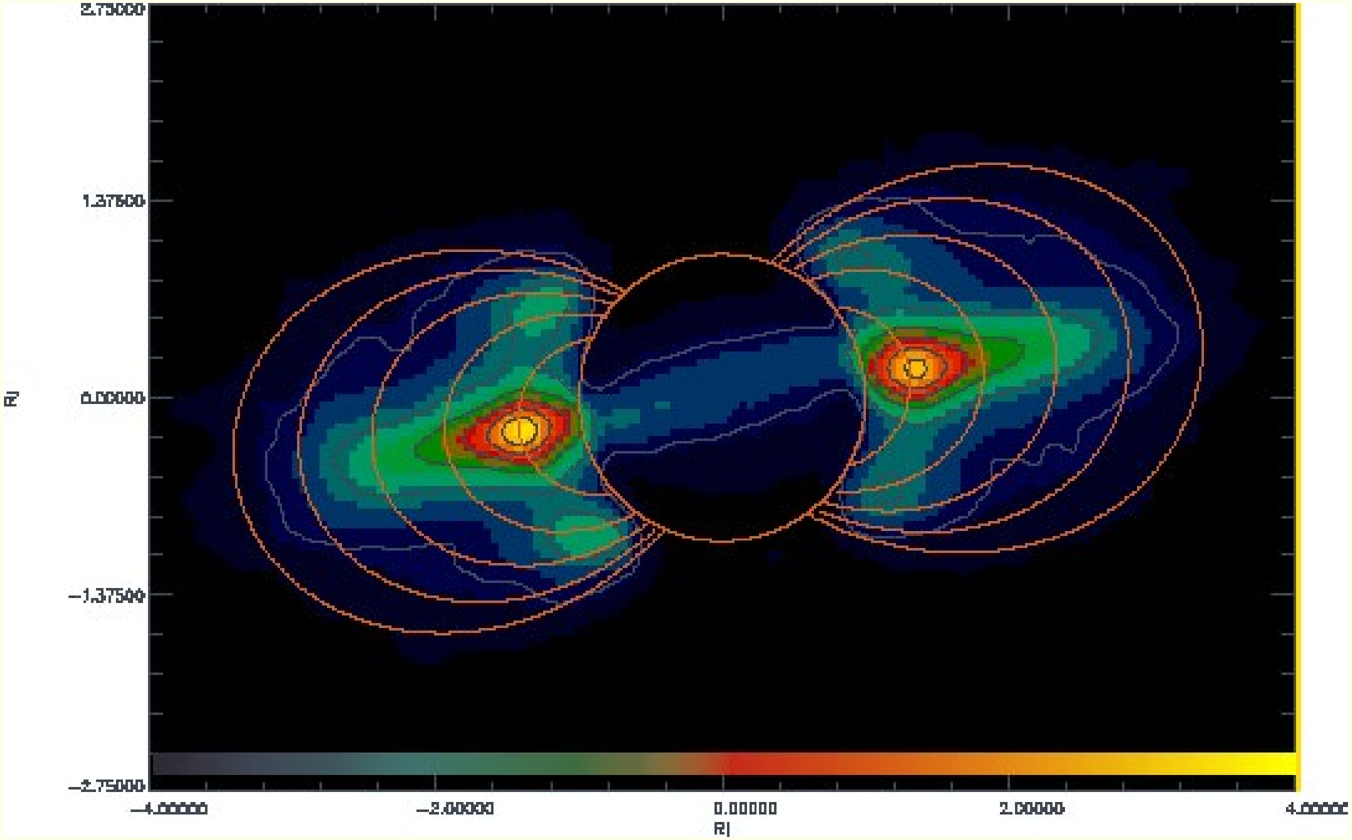
- $r_J = 11{,}2\ r_E$
- $m_J = 318\ m_E$
- distance from sun: 5,20 AU
- gas giant with very dense atmosphere containing hydrogen, helium, ammonium, methane, etc.
- ~ 60 moons (+ weak ring system)

Jupiter's magnetosphere



- high plasma density, Io is an important source
- plasma pressure thus becomes an important factor for balancing the solar wind pressure
- the plasma co-rotates with Jupiter, which gives the magnetosphere a flattened look

Syncrotron radiation from Jupiter's radiation belts

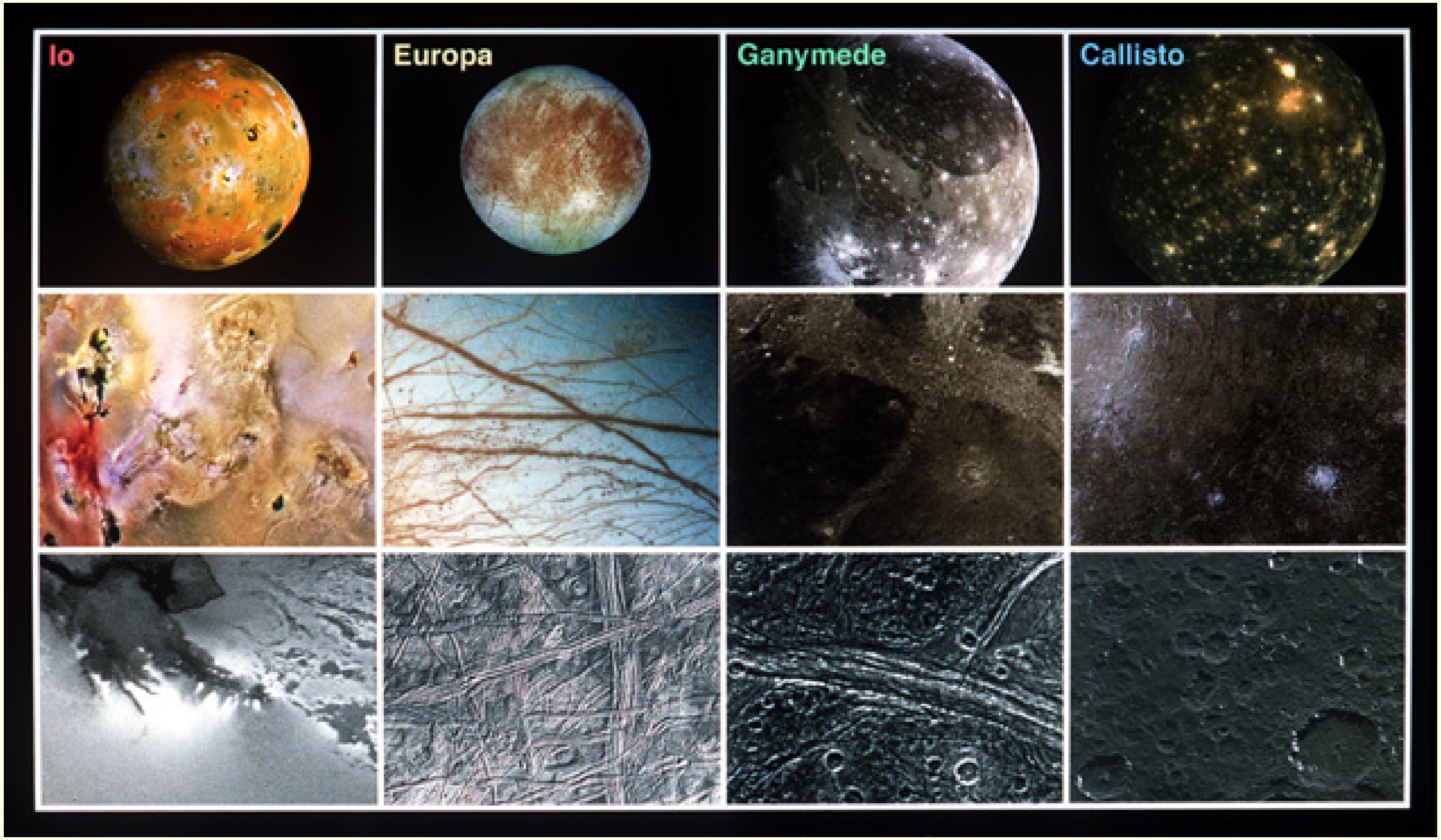


- Gyrating electrons emit "syncrotron radiation" with frequencies ~ $f_{ce} = eB/(2\pi m_e)$
- The emitted power is proportional to the electron temperature:

  $P = CT_e$

- In this way you can get a picture of the radiation belts

Galilean satellites



Volcanic activity, source for plasma.

Oceans under the ice?

Has its own magnetosphere the size of Mercury's

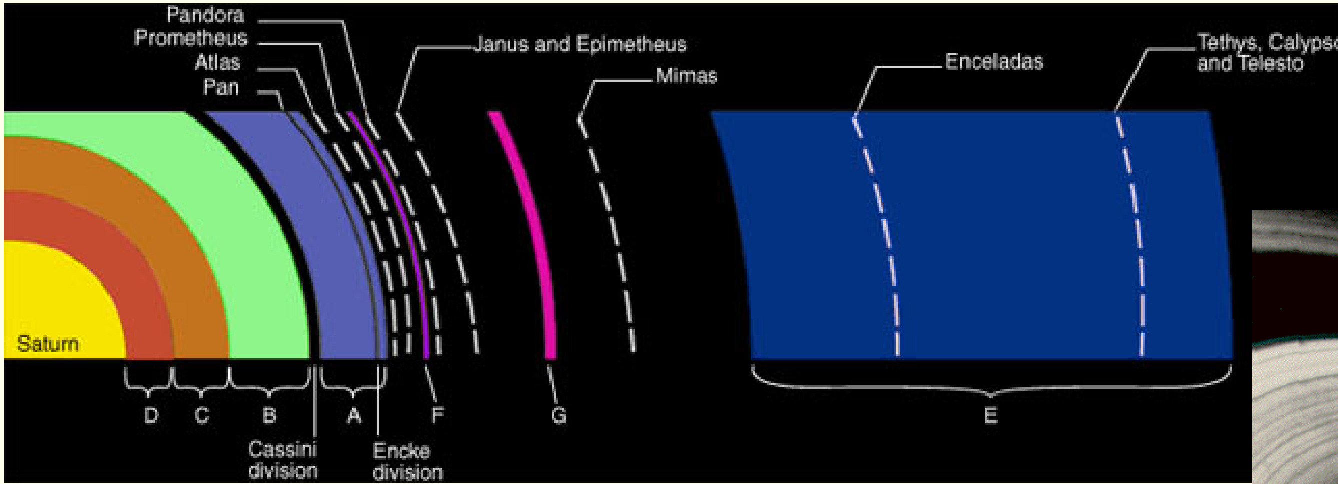
Weak magnetic field

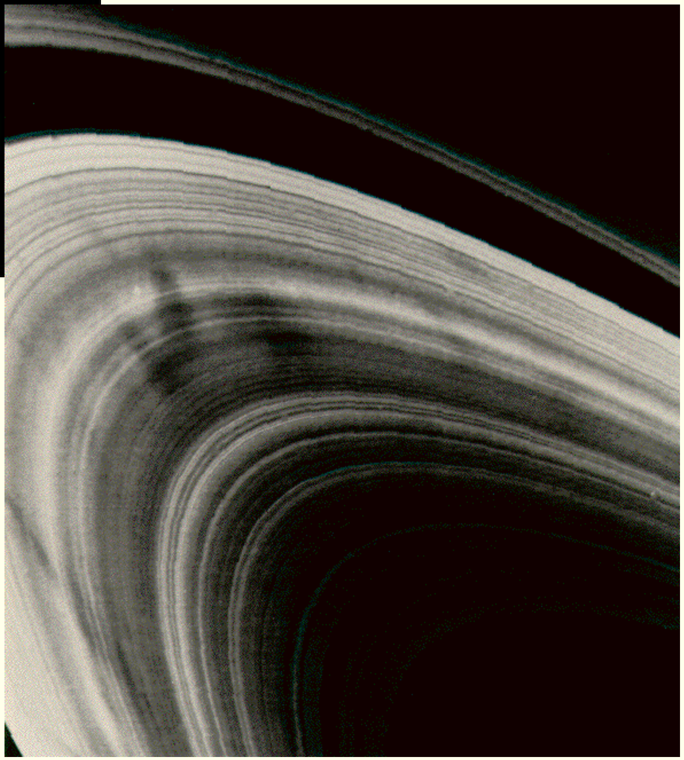
Saturn

Photo from Hubble Space Telescope

- $r_s = 9{,}42\ r_E$
- $m_S = 95\ m_E$
- distance from sun: 9.53 AU
- gas giant with very dense atmosphere containing sulphur, hydrogen, helium, ammonium, methane, etc.
- ~ 31 moons + ring system

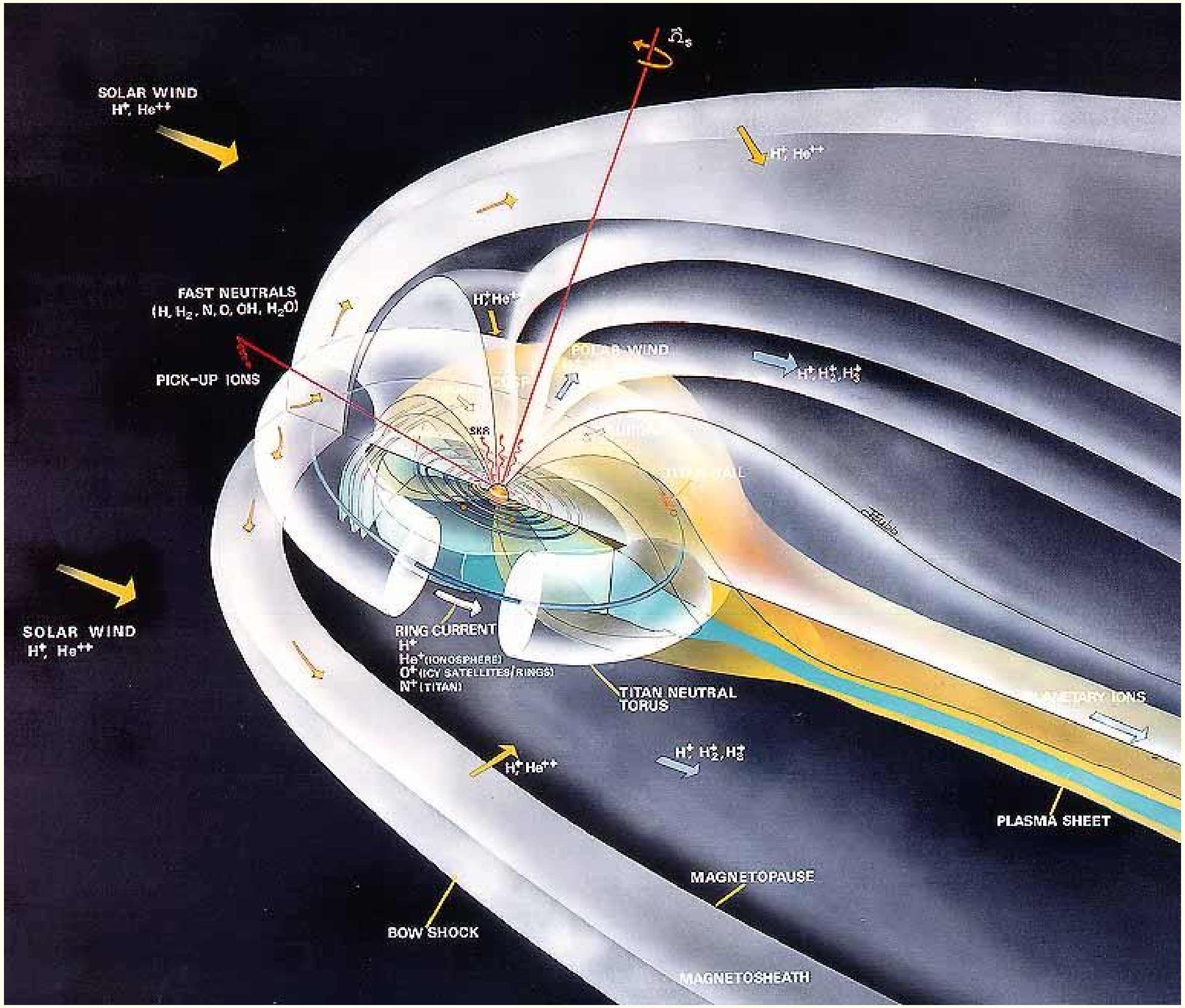
Saturn’s rings



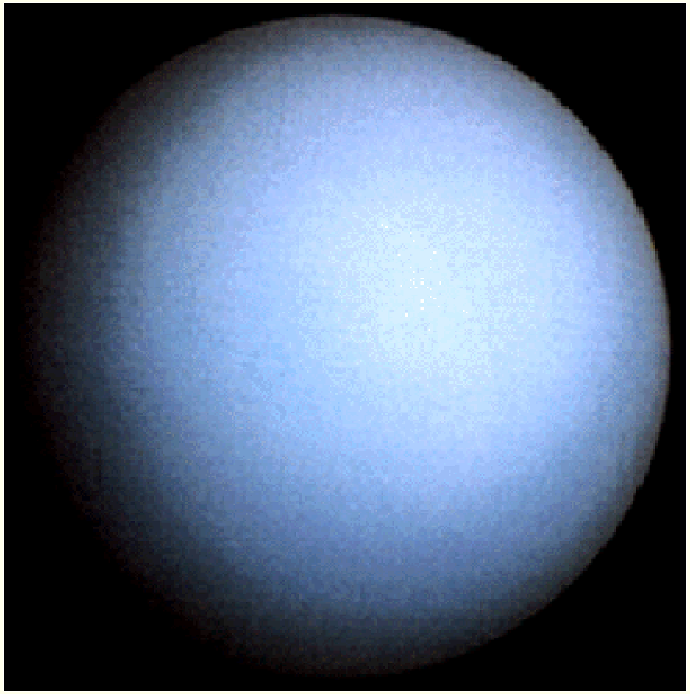


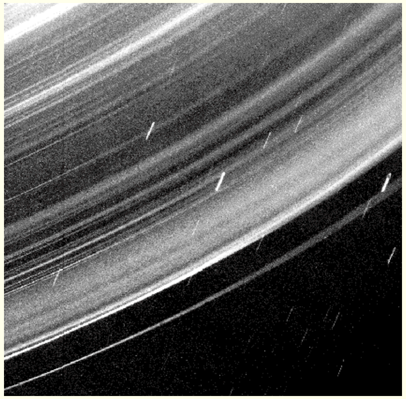
- ring systemet is made up of ice and mineral particles from ~ 1 cm to ~ 1 km
- rings are only 1.5 km thick

Saturn's magnetosphere



- ring systemet is both source and sink for plasma and limits the size of the plasmasphere
- plasma pressure important for balancing solar wind pressure, just as for Jupiter
- Intense radiation belt with 50 MeV protons

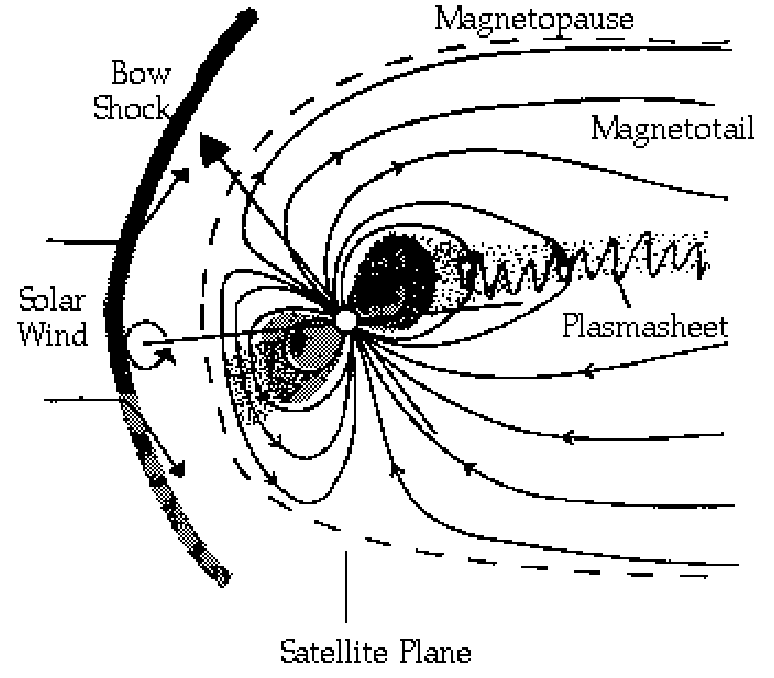


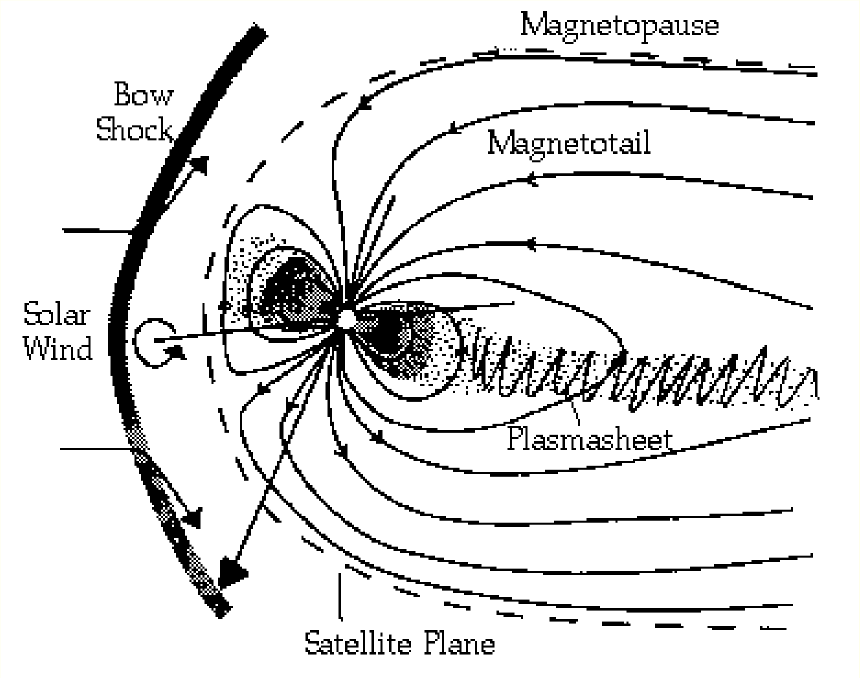


Uranus

- $r_U = 3{,}84\ r_E$
- $m_U = 14{,}5\ m_E$
- distance from sun: 19,2 AU
- gas giant with very dense atmosphere containing mainly methane
- 20 moons + ring system

Uranus





- Uranus' rotational axis almost in ecliptic plane
- magnetic field axis makes an angle of around 60° with rotational axis
- this gives enourmous daily variations of the structure of the magnetosphere

Neptune

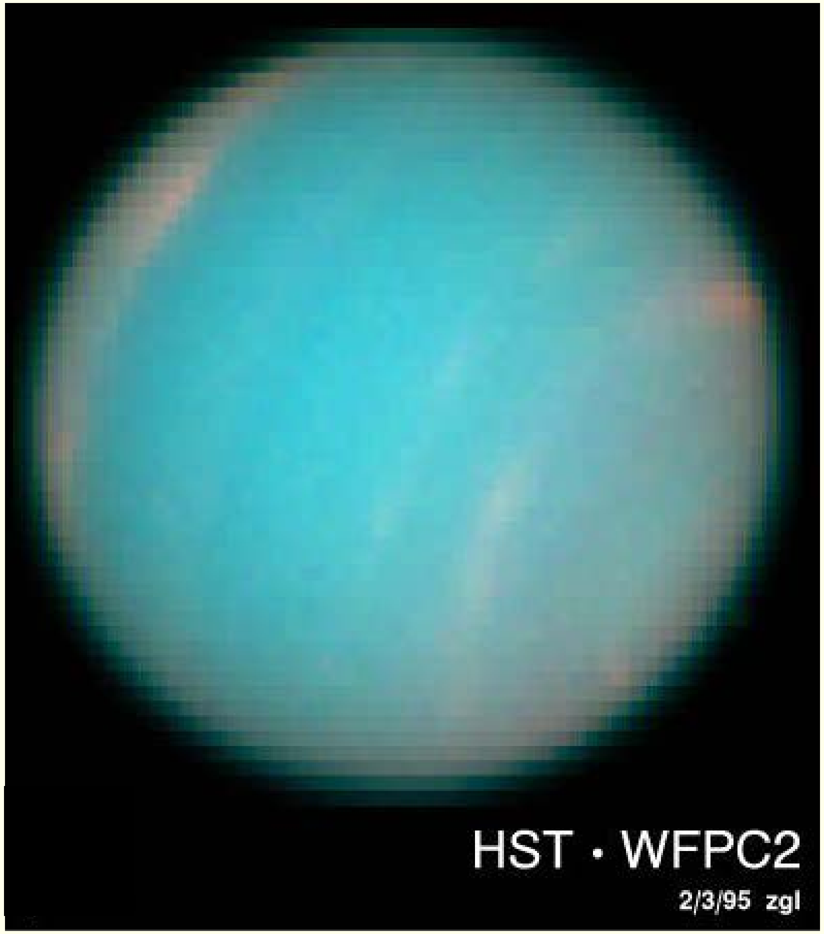
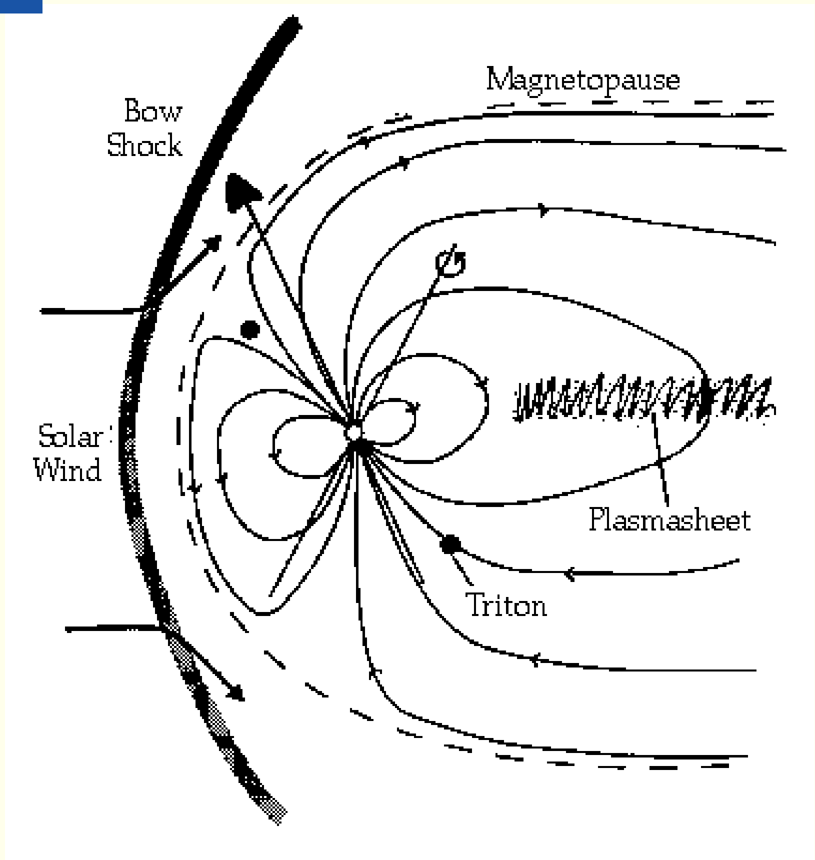
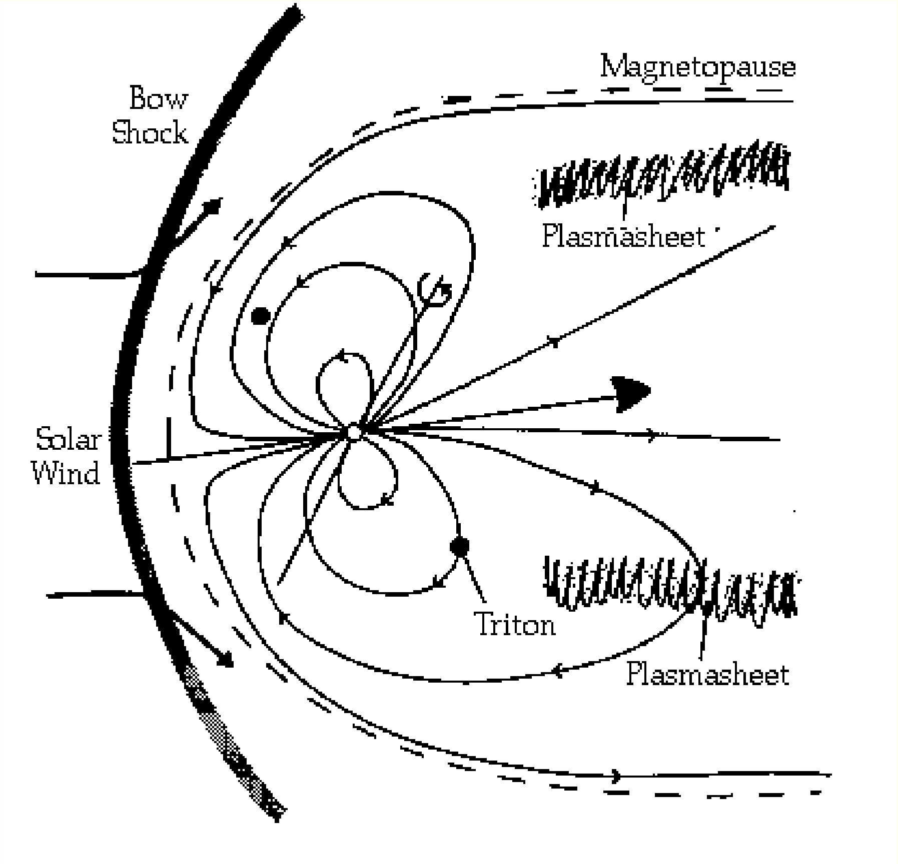


Photo from Hubble Space Telescope

- $r_N = 3{,}93\ r_E$
- $m_N = 17{,}2\ m_E$
- distance from sun : 30.1 AU
- gas giant with very dense atmosphere containing mainly methane
- 11 moons + ring system

Neptune

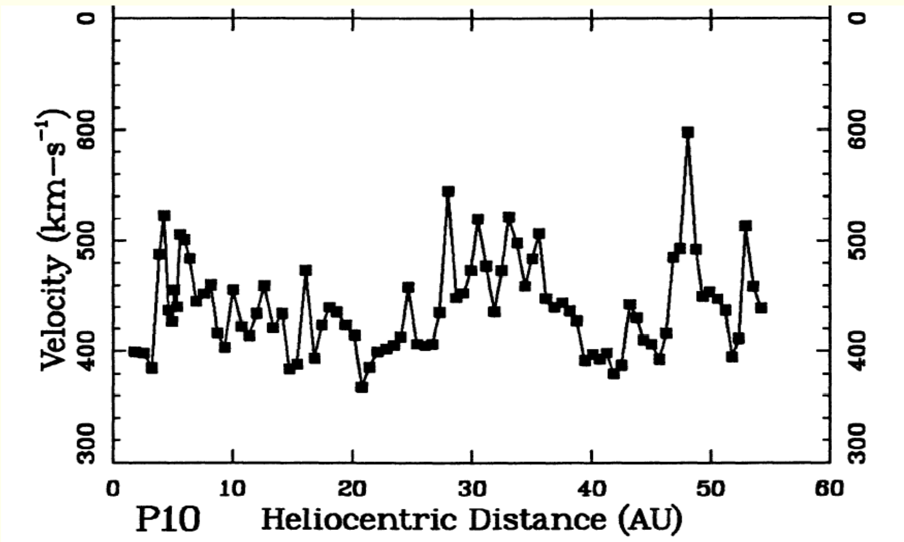




- magnetic field axis makes an angle of around 43° with rotational axis
- this gives enourmous daily variations of the structure of the magnetosphere also for Neptune

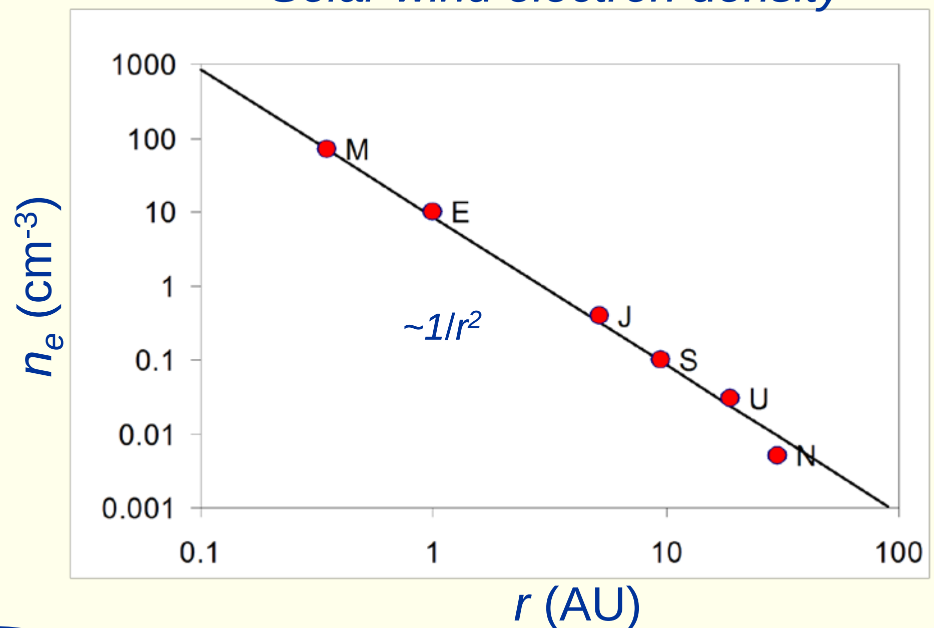
Comparative magnetospheres
Solar wind properties

Solar wind velocity

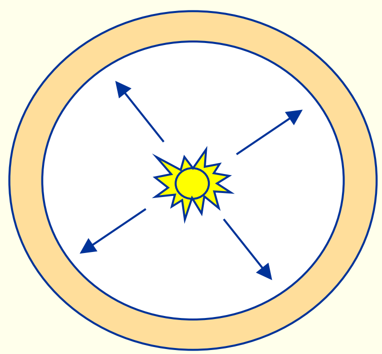


Pioneer 10, measurements
[Grazin et al., 1994]

Solar wind electron density



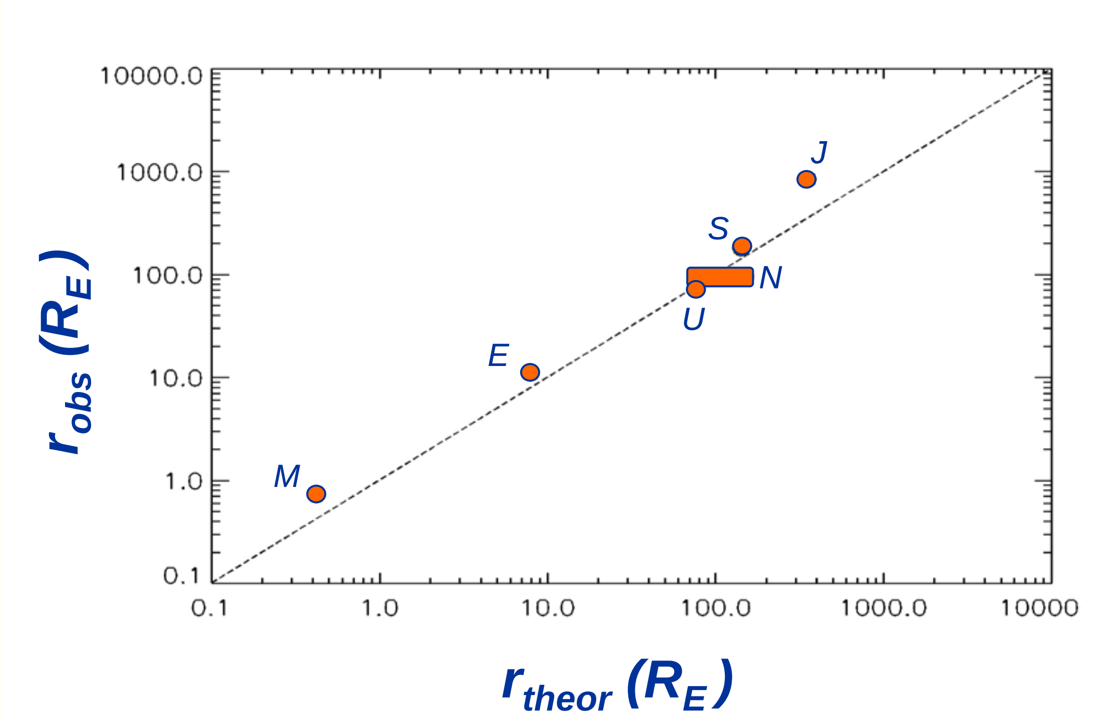
[Blanc et al., 2005]



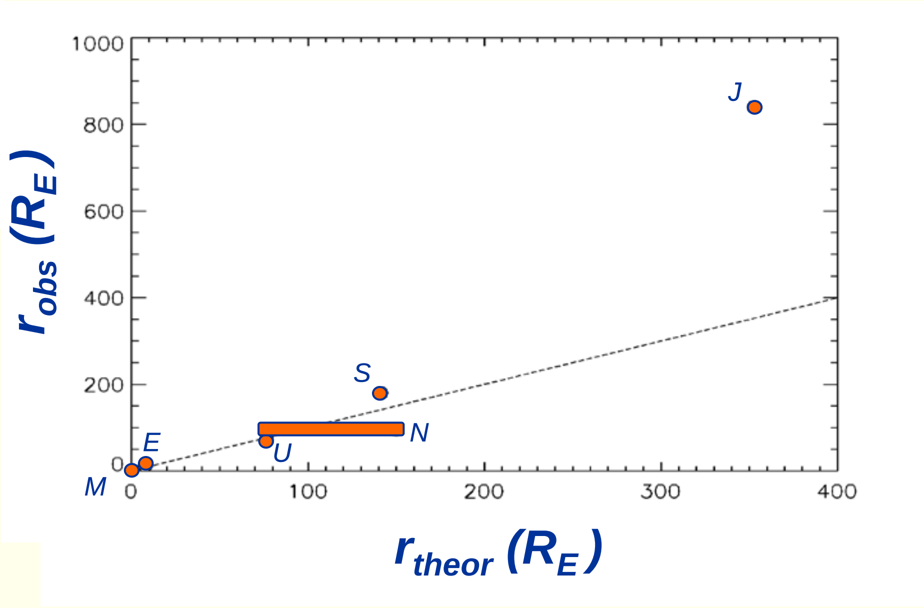
$$dV = 4\pi r^2 dr$$

Comparative magnetospheres

Observed vs. theoretical standoff-distance



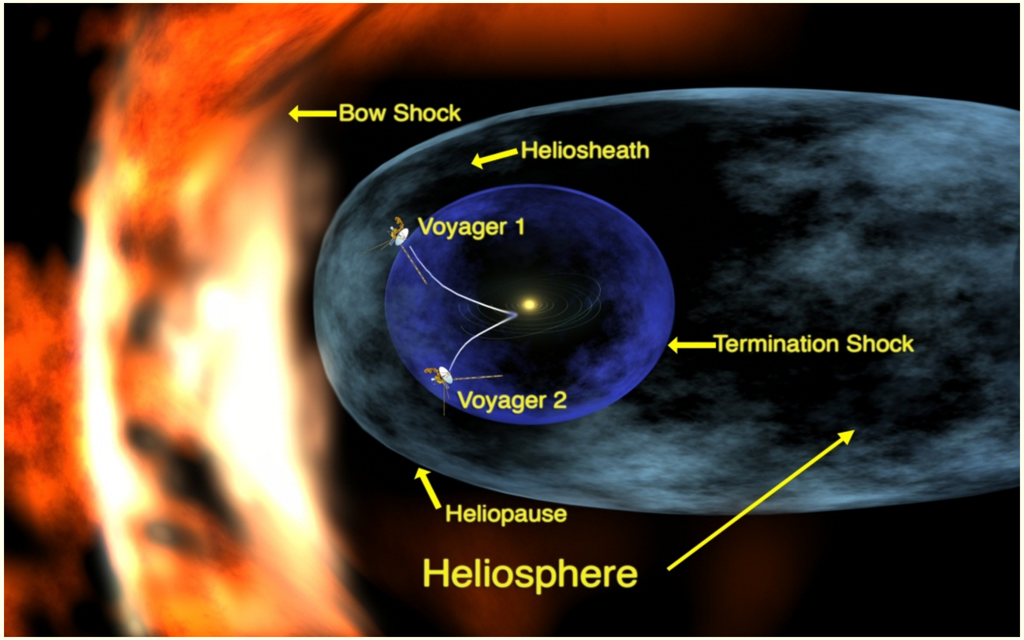
$$r_{theor} = \left(\frac{\mu_0 a}{4\pi}\right)^{1/3} \left(2\mu_0 \rho_{SW} v_{SW}^2\right)^{-1/6}$$

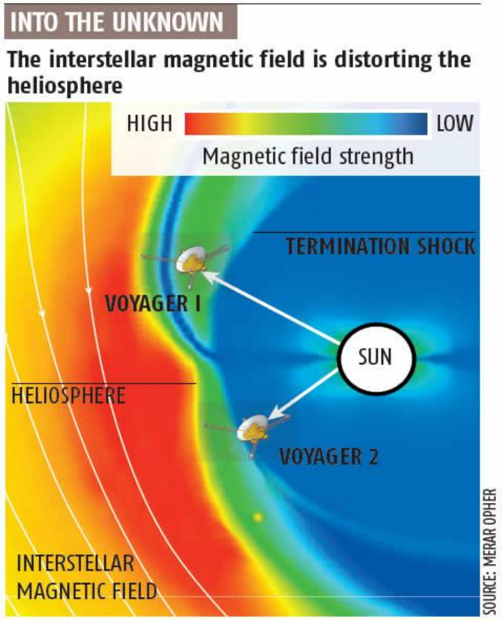


- *Model reasonably valid over three orders of magnitude*
- *Size of Jupiter's (and maybe Saturn's) magnetosphere underestimated*

Other other magnetospheres
Heliosphere

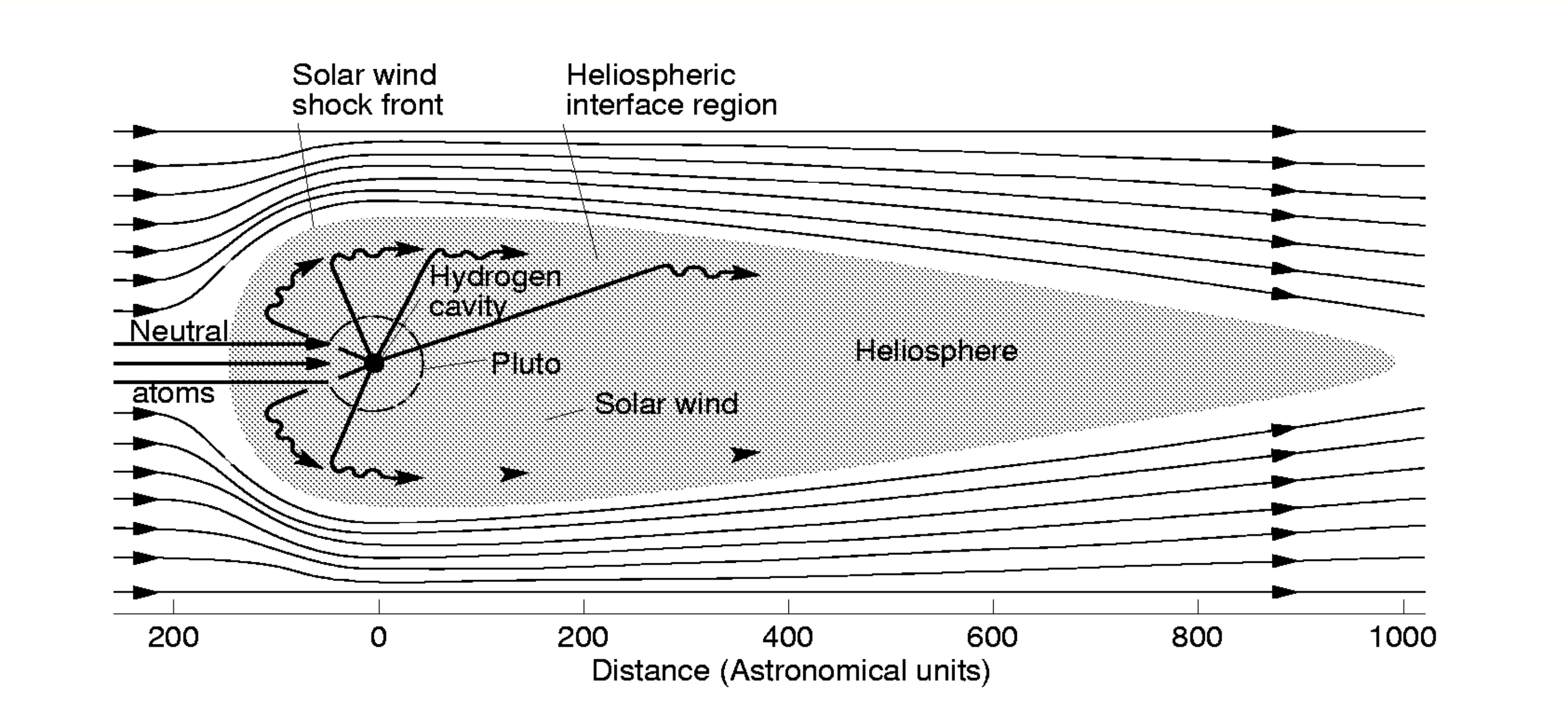
Interstellar wind





[Opher, 2007]

Heliosphere



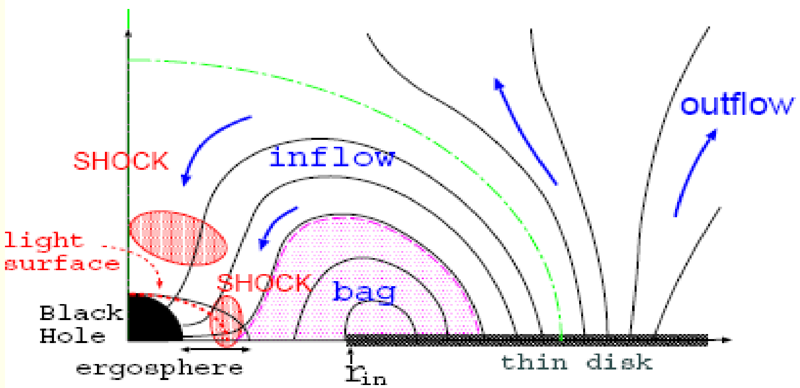
- Reaches approximately 100 AU into space (=1.5×10^{13} m)
- Voyager sonds are approaching/encountering the heliopause right now

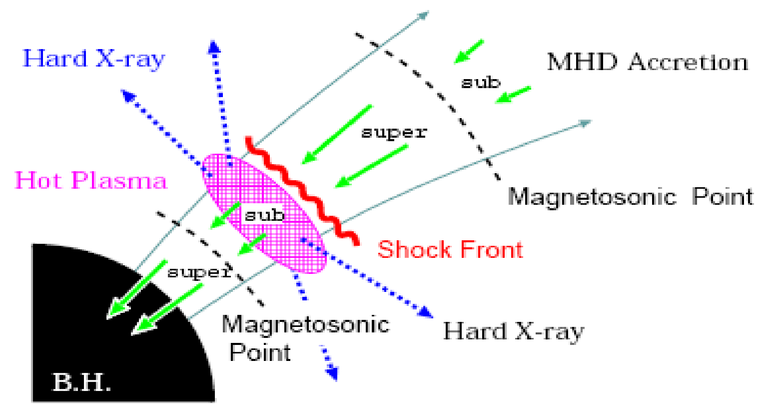
Other other magnetospheres

Exotic magnetospheres

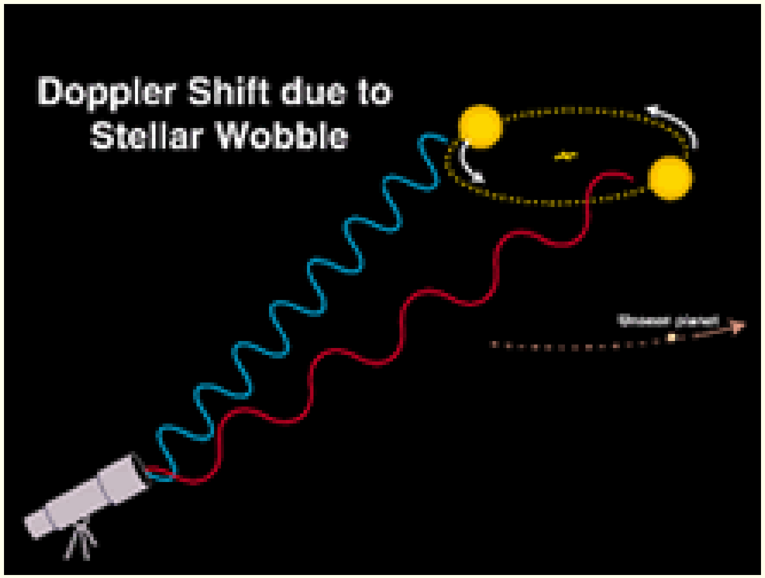
"Black Hole Magnetosphere" =

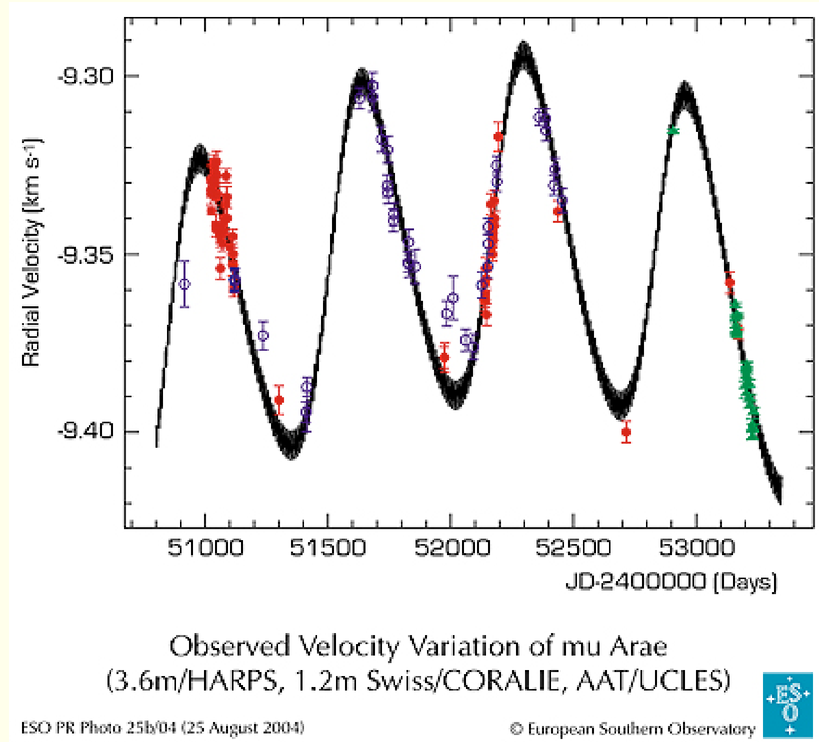
Loop (bag) + Disk–BH (inflow) + BH (Disk) − r_∞ (outflow)





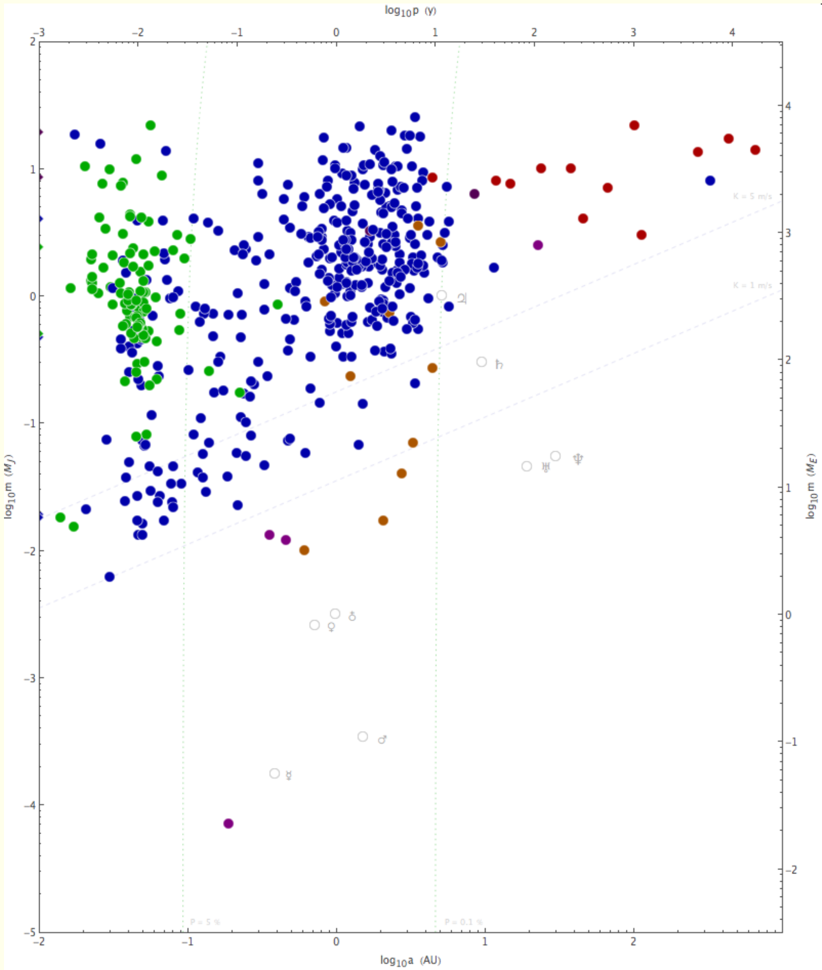
Exoplanets





Mostly detected by doppler shifts in starlight, or by dimming of the emitted starlight during transit of planet.

Magnetospheres of exoplanets

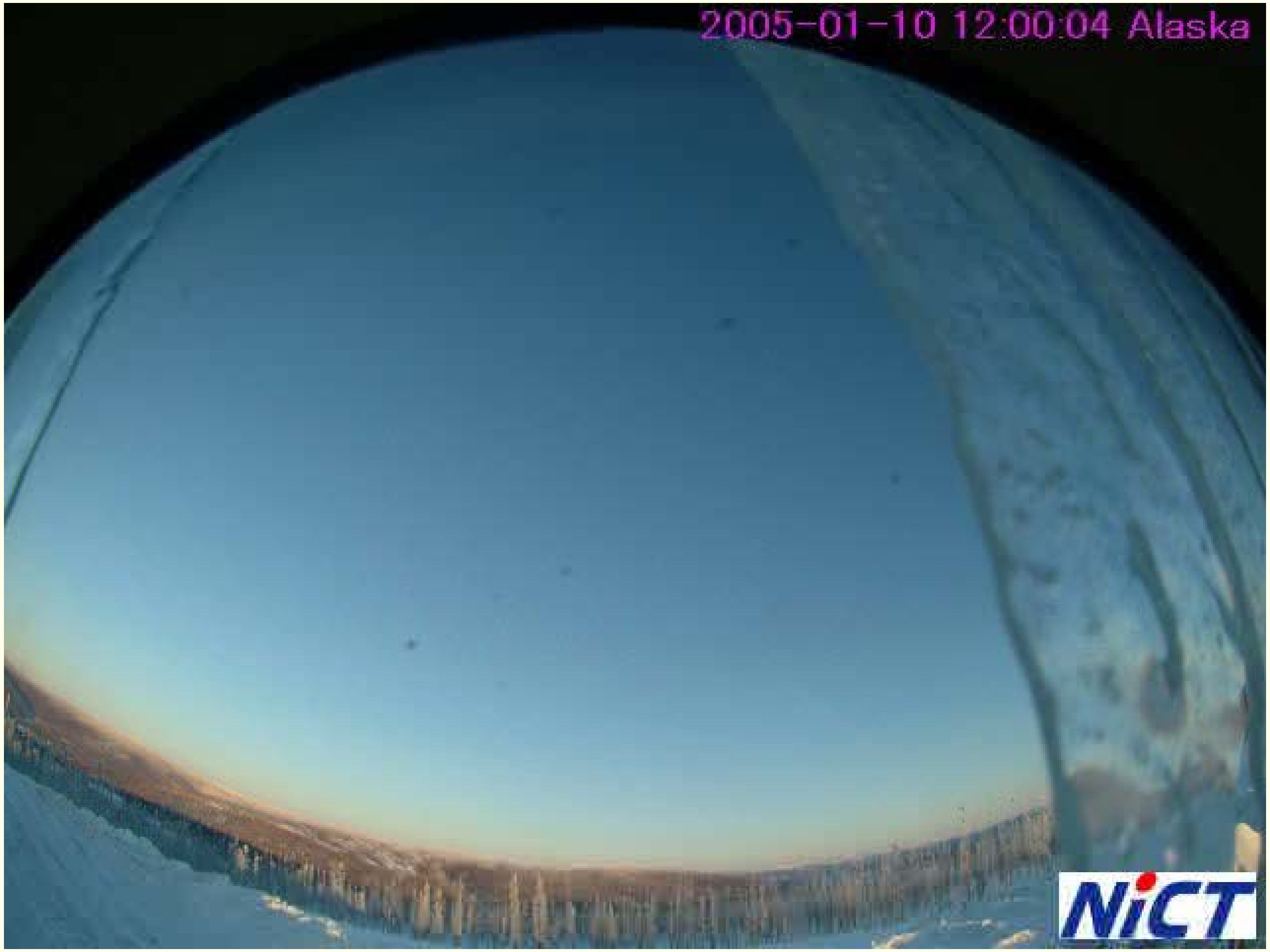


log *m*

log *a*

- 728 exoplanets found up to January 23, 2012
- Nothing is known about their atmospheres, ionospheres or magnetospheres
- Can possibly be detected by
  - radio emissions from auroral activity or from radiation belts
  - spectroscopy when the magnetosphere is in front of its sun

The aurora

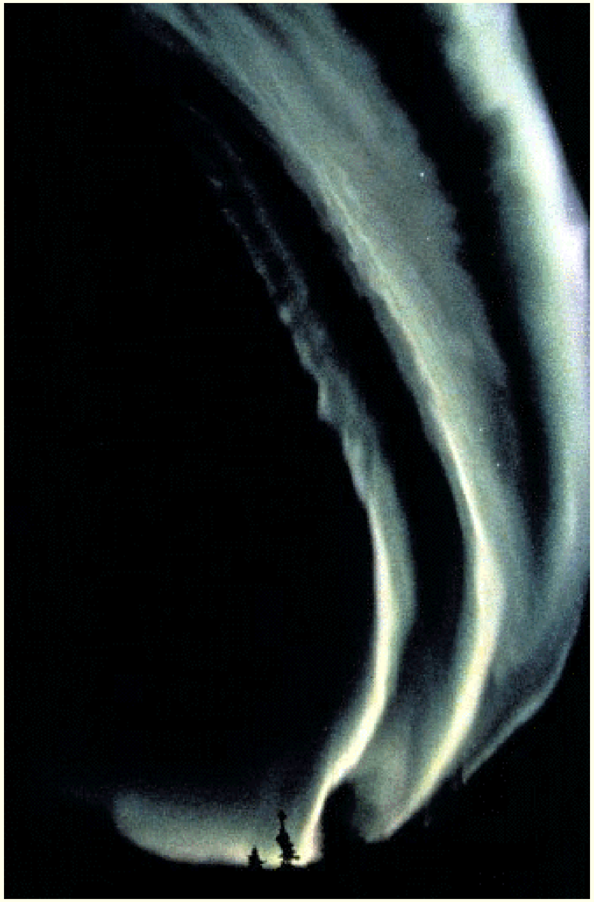


The aurora

The aurora

The aurora

Homogenous auroral arcs

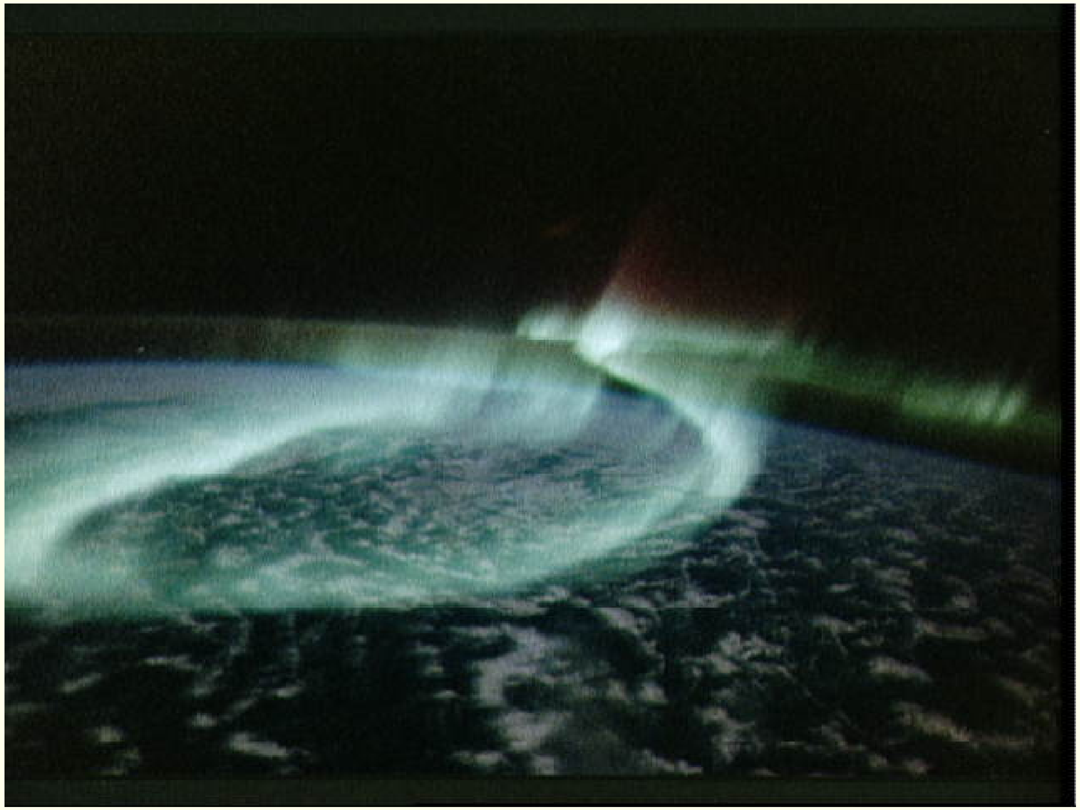


Rays, curtains

Rays are formed in the direction of the local magnetic field.

Drapes develop from homogenous arcs, often when they increase in intensity.

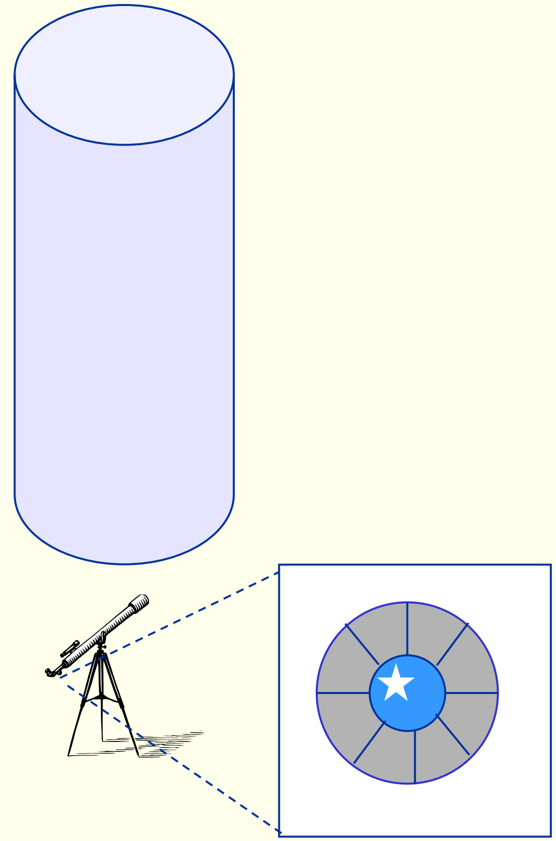
Auroral spirals



Develop when arcs become unstable

Auroral corona

Geometric effect of perspective when you look towards magnetic zenith. Compare the figure.



Aurora - altitude

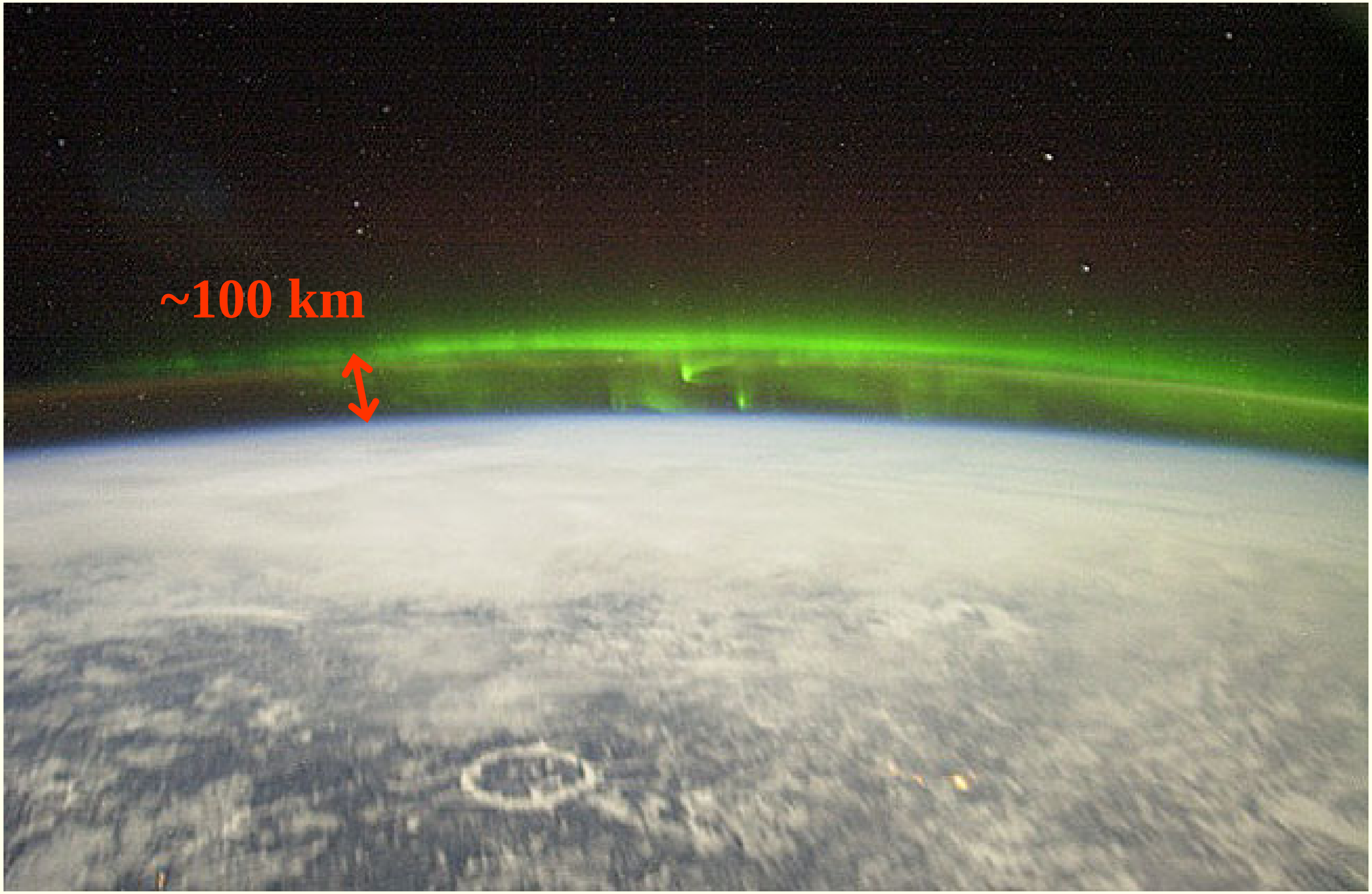


Foto from International Space Station

Early notions

Woodcut from Böhmen 1570.

Anders Celsius documented that compass needles where strongly affected during auroral activity in 1733.

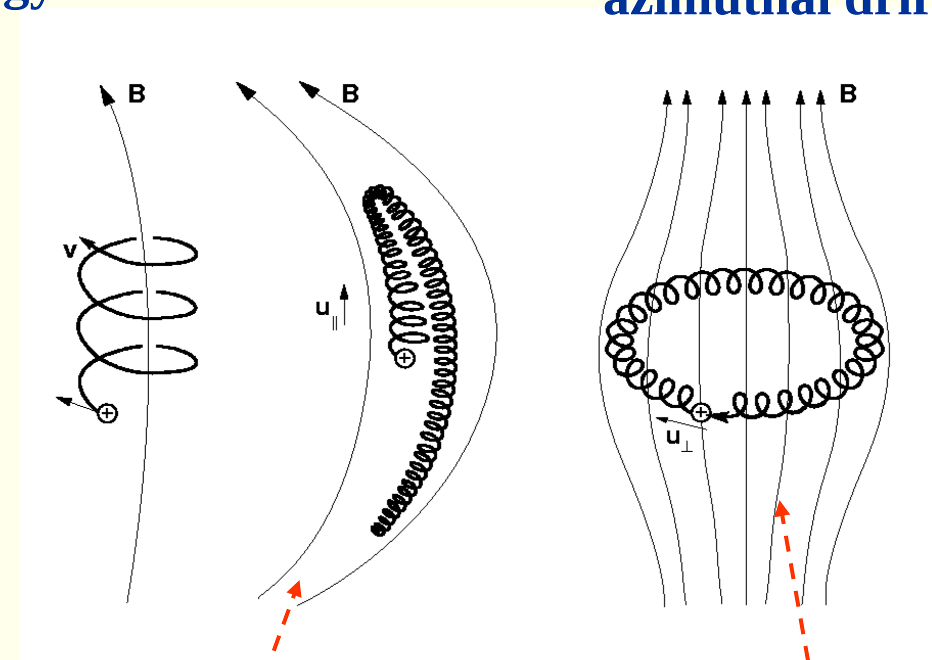
What causes the aurora?

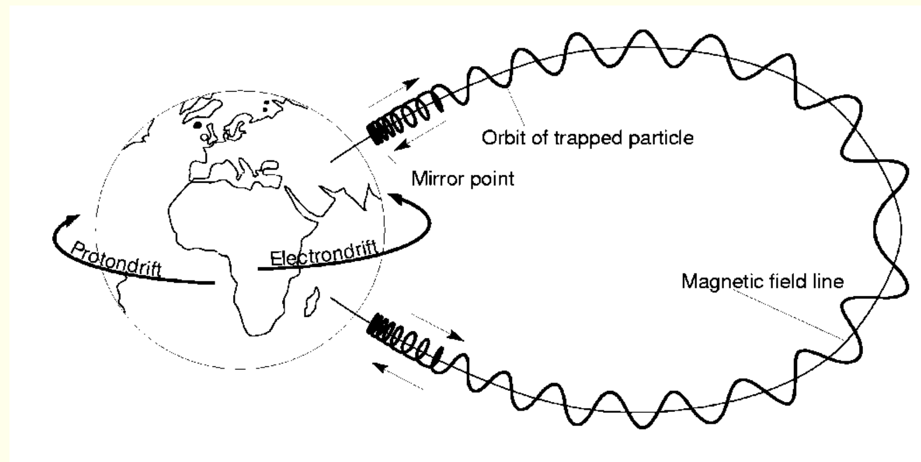
Particle motion in geomagnetic field

longitudinal oscillation

gyration

azimuthal drift

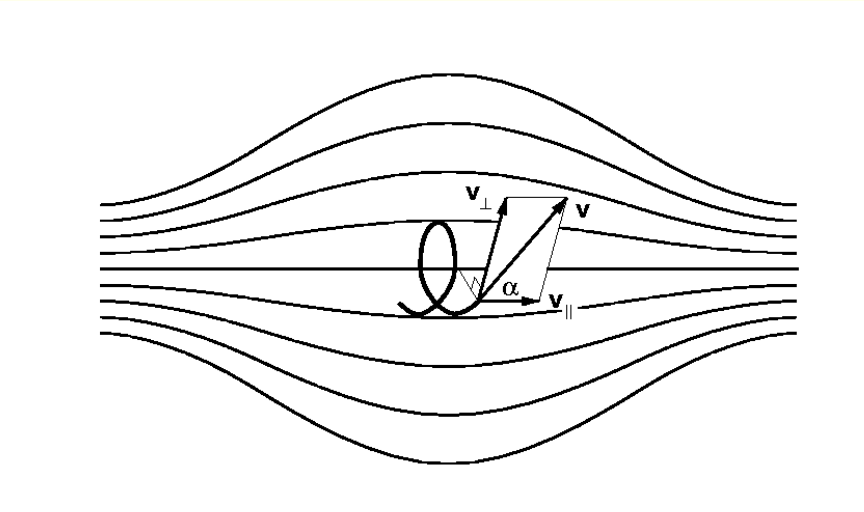




Magnetic mirror

grad B drift

Magnetic mirror



$mv^2/2$ constant (energy conservation) ➡

$$\frac{\sin^2\alpha}{B} = konst$$

particle turns when α = 90° ➡

$$B_{turn} = B/\sin^2\alpha$$

If maximal B-field is B_{max} a particle with pitch angle α can only be turned around if

$$B_{turn} = B/\sin^2\alpha \leq B_{max}$$ ➡

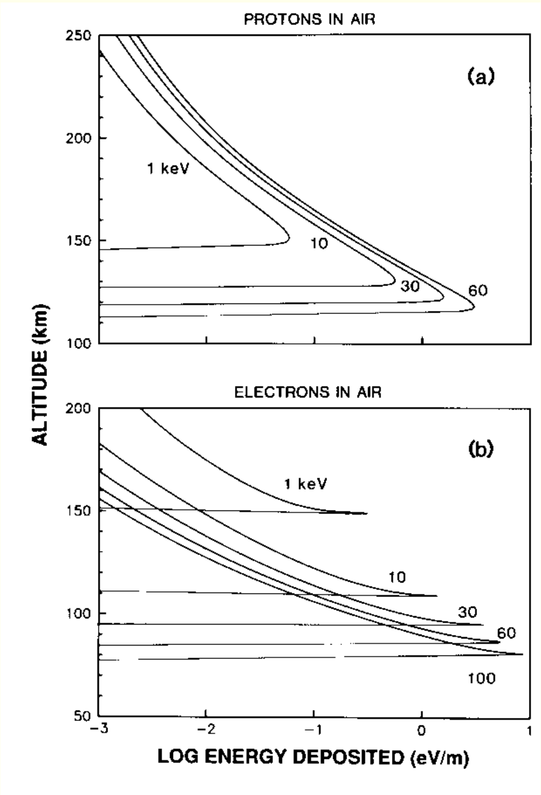
$$\alpha > \alpha_{lc} = \arcsin\sqrt{B/B_{max}}$$

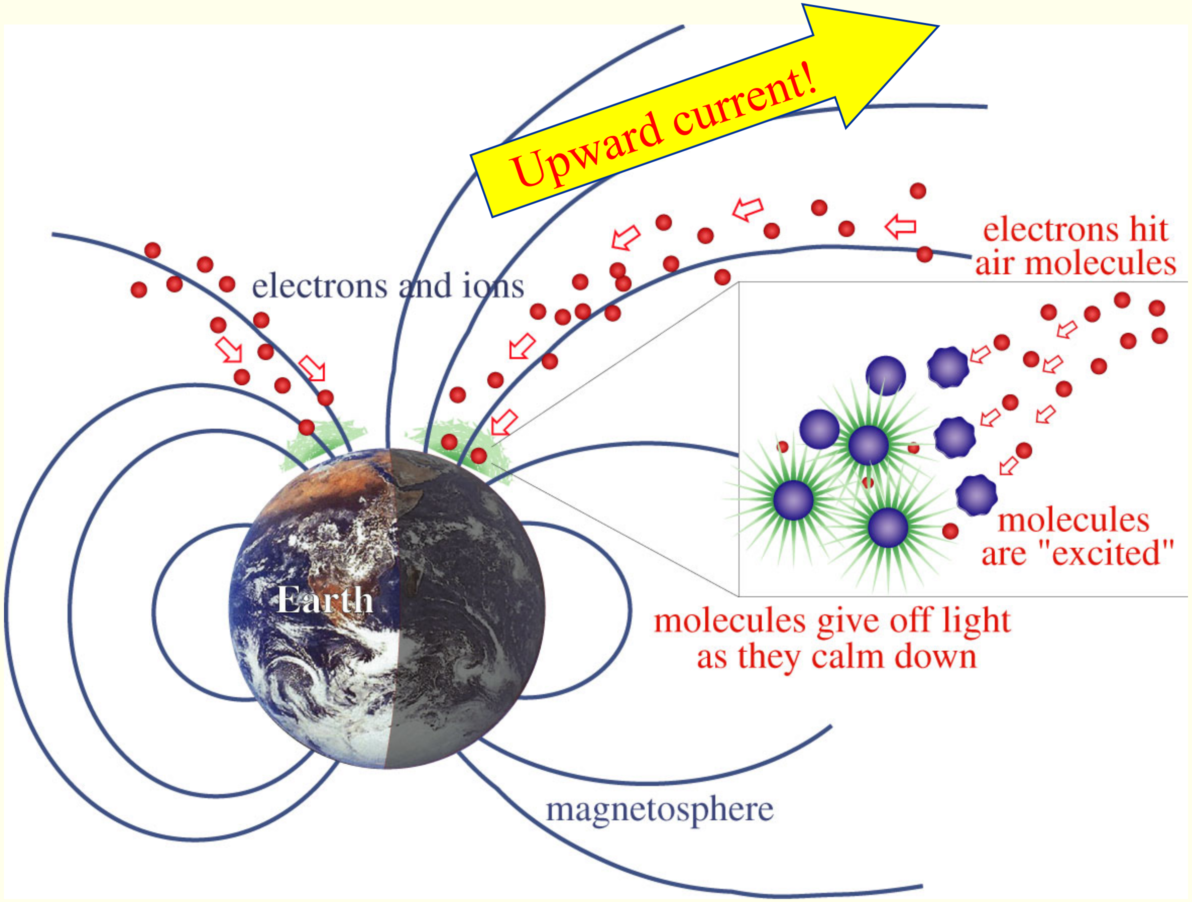
The magnetic moment μ is an ***adiabatic invariant***.

$$\mu = \frac{mv_\perp^2}{2B} = \frac{mv^2\sin^2\alpha}{2B}$$

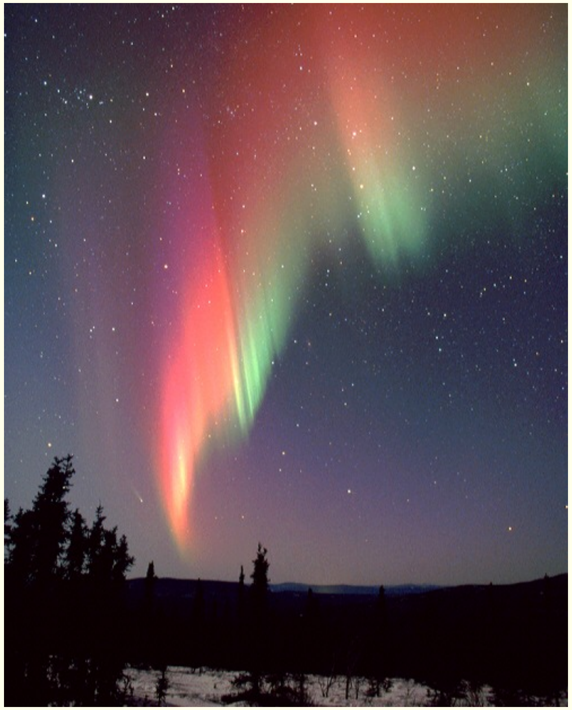
Particles in ***loss cone*** : $\alpha < \alpha_{lc}$

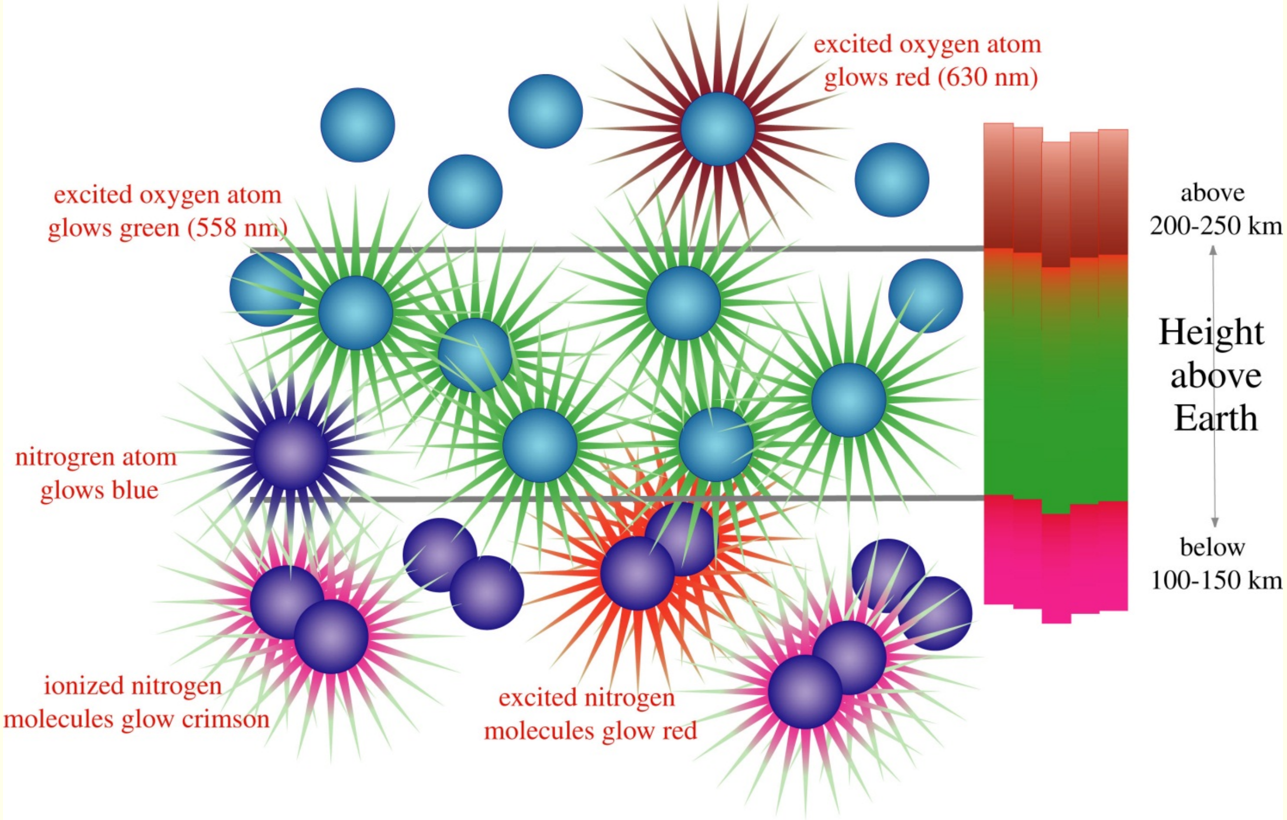
Collisions - emissions



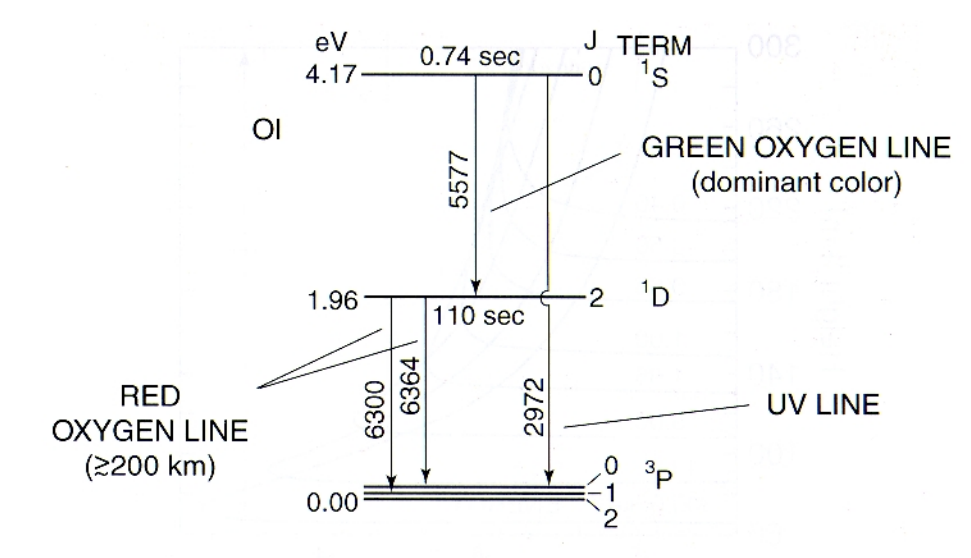


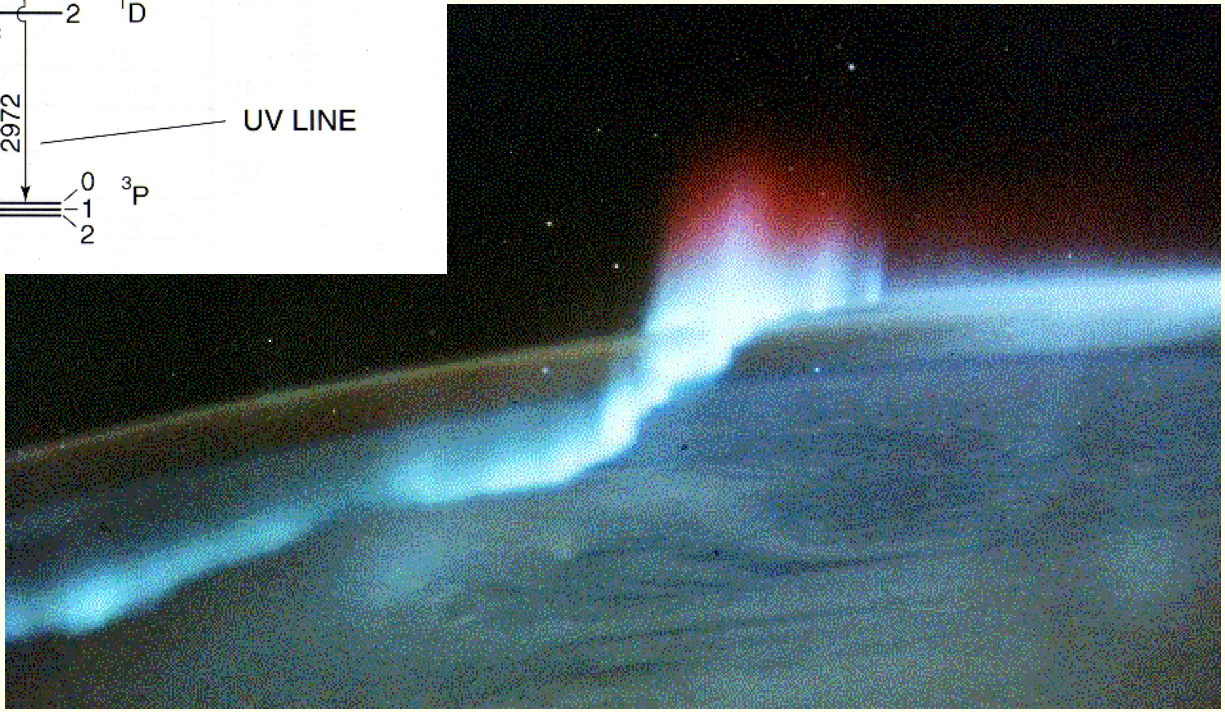
Emissions



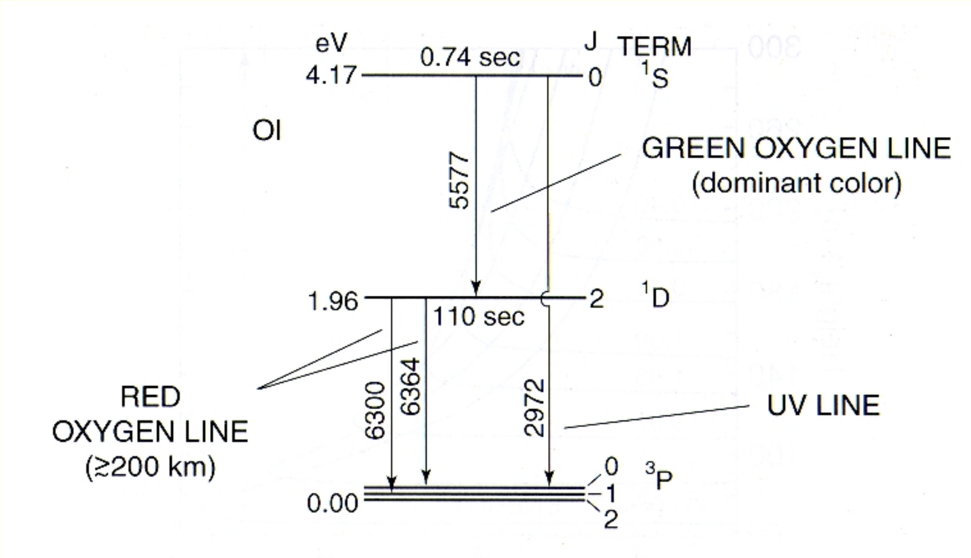

excited oxygen atom
glows red (630 nm)
excited oxygen atom
glows green (558 nm)
nitrogren atom
glows blue
ionized nitrogen
molecules glow crimson
excited nitrogen
molecules glow red
above
200-250 km
Height
above
Earth
below
100-150 km

Oxygen emissions

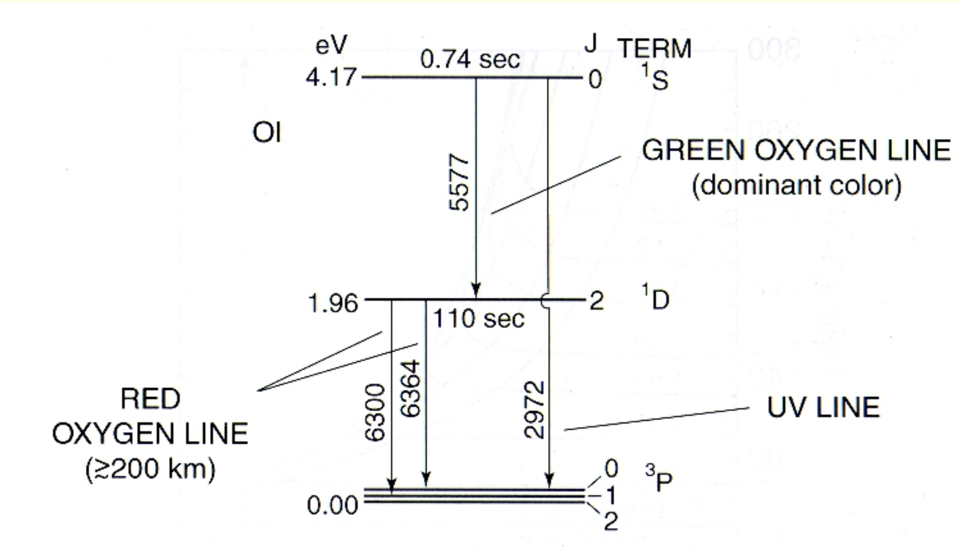


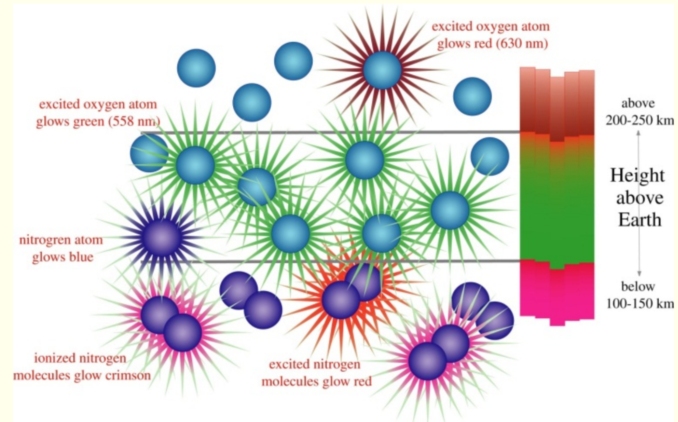


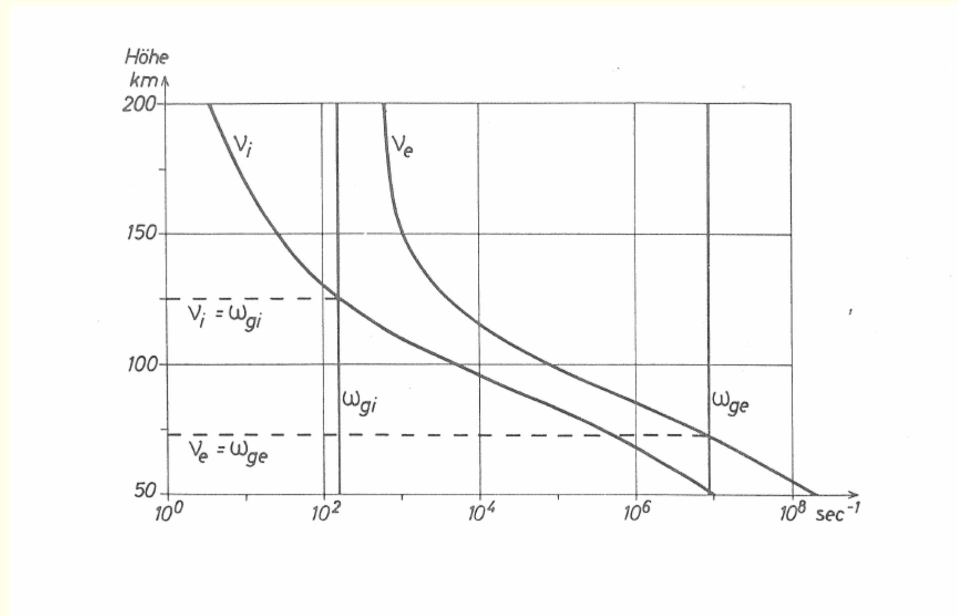
Why is there no red emissions at lower altitude?



Oxygen emissions







The red emission line is suppressed by collisions at lower altitudes due the its long transition time. (When an excited atom collides with another atom, is is de-excited without any emission.)

Larger scales

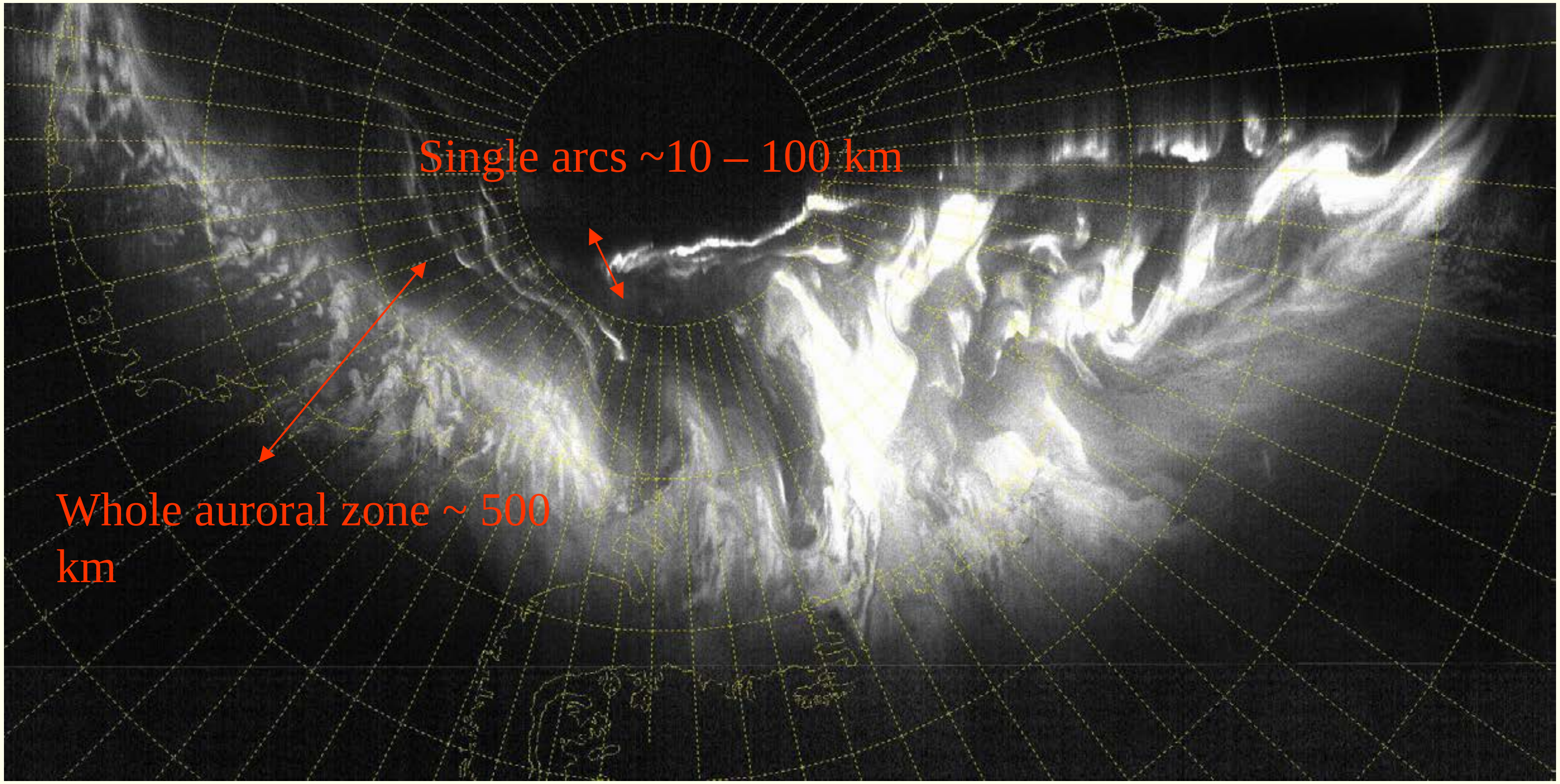
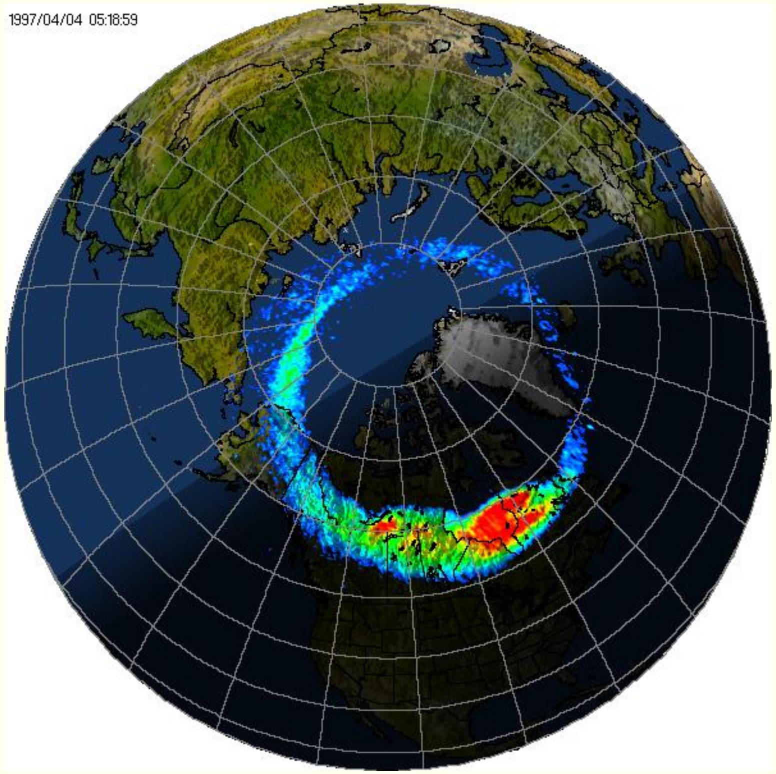
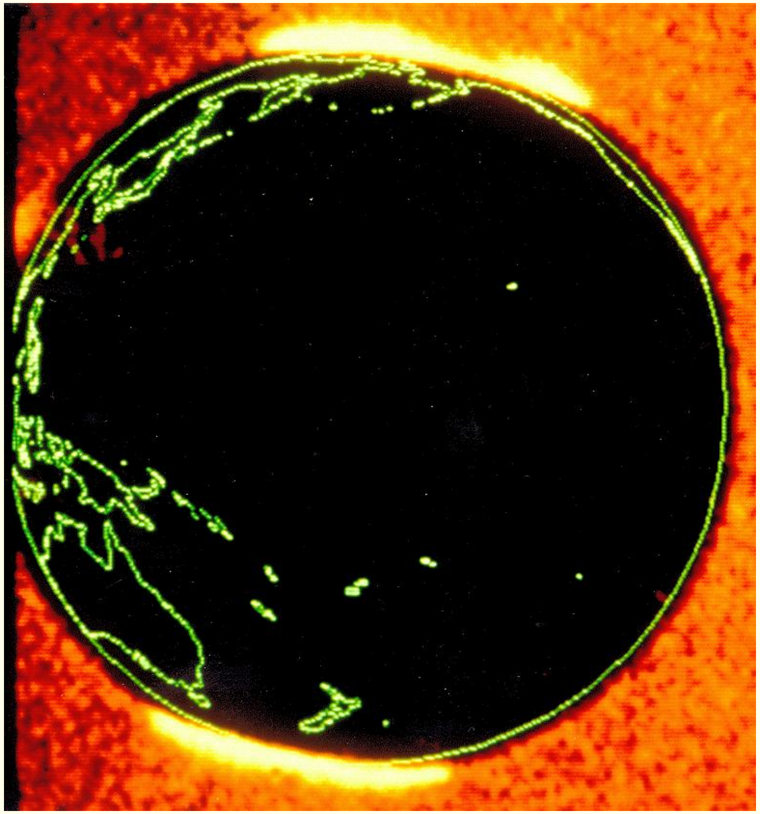


Foto från DMSP-satelliten

Auroral ovals



Polar

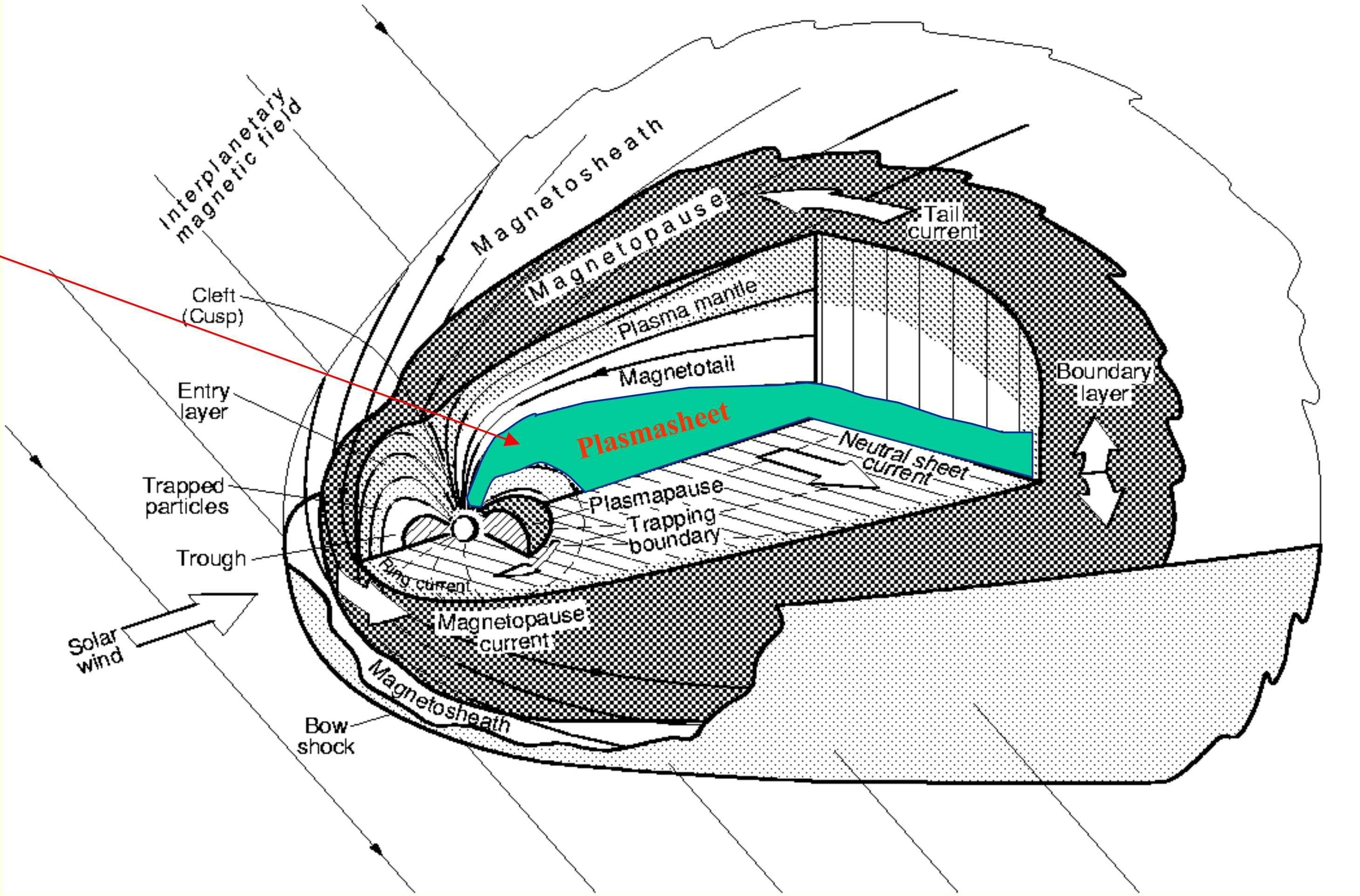


Dynamics Explorer

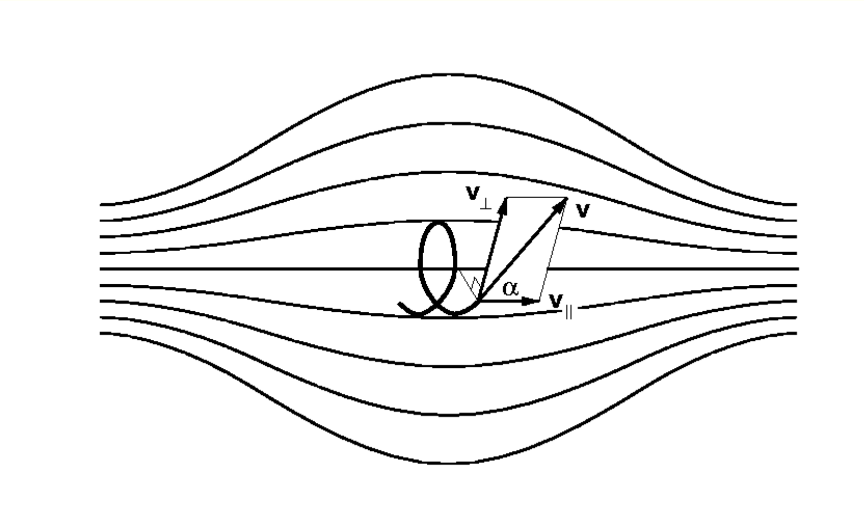
The auroral oval is the projection of the plasmasheet onto the atmosphere

Mystery!

The particles in the plasmasheet do not have high enough energy to create aurora visible to the eye.



Magnetic mirror



mv²/2 constant (energy conservation) →

$$\frac{\sin^2\alpha}{B} = konst$$

particle turns when α = 90° →

$$B_{turn} = B/\sin^2\alpha$$

If maximal B-field is B_{max} a particle with pitch angle α can only be turned around if

$$B_{turn} = B/\sin^2\alpha \le B_{max}$$ →

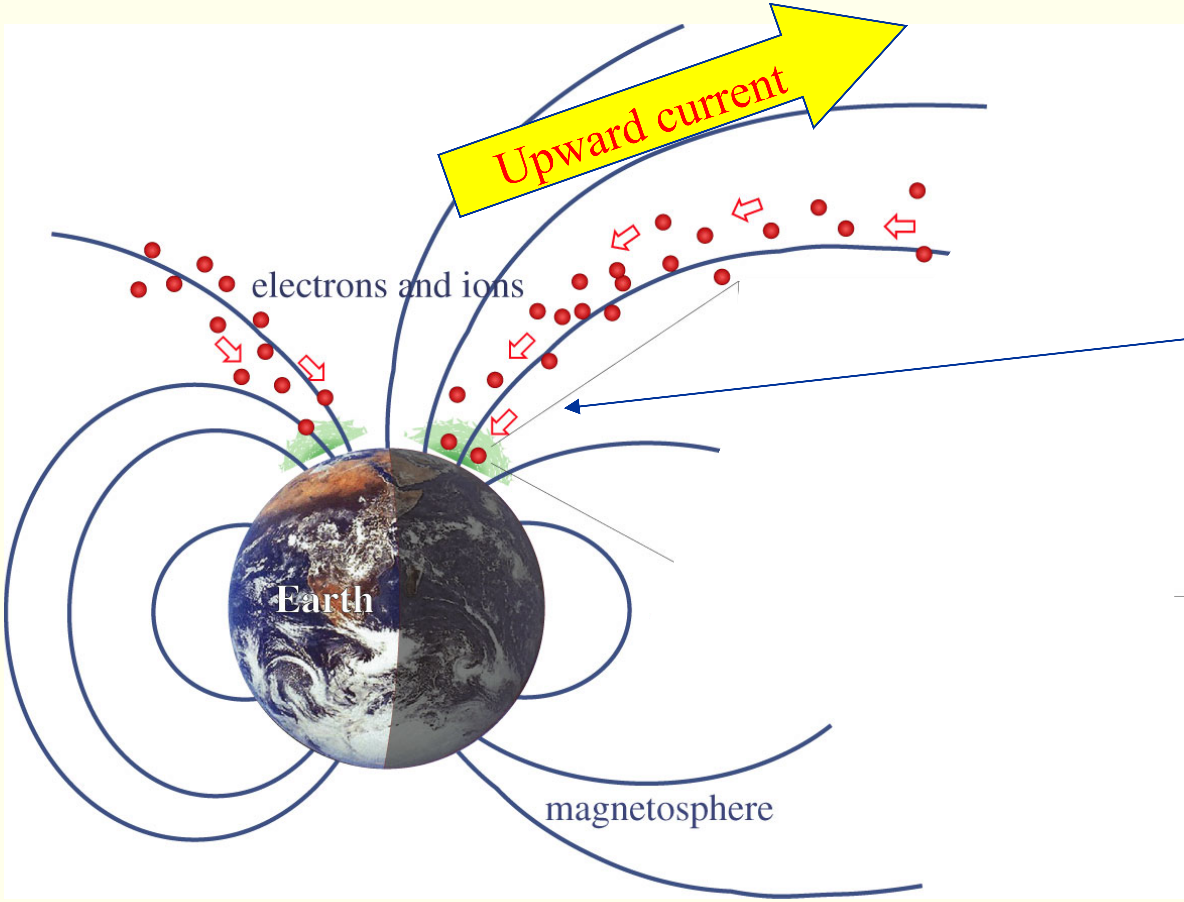
$$\alpha > \alpha_{fl} = \arcsin\sqrt{B/B_{max}}$$

The magnetic moment μ is an ***adiabatic invariant***.

$$\mu = \frac{mv_\perp^2}{2B} = \frac{mv^2\sin^2\alpha}{2B}$$

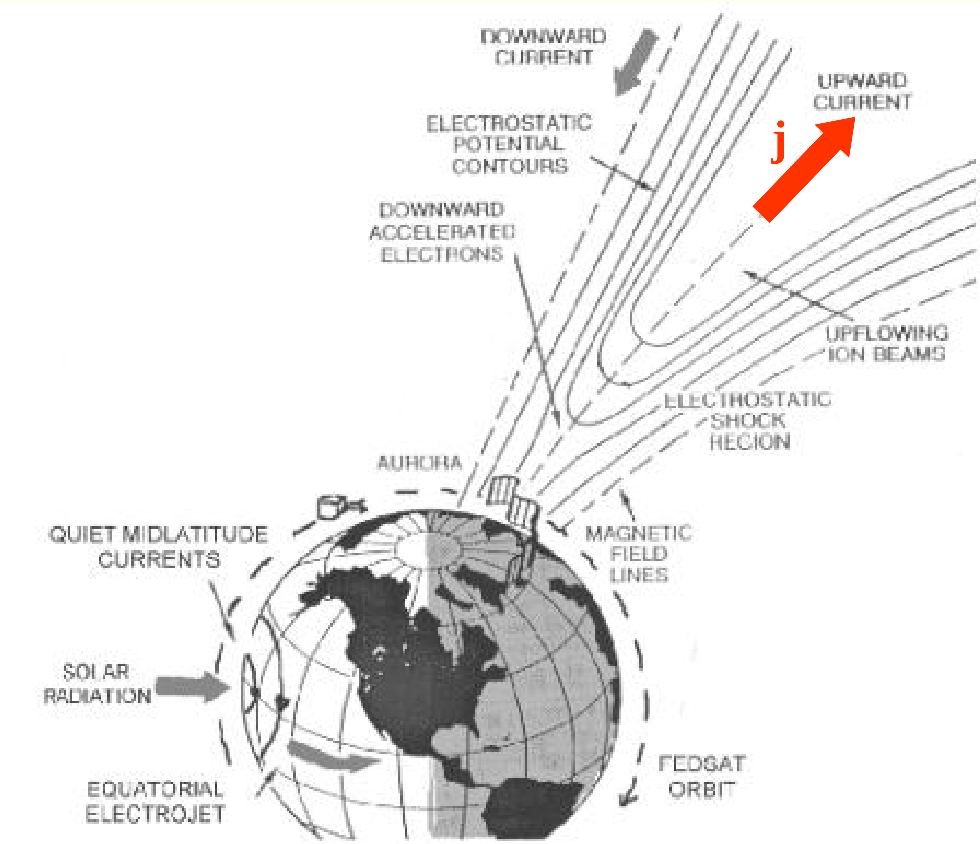
Particles in ***loss cone*** : $\alpha < \alpha_{fl}$

Why particle acceleration?



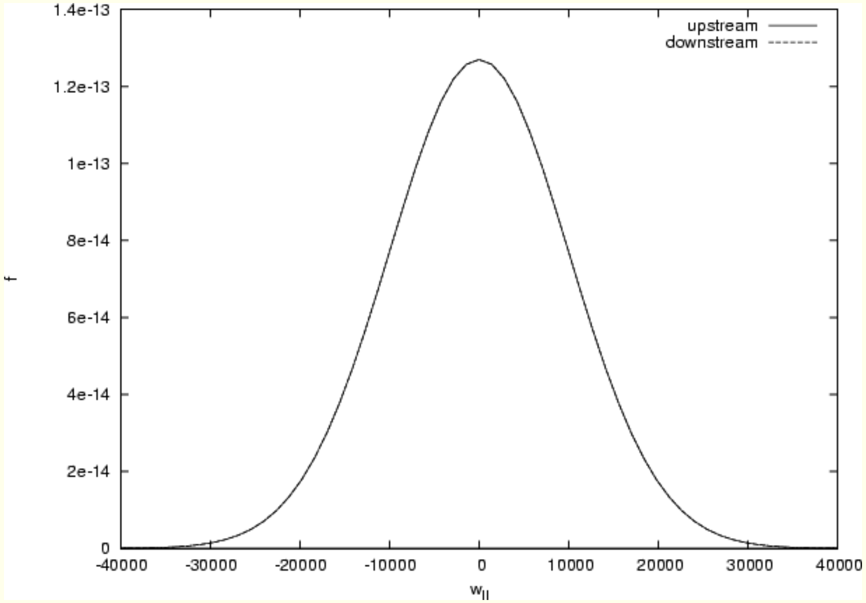
- The magnetosphere often seems to act as a current generator.
- The lower down you are on the field line, the more particles have been reflected by the magnetic mirror.
- At low altitudes there are not enough electrons to carry the current.

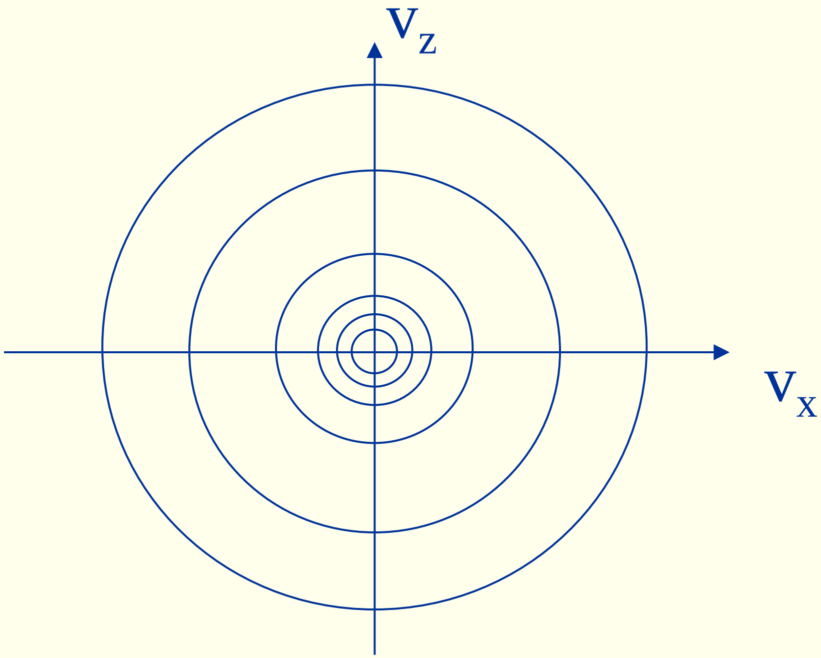
Why particle acceleration?



- Electrons are accelerated downwards by upward E-field.
- This increases the pitch-angle of the electrons, and more electrons can reach the ionosphere, where the current can be closed.

Distribution function

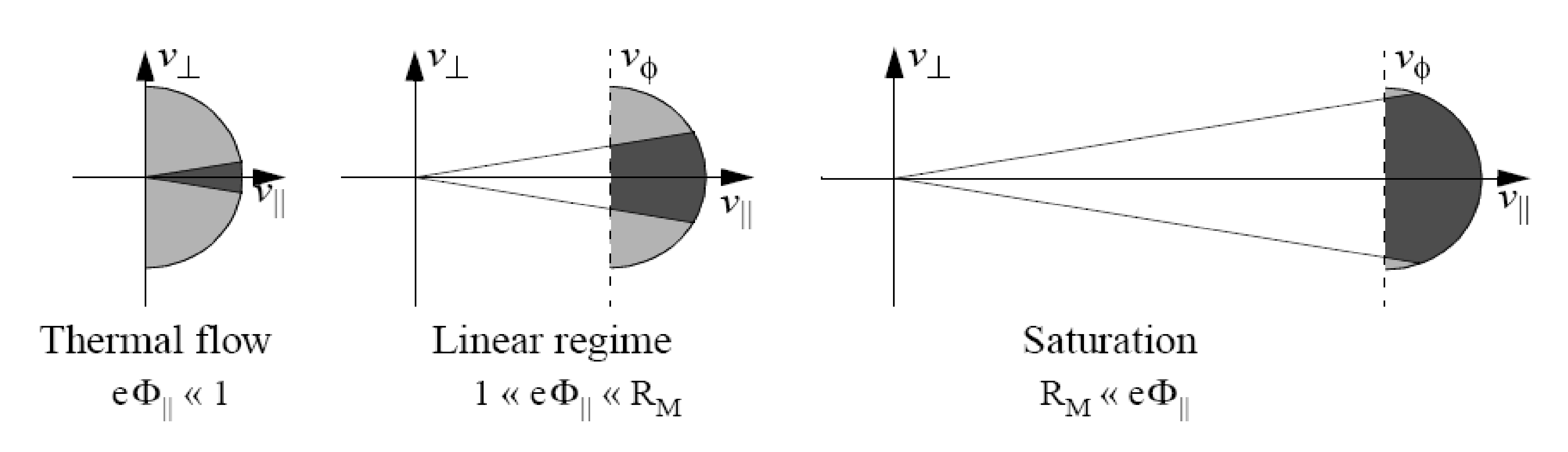




Example:
Maxwellian distribution

$$f = \frac{n}{\sqrt{(2\pi RT)^3}} \exp\left(-\frac{m(v_x^2 + v_y^2 + v_z^2)}{2kT}\right)$$

Why particle acceleration?



- Electrons are accelerated downwards by upward E-field.
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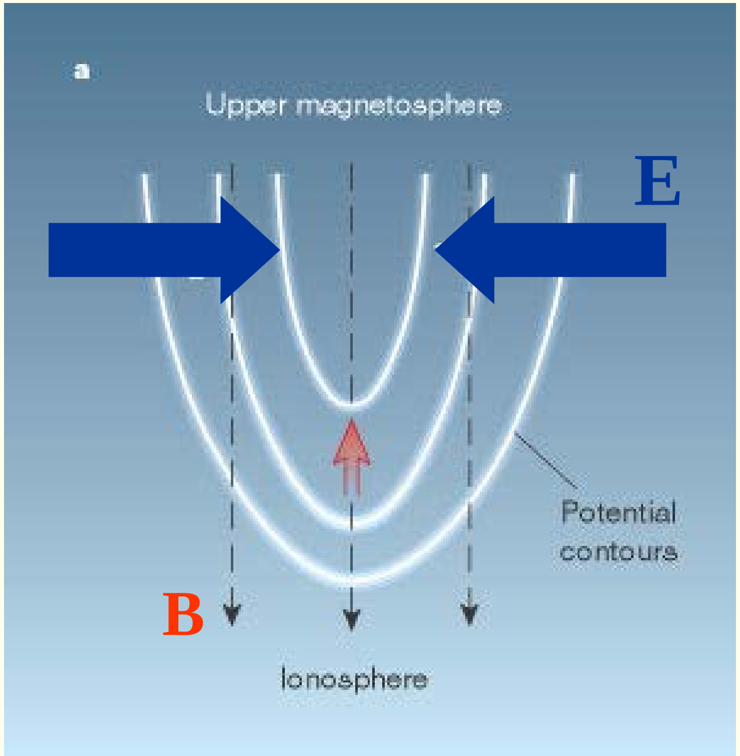
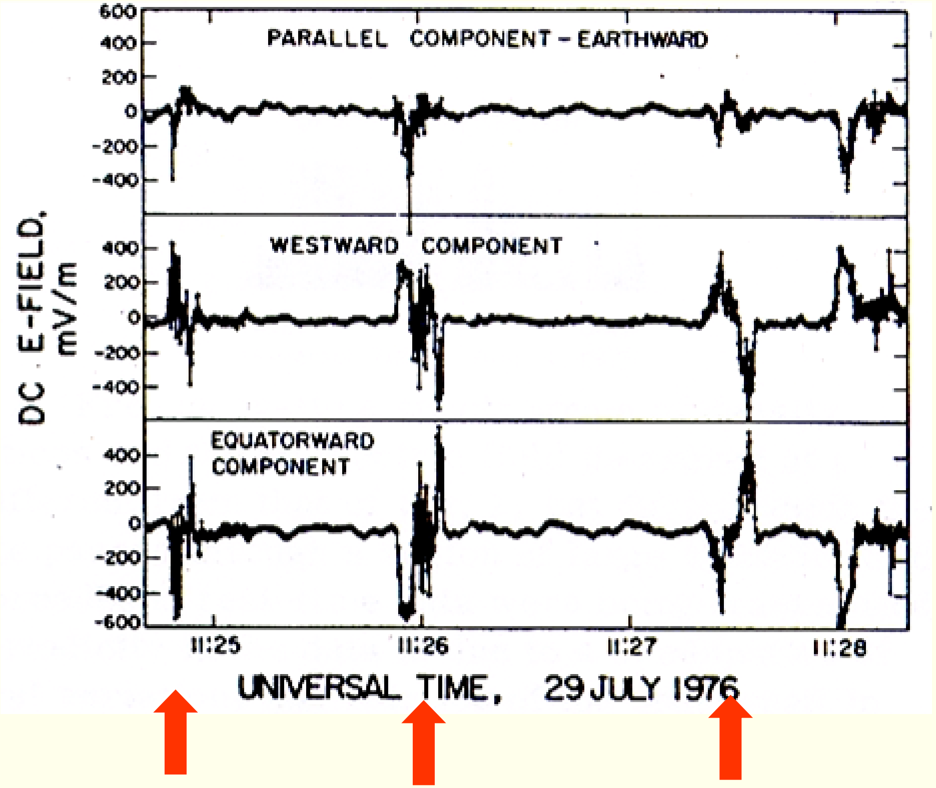
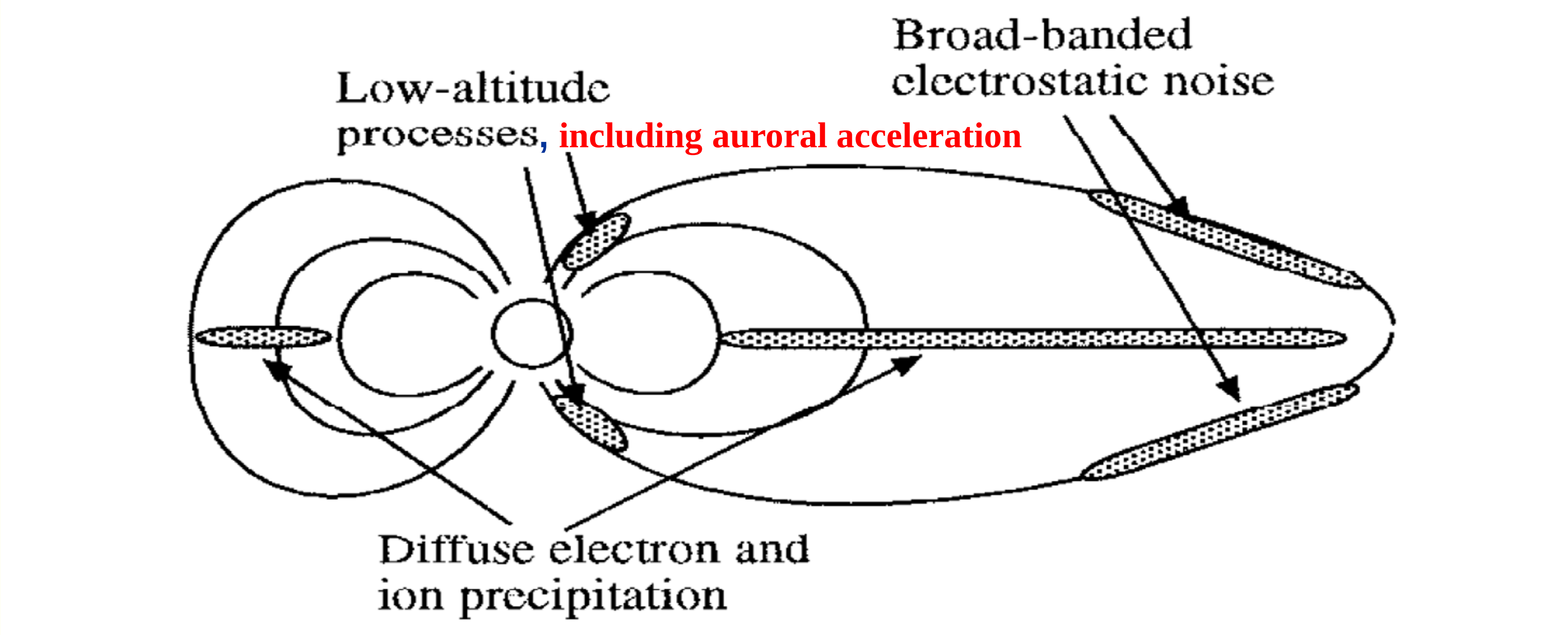
Satellite signatures of U potential





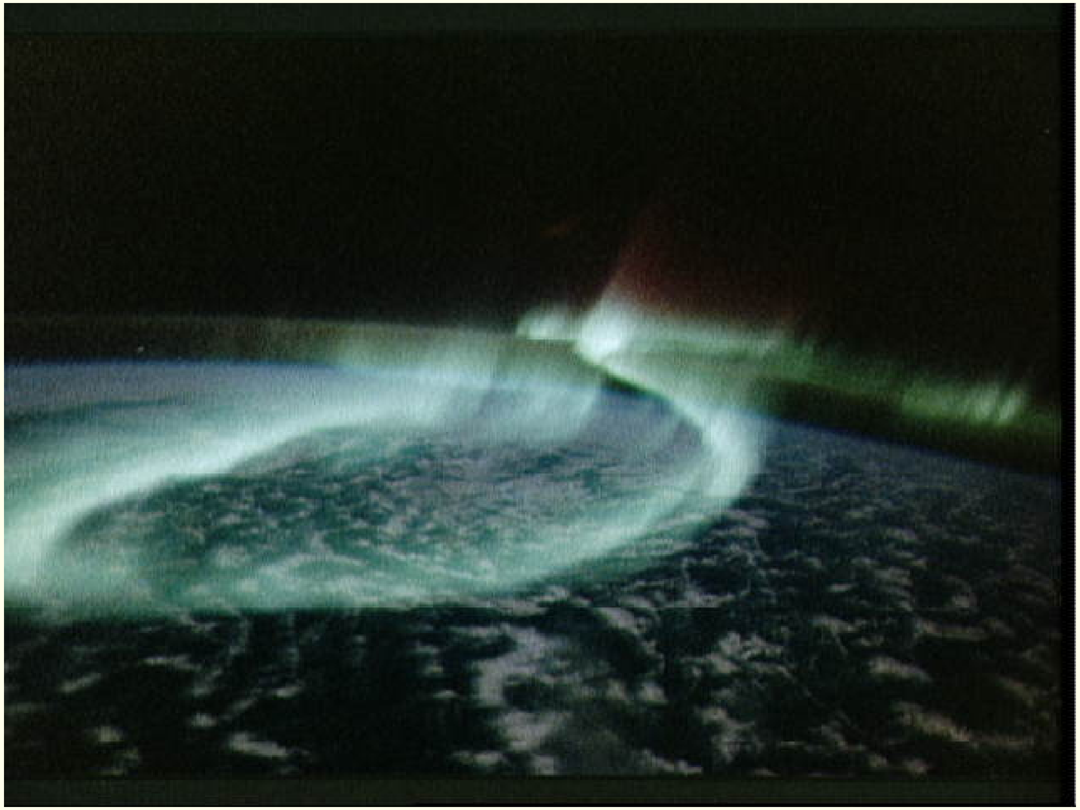
Measurements made by the ISEE satellite
(Mozer et al., 1977)

Acceleration regions



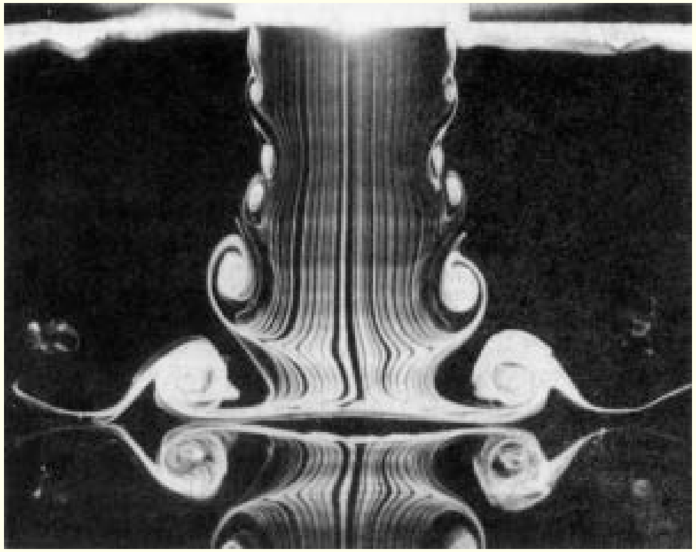
Auroral acceleration region typically situated at altitude of 1-3 R_E

Auroral spirals

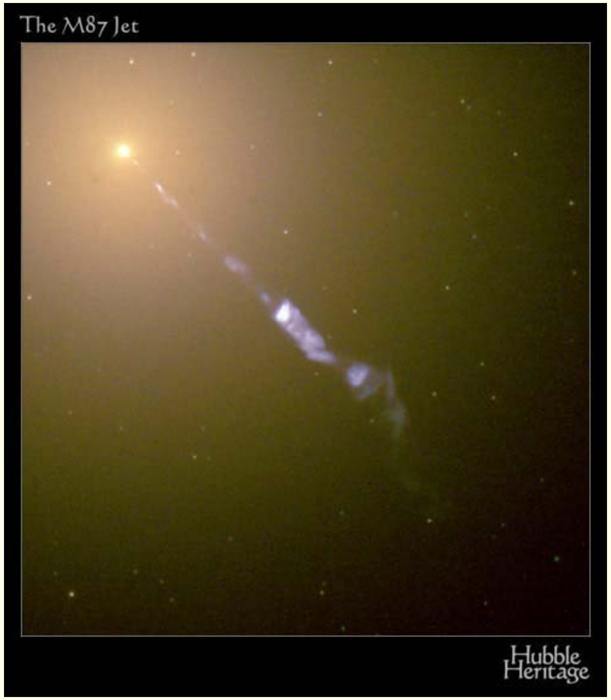


Develop when arcs become unstable

Kelvin-Helmholz-instability – a general phenomenon



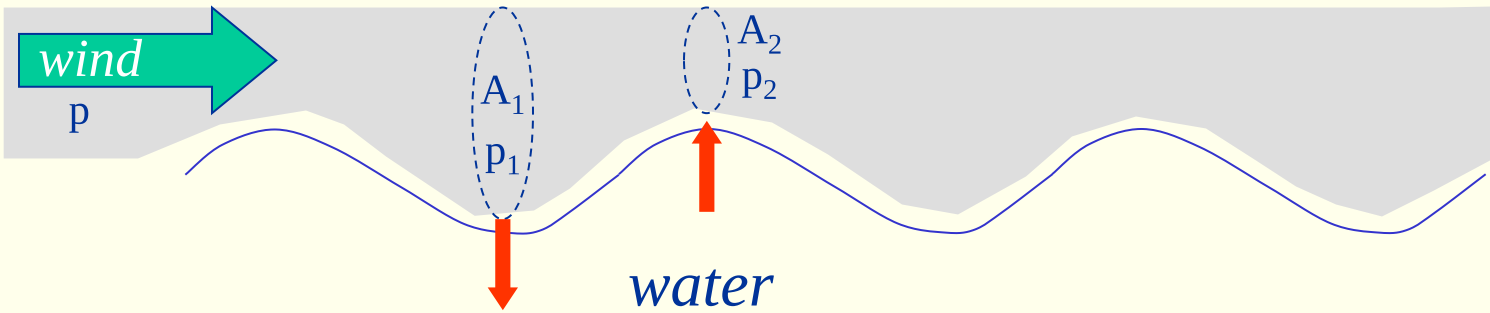
Aero- and fluid dynamics



Extragalactic jet (M87)

Cluds

Kelvin-Helmholz instability
Example: water waves



Continuity equation:

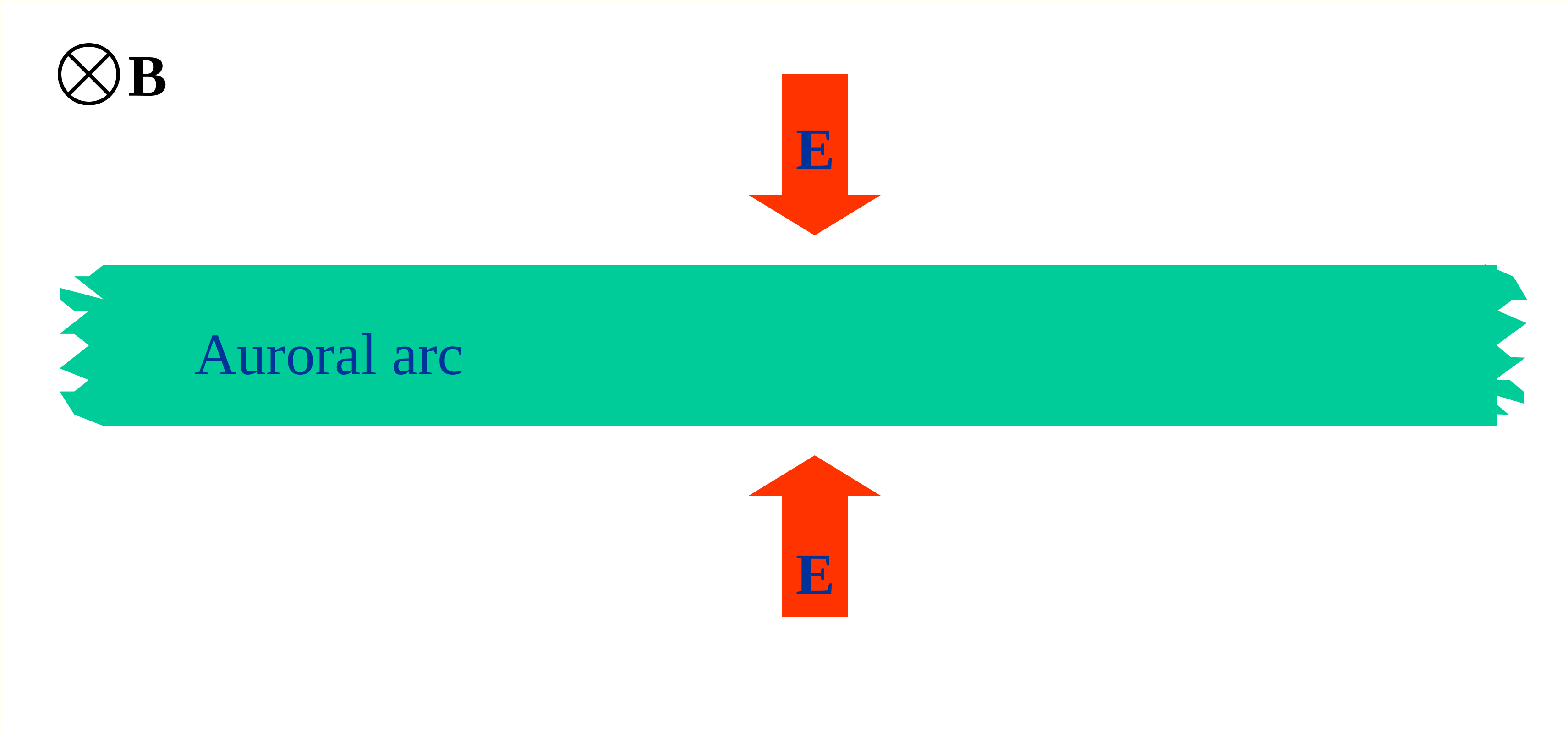
$$A_1 v_1 = A_2 v_2$$

Bernoulli's equation:

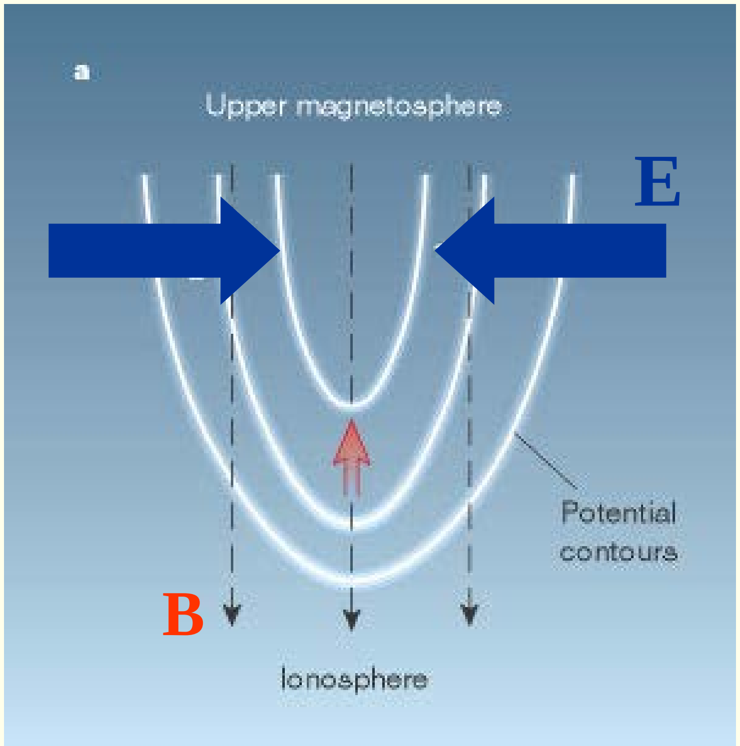
$$p_1 + \rho v_1^2 = p_2 + \rho v_2^2 = const.$$

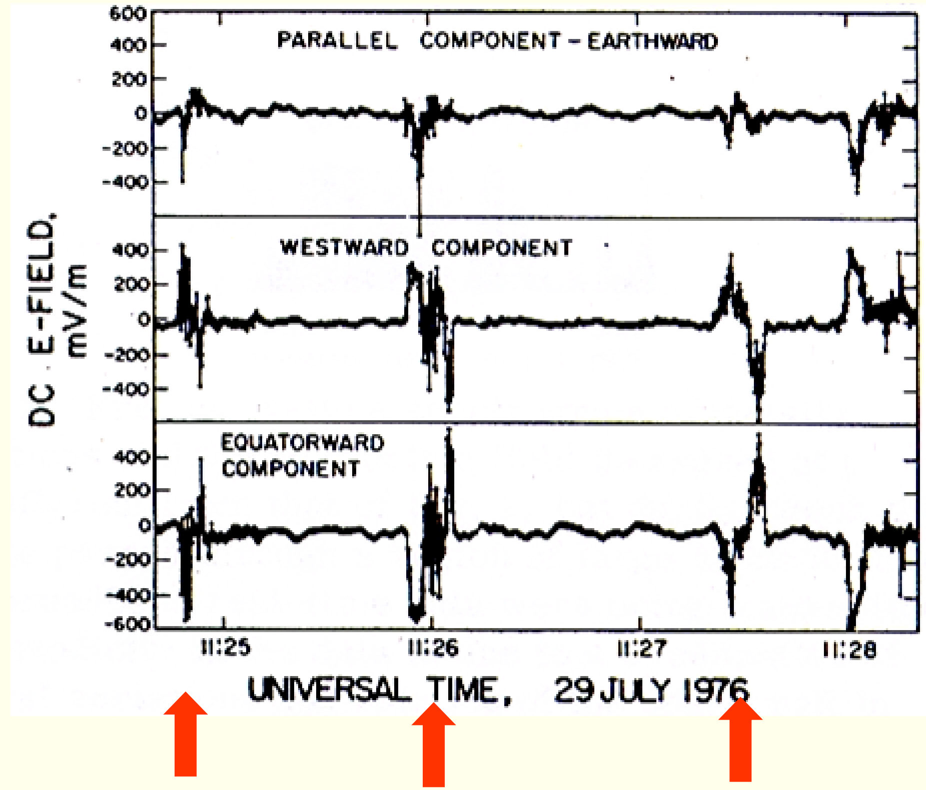
$$\because\ p_1 > p > p_2$$

Spirals – Kelvin-Helmholz instability



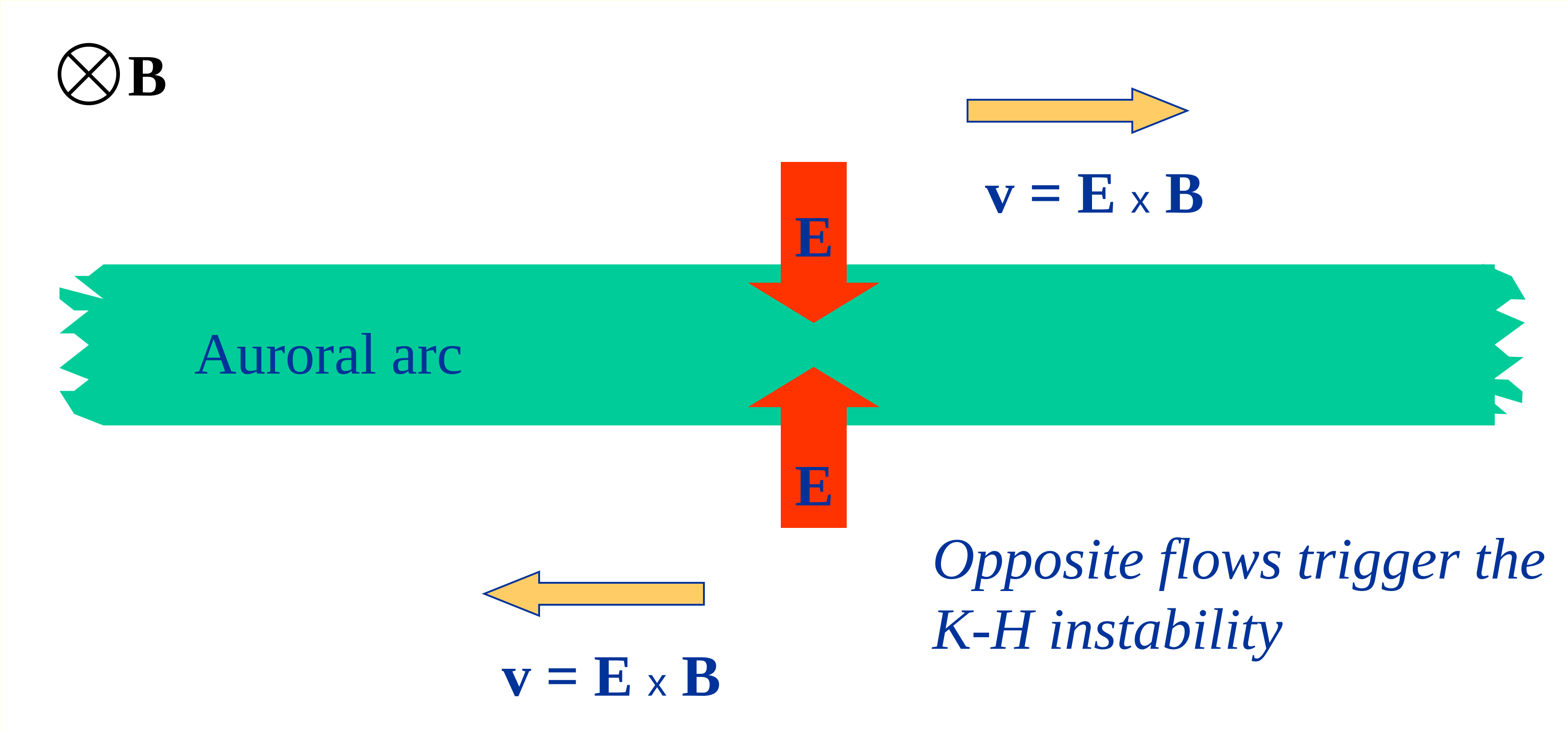
Satellite signatures of U potential



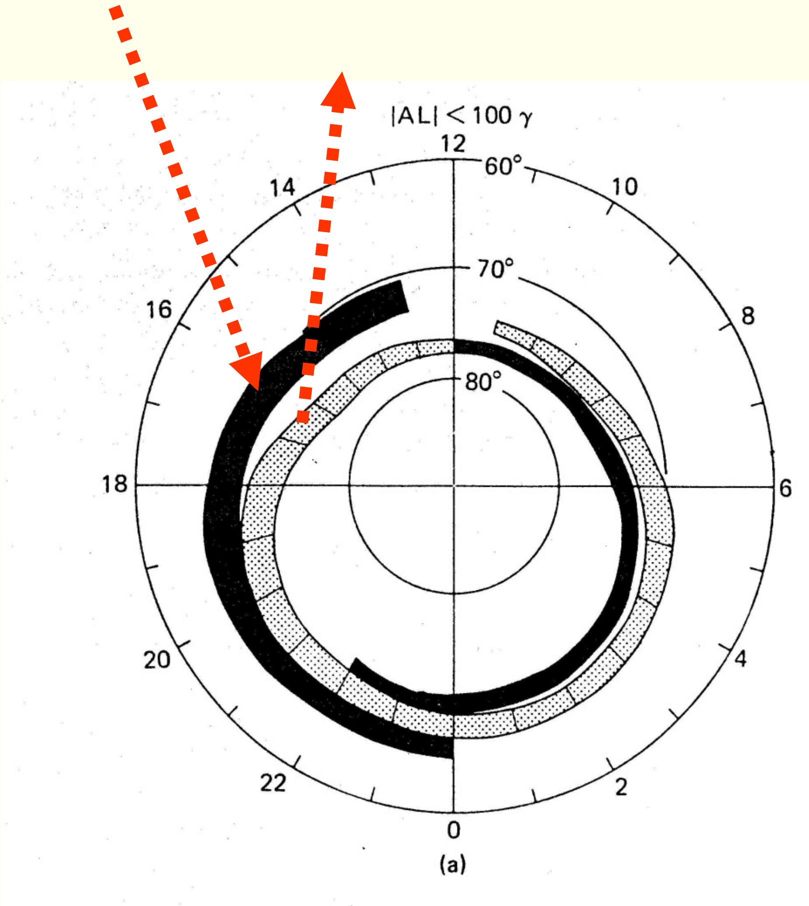


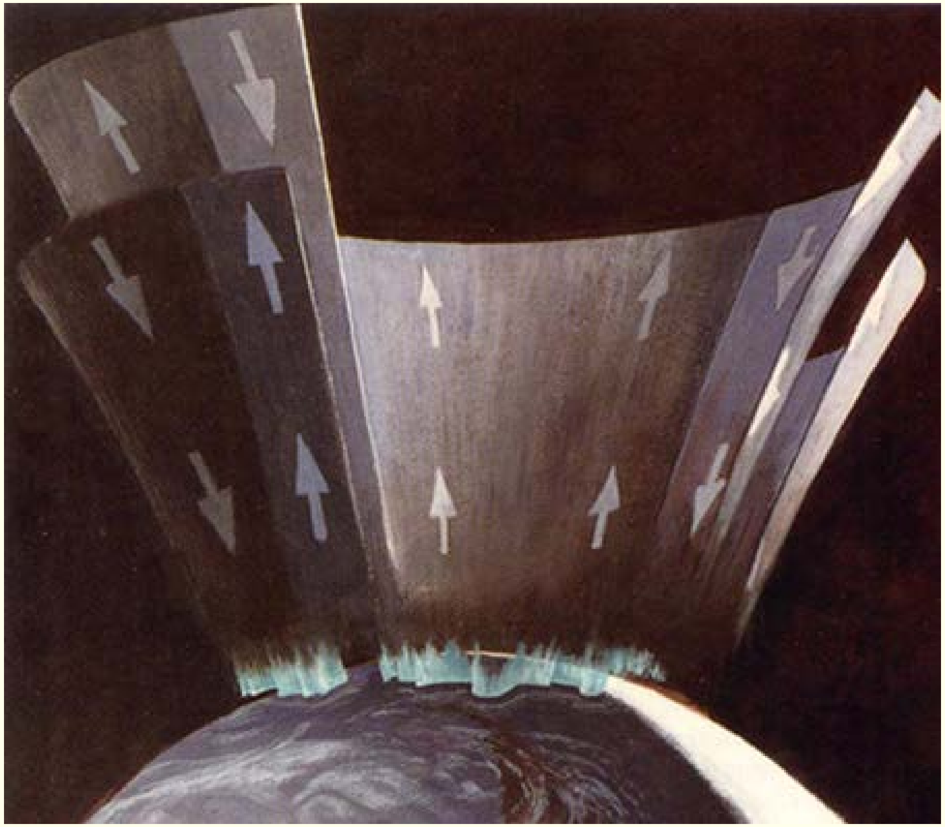
Measurements made by the ISEE satellite (Mozer et al., 1977)

Spirals – Kelvin-Helmholz instability

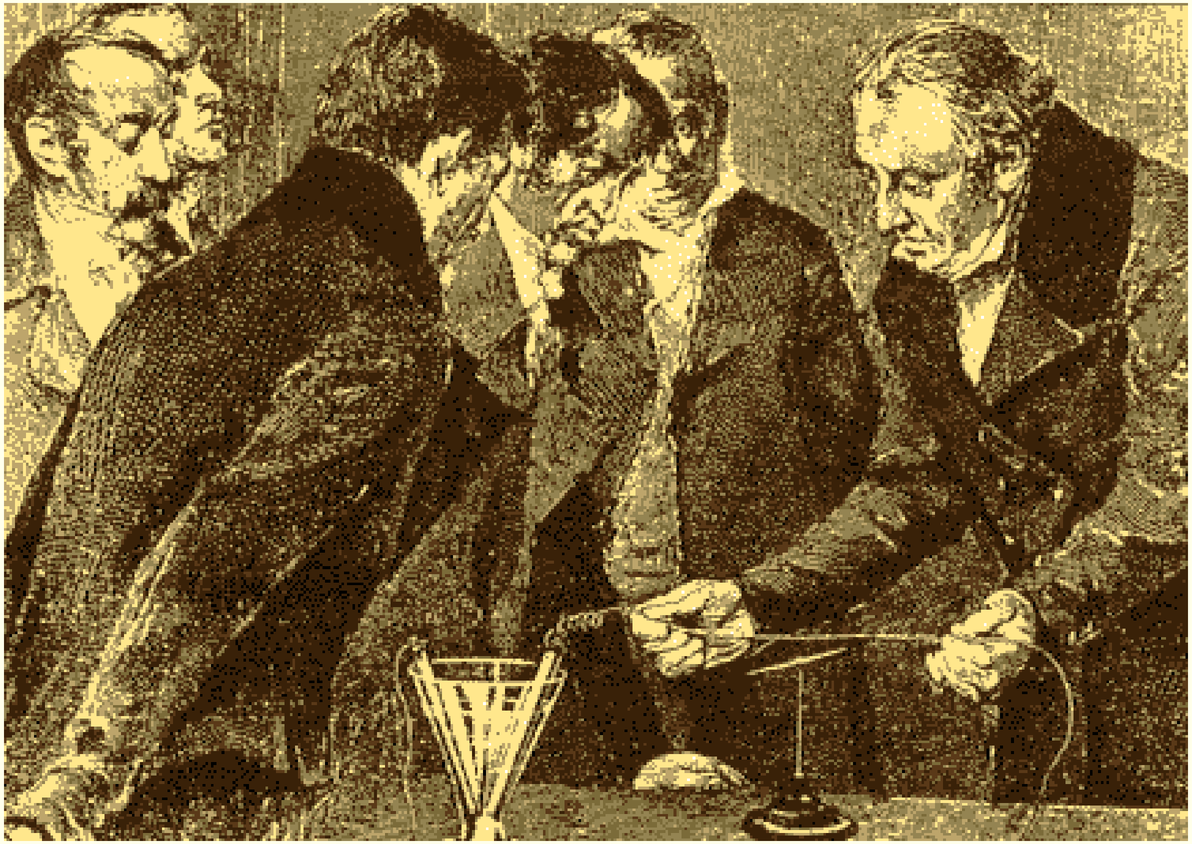


Birkeland currents in the auroral oval

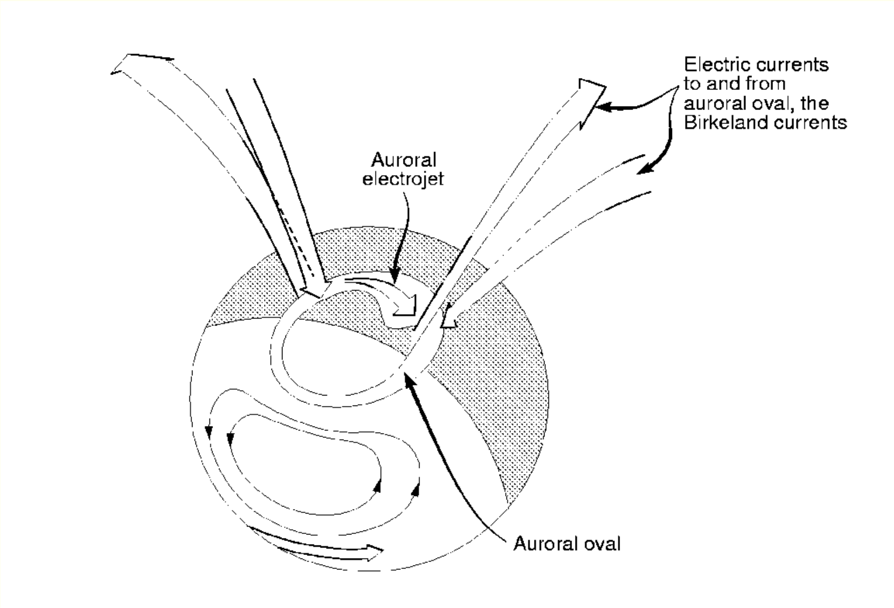




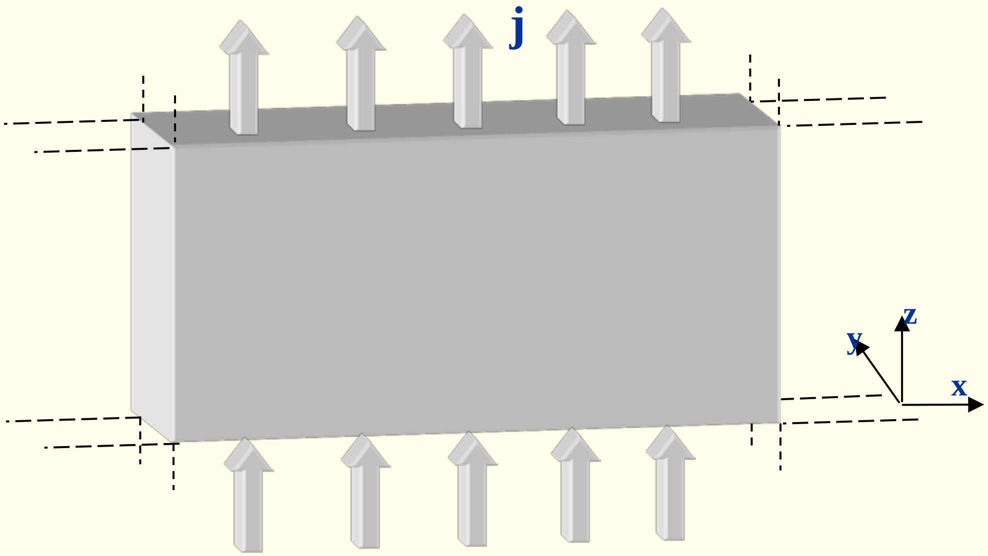
How can you measure currents in space?



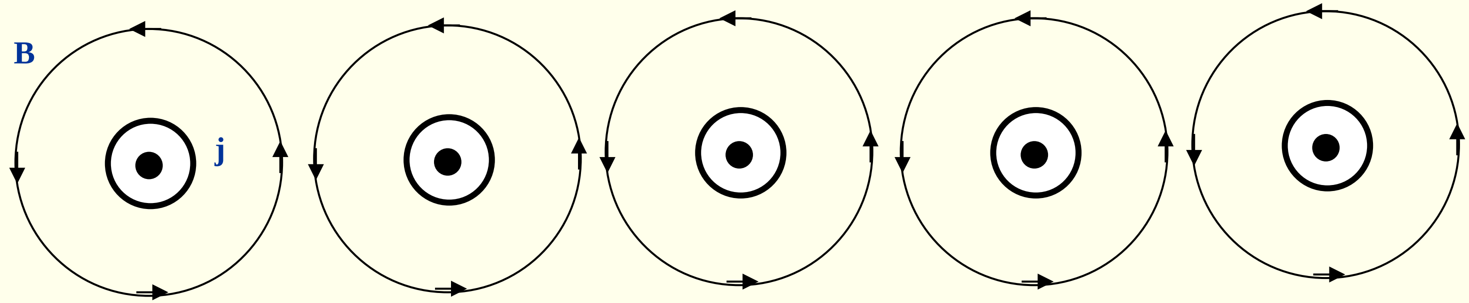
Current sheet approximation



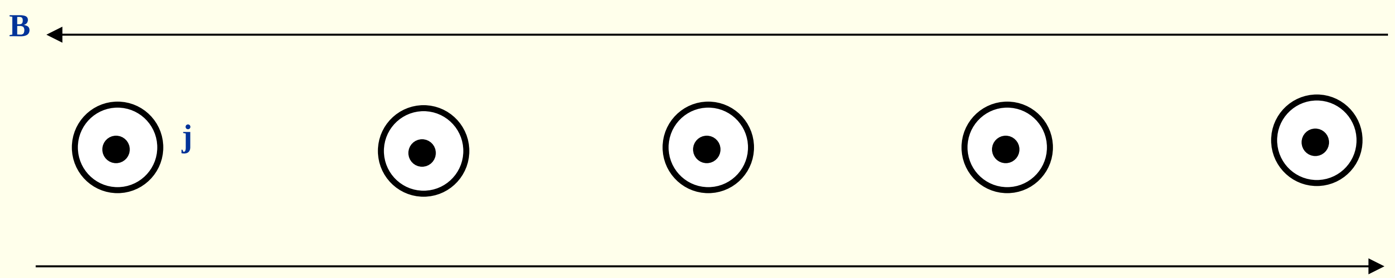
Approximate currents by thin current sheets with infinite size in the x- och z-directions.



Current sheet approximation

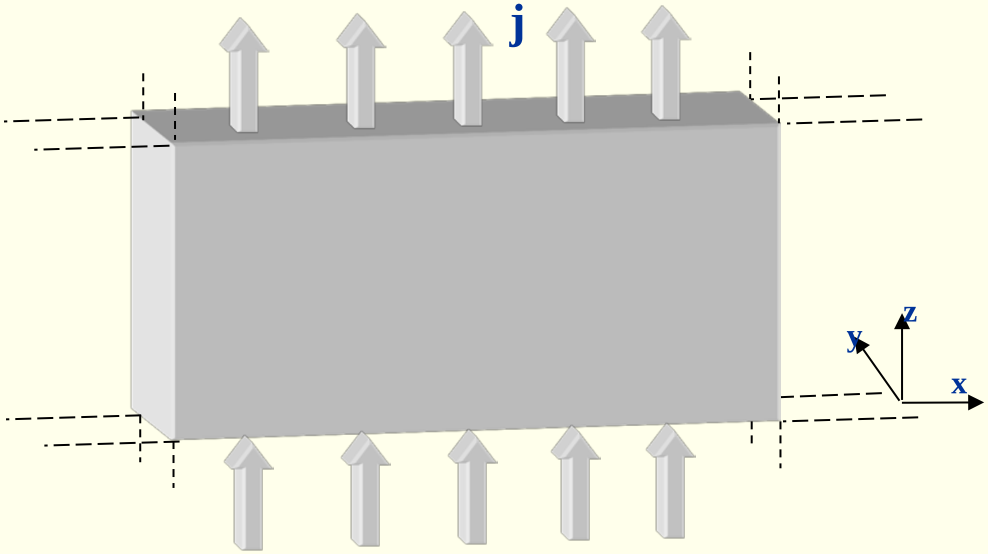


What will the magnetic field around such a current configuration be? Start by approximating with line currents to get a qualitative picture.



The closer you place the line currents, the more the magnetic fields between the line currents will cancel

Current sheet approximation and Ampére's law



$$\left(\frac{\partial B_z}{\partial y}-\frac{\partial B_y}{\partial z},\ \frac{\partial B_x}{\partial z}-\frac{\partial B_z}{\partial x},\ \frac{\partial B_y}{\partial x}-\frac{\partial B_x}{\partial y}\right)=\mu_0\left(j_x,\ j_y,\ j_z\right)$$

But $\frac{\partial}{\partial x}=0$ and $\frac{\partial}{\partial z}=0$

$$\left(\frac{\partial B_z}{\partial y},\ 0,\ -\frac{\partial B_x}{\partial y}\right)=\mu_0\left(0,\ 0,\ j_z\right)$$

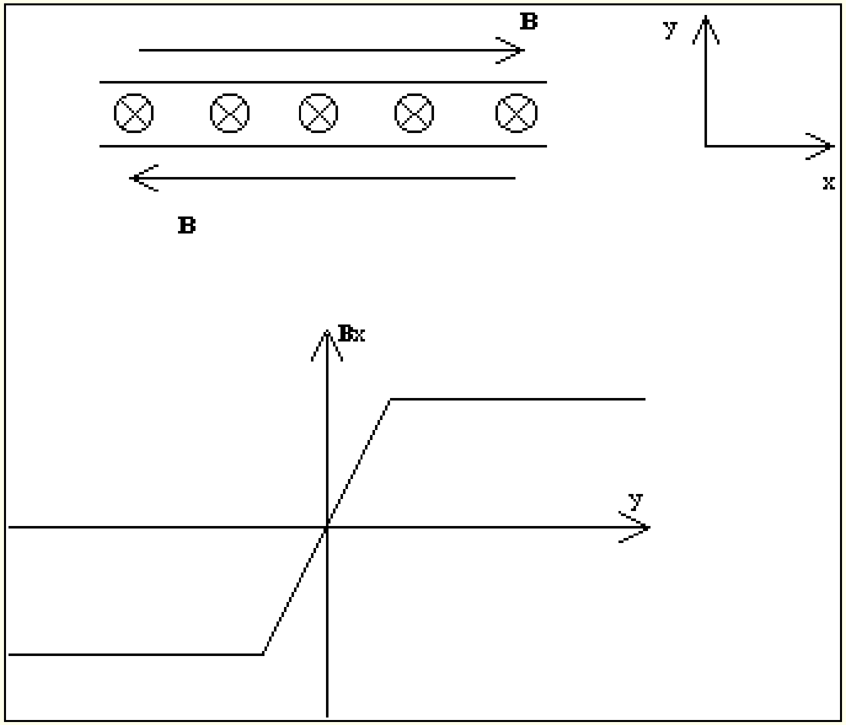
eller

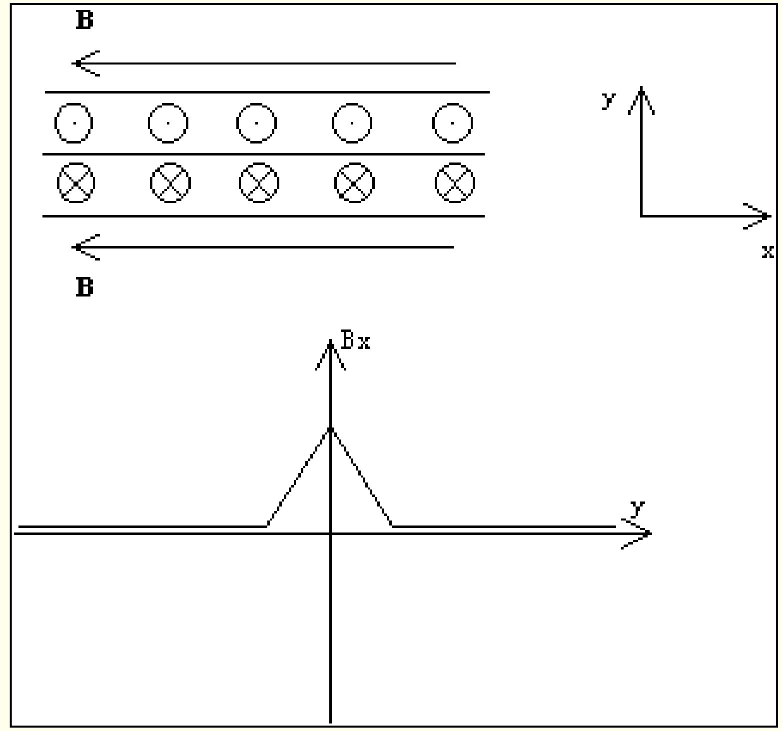
Ampére's law (no time dependence):

$$\nabla\times\mathbf{B}=\mu_0\mathbf{j}$$

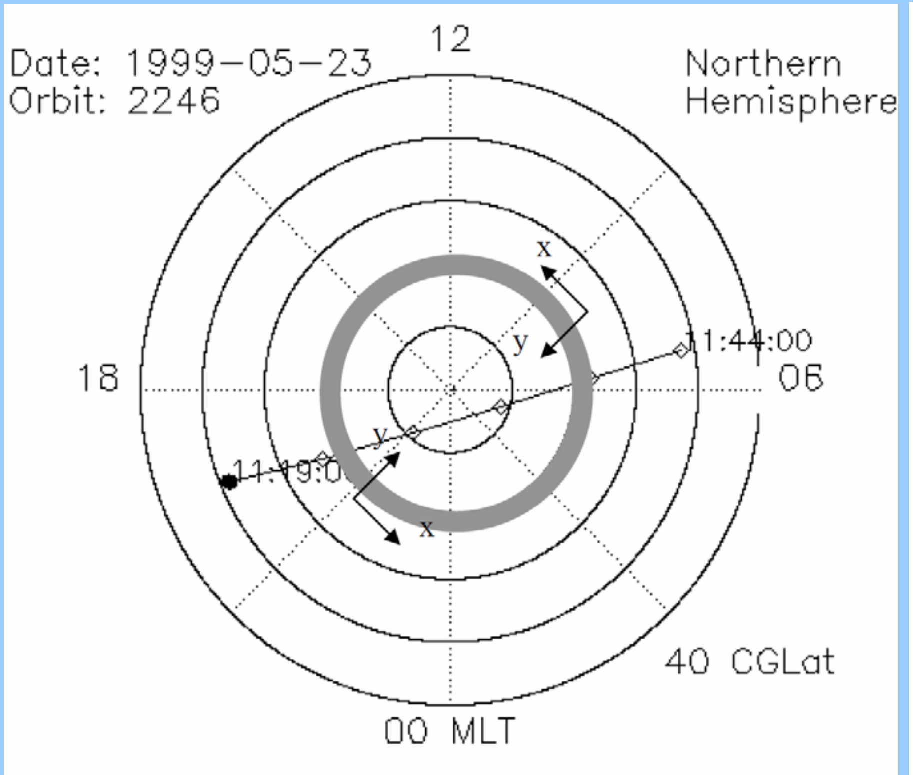
$$j_z=-\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

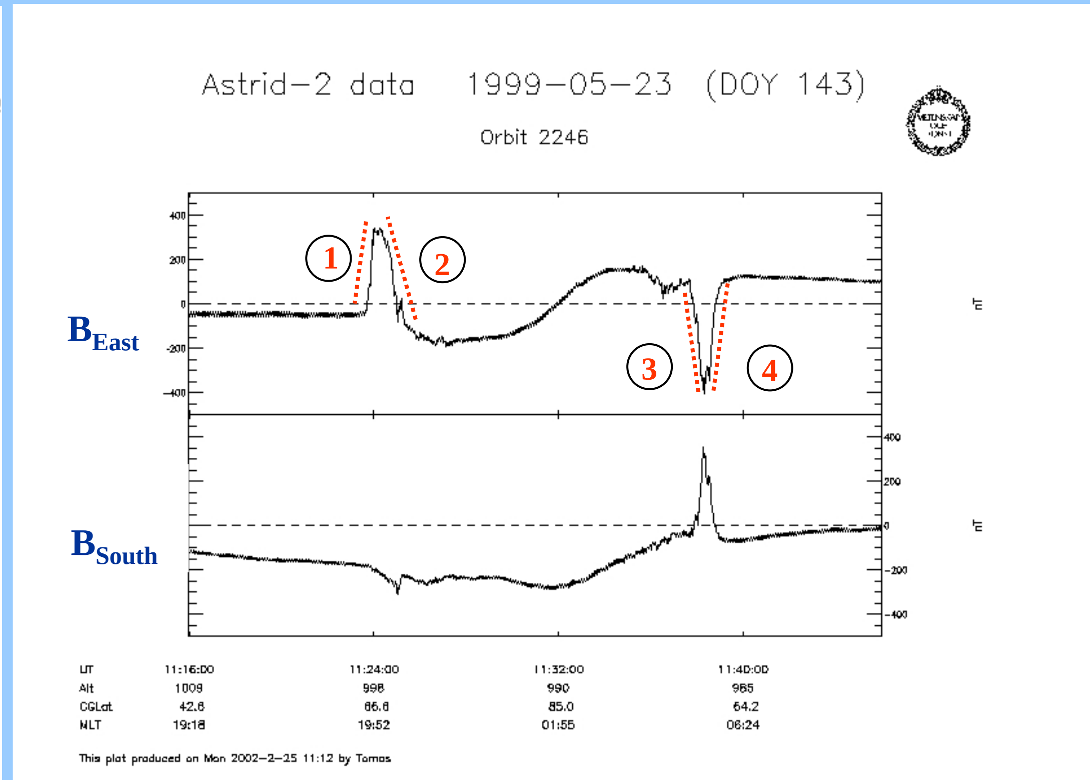
Current sheet - example





$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$





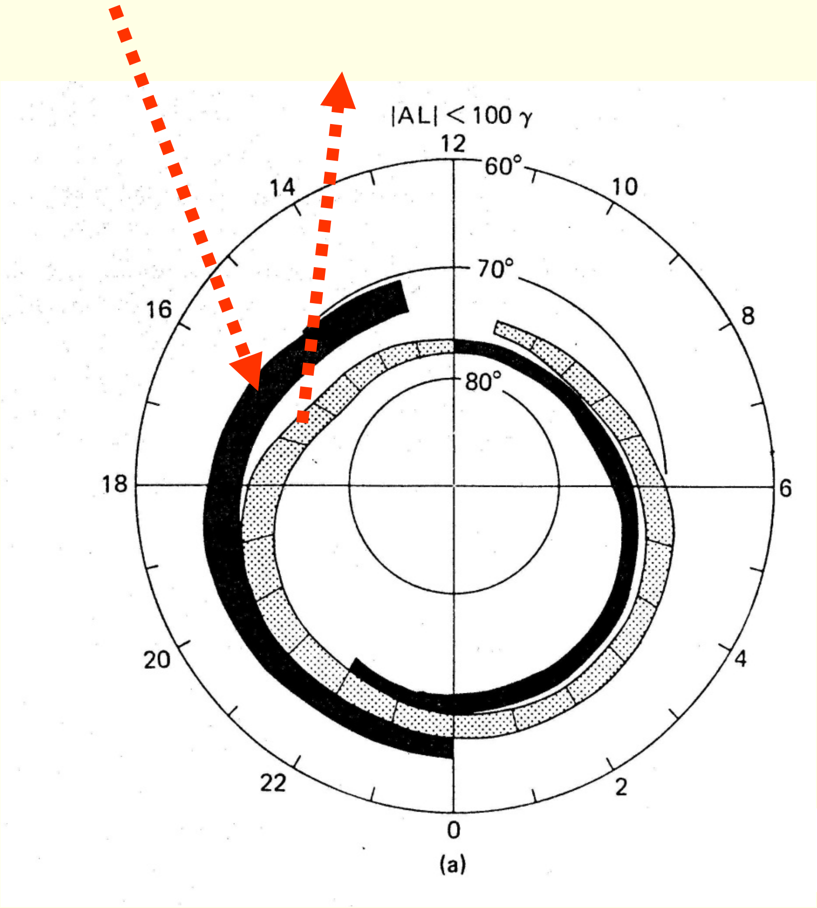
$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

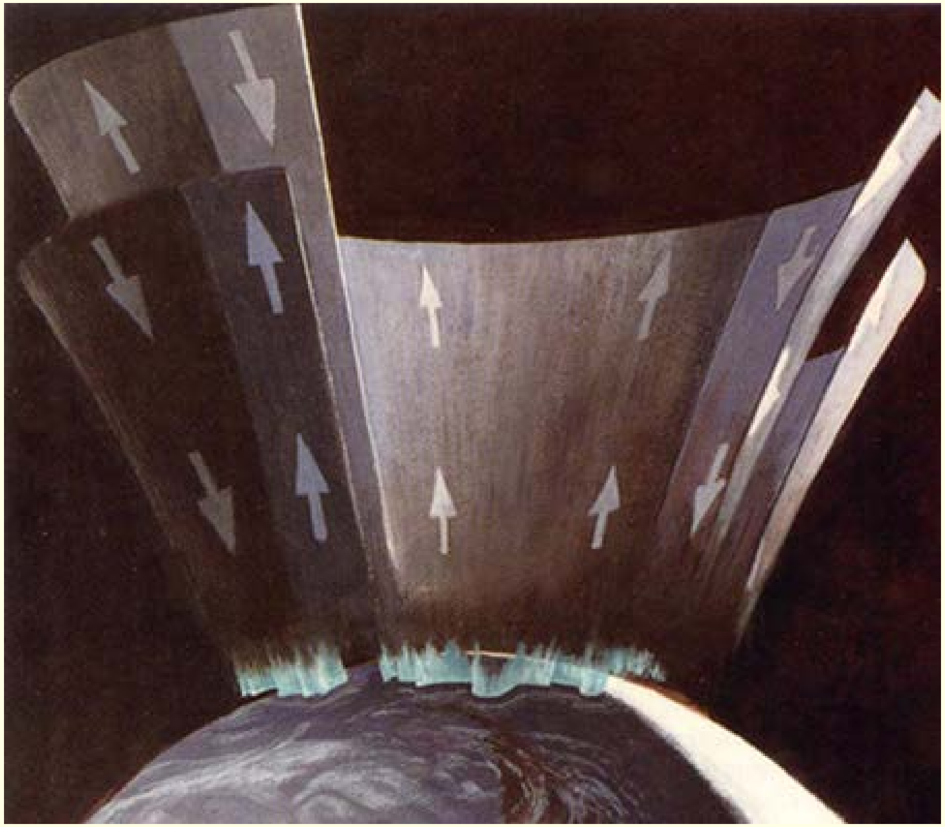
What is the direction of the current in current sheet 1?

Blue — Into the ionosphere

Red — Out of the ionosphere

Birkeland currents in the auroral oval





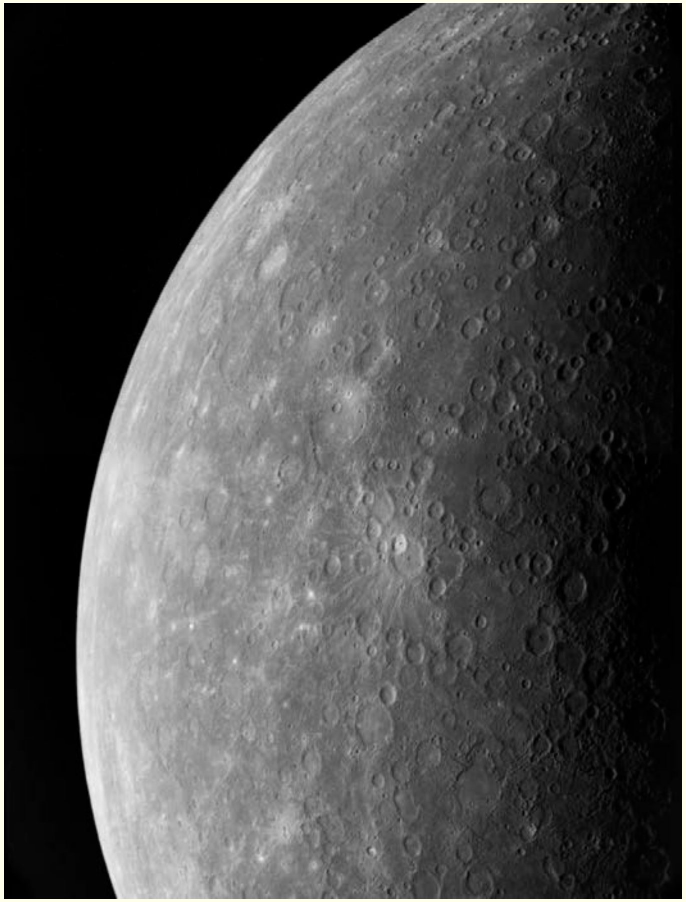
At what planets do you expect aurora to exist?

Blue	Earth, Mercury, Jupiter, Saturn
Yellow	Earth, Venus, Jupiter, Saturn, Uranus, Neptune
Green	Earth, Mars, Jupiter, Saturn, Uranus, Neptune
Red	Earth, Jupiter, Saturn, Uranus, Neptune

What do we need to have an aurora?

- Magnetic field (to guide the plasma particles towards the planet)
- Atmosphere (to create emissions)

Mercury



- No atmosphere
- X-ray aurora???
  Can possibly be created by electrons colliding directly with the planetary surface and lose their energy in one single collision.

Jupiter aurora

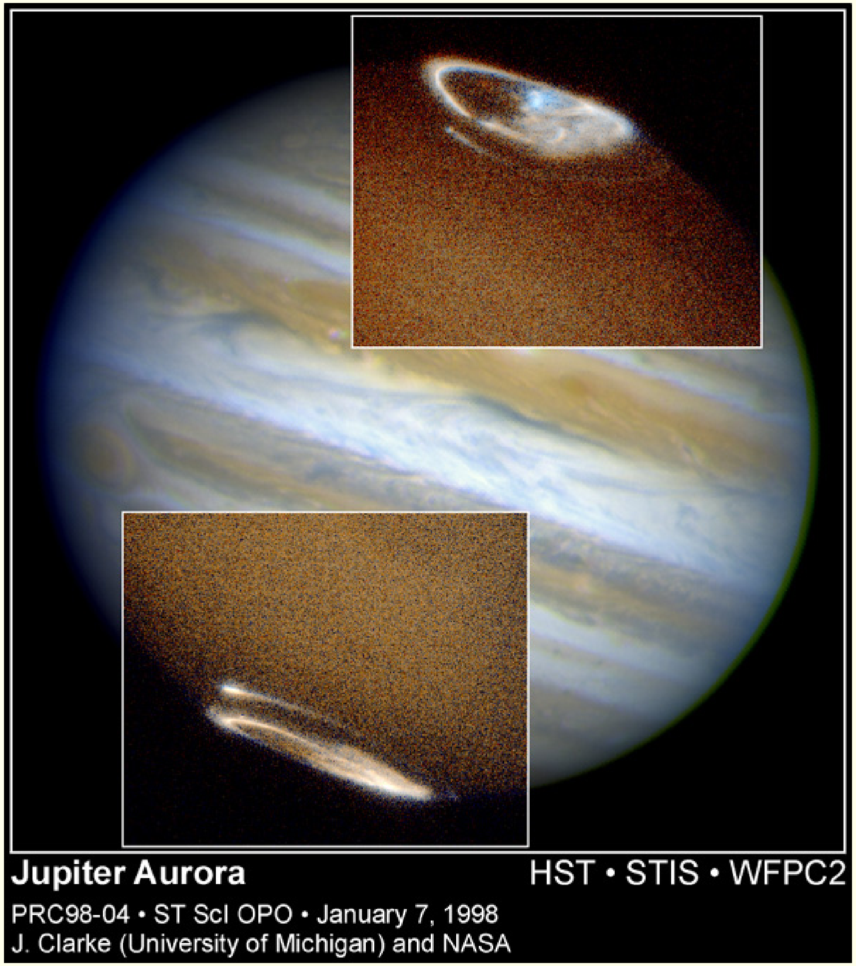
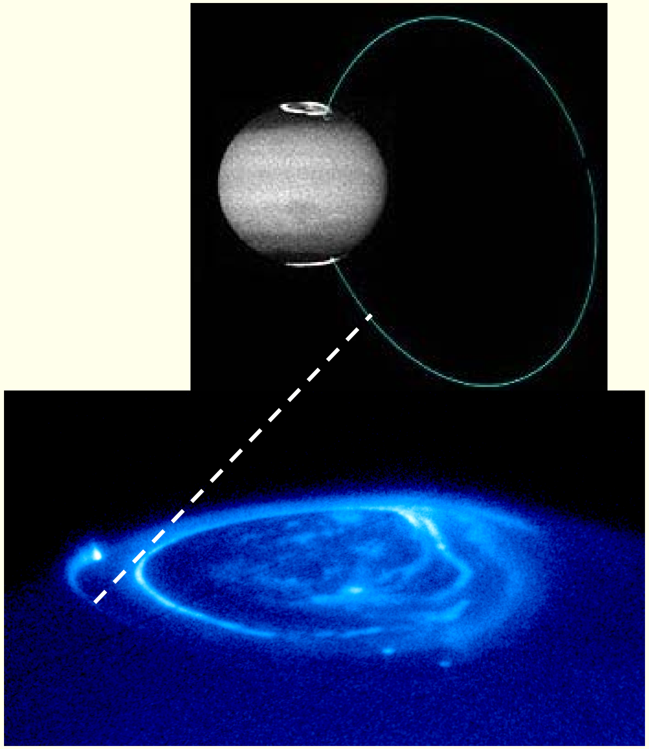


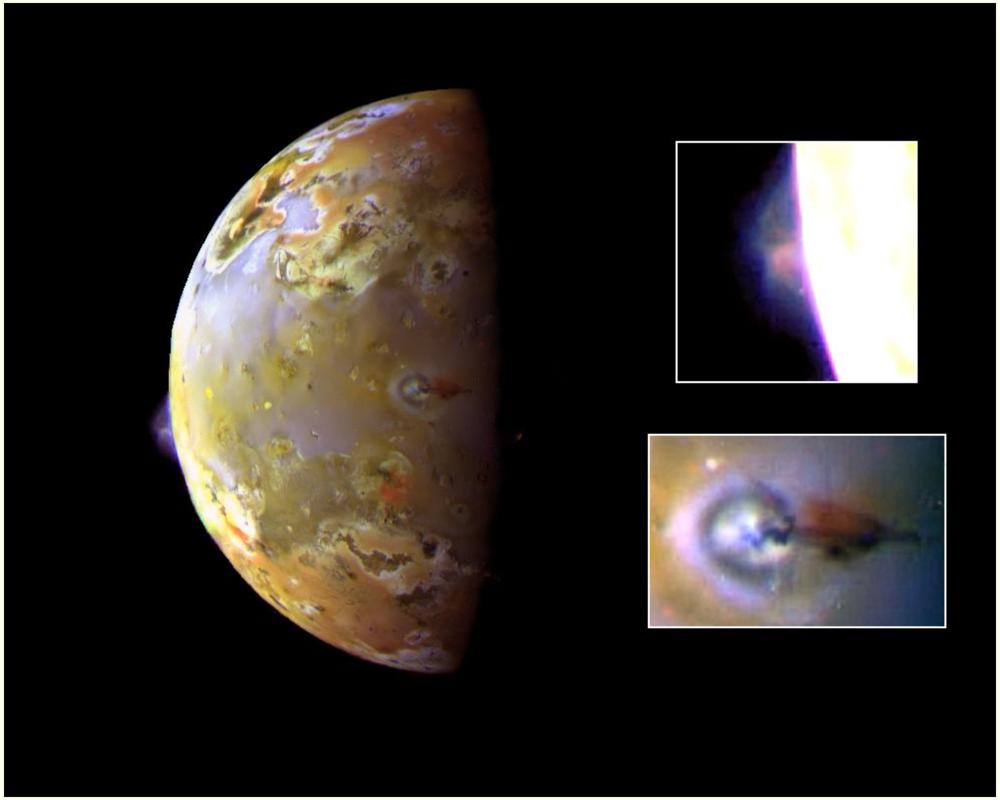
Foto från Hubble Space Telescope

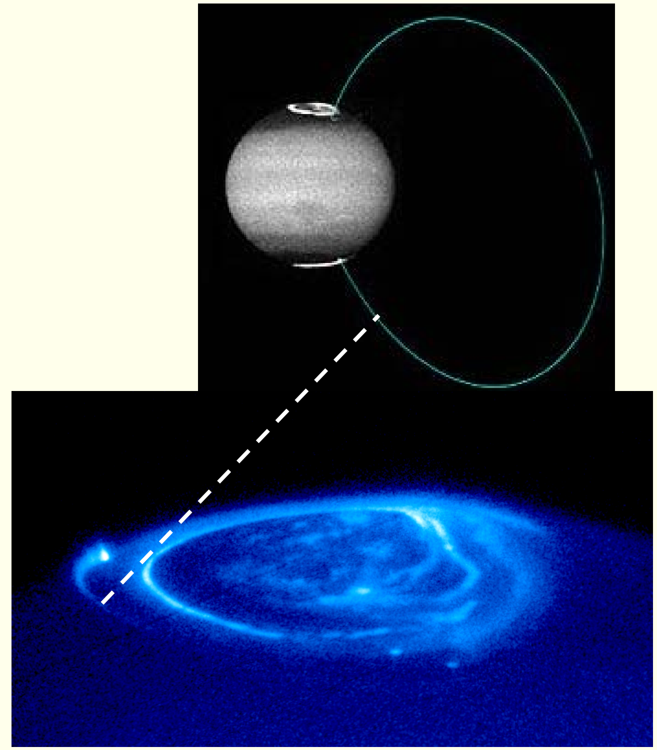
- Jupiter's aurora has a power of ~1000 TW *(compare Earth: ~100 GW, nuclear power plant: ~1 GW)*
- Note the "extra" oval on Io's flux tube!



Jupiter and Io

Photo from rymdsonden Galileo

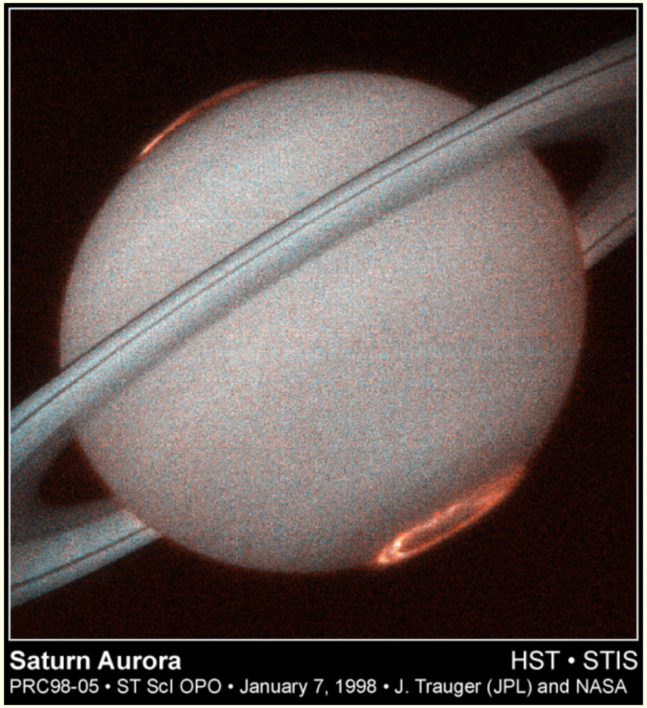




The Jupiter moon Io is very volcanically active, and deposes large amounts of dust and gas in Jupiter's magnetosphere. This is ionized by the sunlight, and the charged plasma partícles follow Jupiter's magnetic field lines towards the atmosphere and cause auroral emissions.

Aurora of the other planets

Saturn



Saturnus' aurora: not noticeably different from Jupiter's, but much weaker. (Total power about the same as Earth's aurora.)

Uranus: *Auora detected in UV. Probably associated with Uranus' ring current/radiotion belts and not very dynamic.*

Neptunus: *weak UV aurora detected.*

Mars, Venus: *No aurora.*

Prerequisites for…

Life

- Energy source (sun)
- Atmosphere
- Magnetic field
- Water

Aurora

- Energy source (sun)
- Atmosphere
- Magnetic field

On space weather and viewing aurora

Some space weather sites

http://spaceweather.com/

http://www.esa-spaceweather.net/

http://sunearthday.nasa.gov/swac/

http://www.noaawatch.gov/themes/space.php

http://www.windows2universe.org/spaceweather/more_details.html

Kiruna

Kiruna all-sky camera:
http://www.irf.se/allsky/rtasc.php

http://sunearthday.nasa.gov/swac/tutorials/aur_kiruna.php

Forecasts:
http://flare.lund.irf.se/rwc/aurora/

http://www.irf.se/Observatory/?link[All-skycamera]=Aurora_sp_statistics

Last Minute!

Last Minute!

- What was the most important thing of today's lecture? Why?
- What was the most unclear or difficult thing of today's lecture, and why?
- Other comments